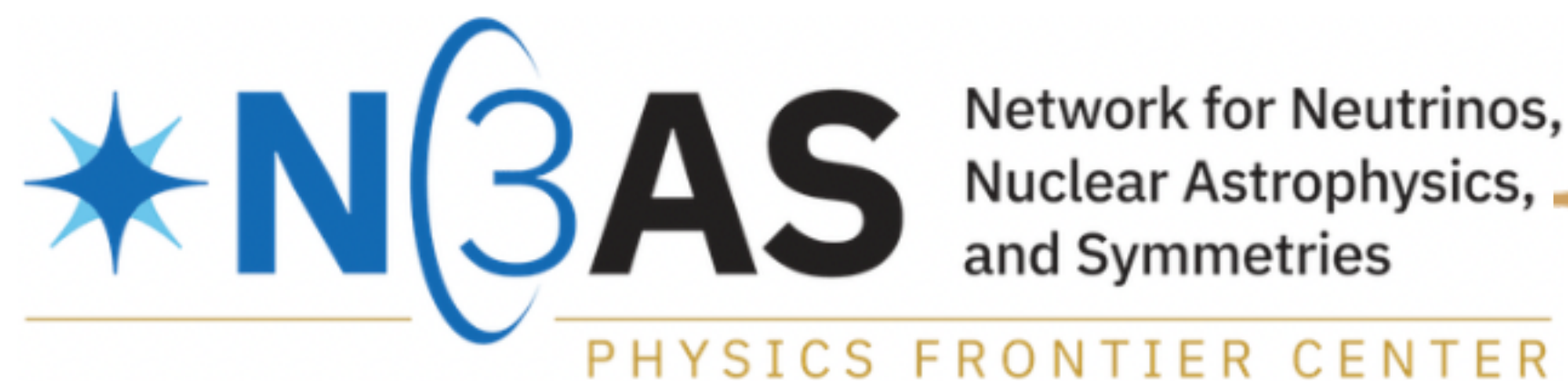


Framework for phase transitions between the Maxwell and Gibbs constructions

Tianqi Zhao 趙天奇

Collaborators: Constantinos Constantinou, Mirco Guerrini, Madappa Prakash, Sophia Han, Christian Drischler, Sanjay Reddy, James Lattimer

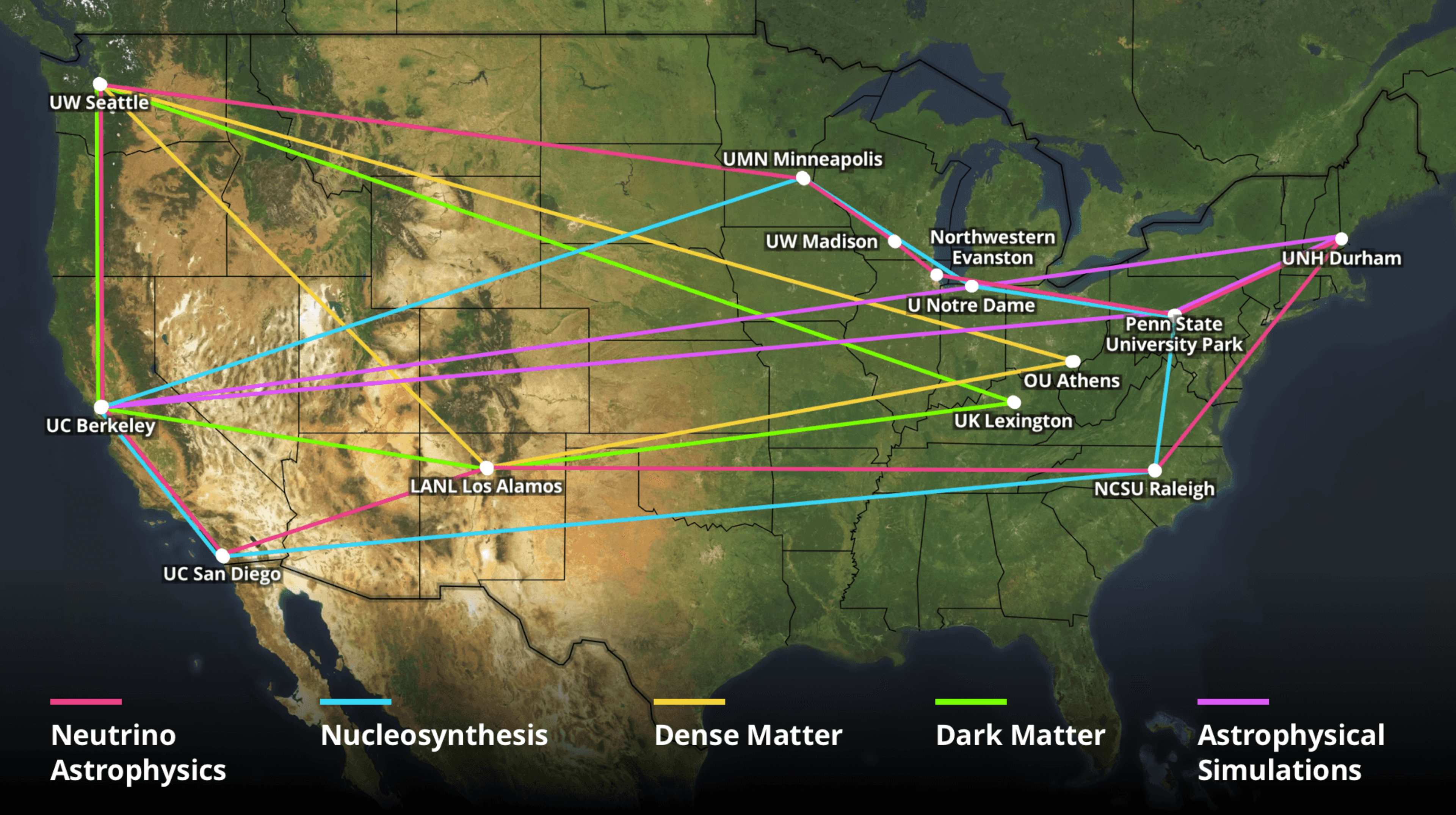
CSQCD at YITP, Oct 7, 2024



OHIO
UNIVERSITY



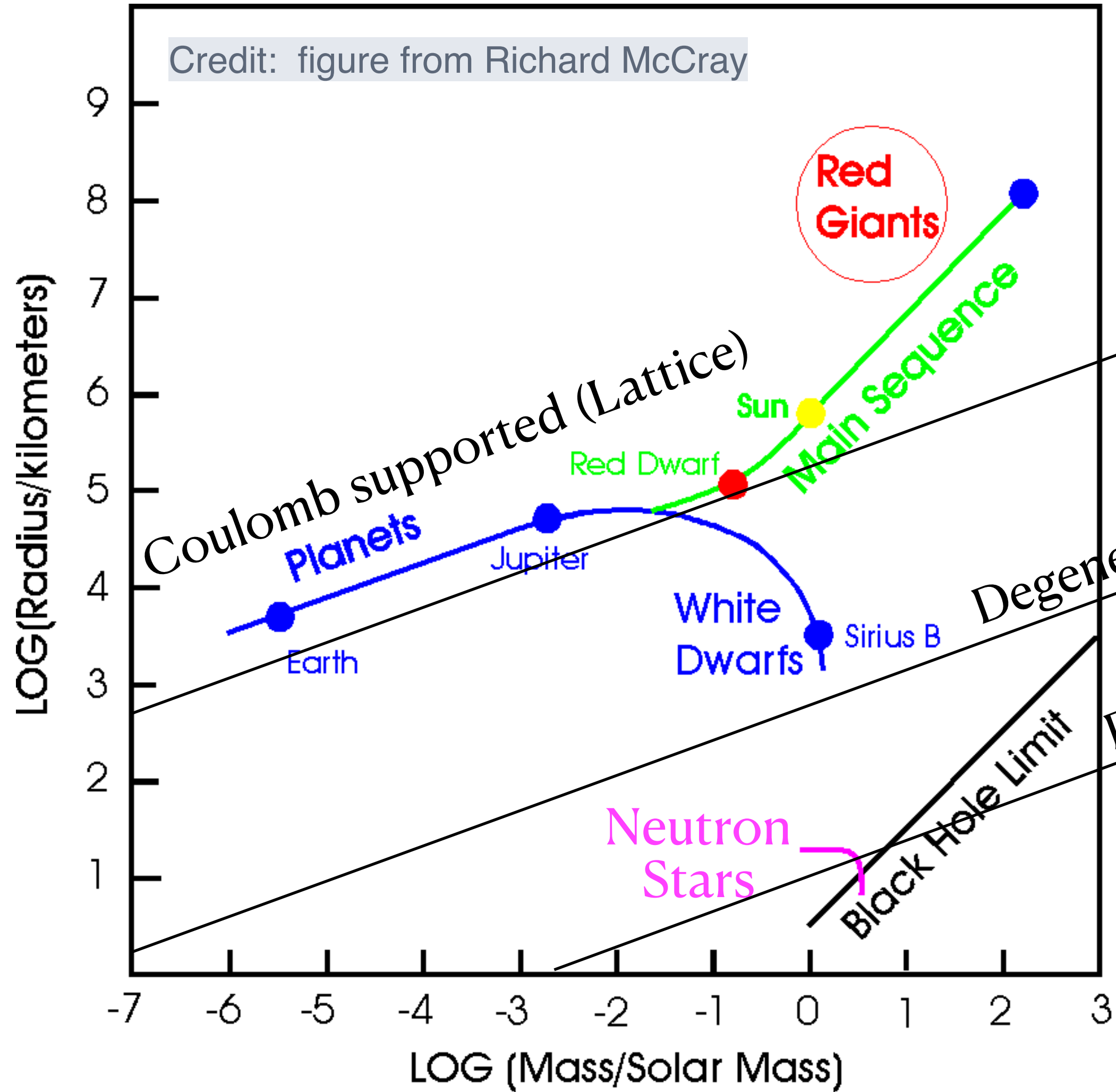
N3AS



Neutron Star

The densest static environment

The environment around Proto-NS, stellar accretion and X-ray bursts



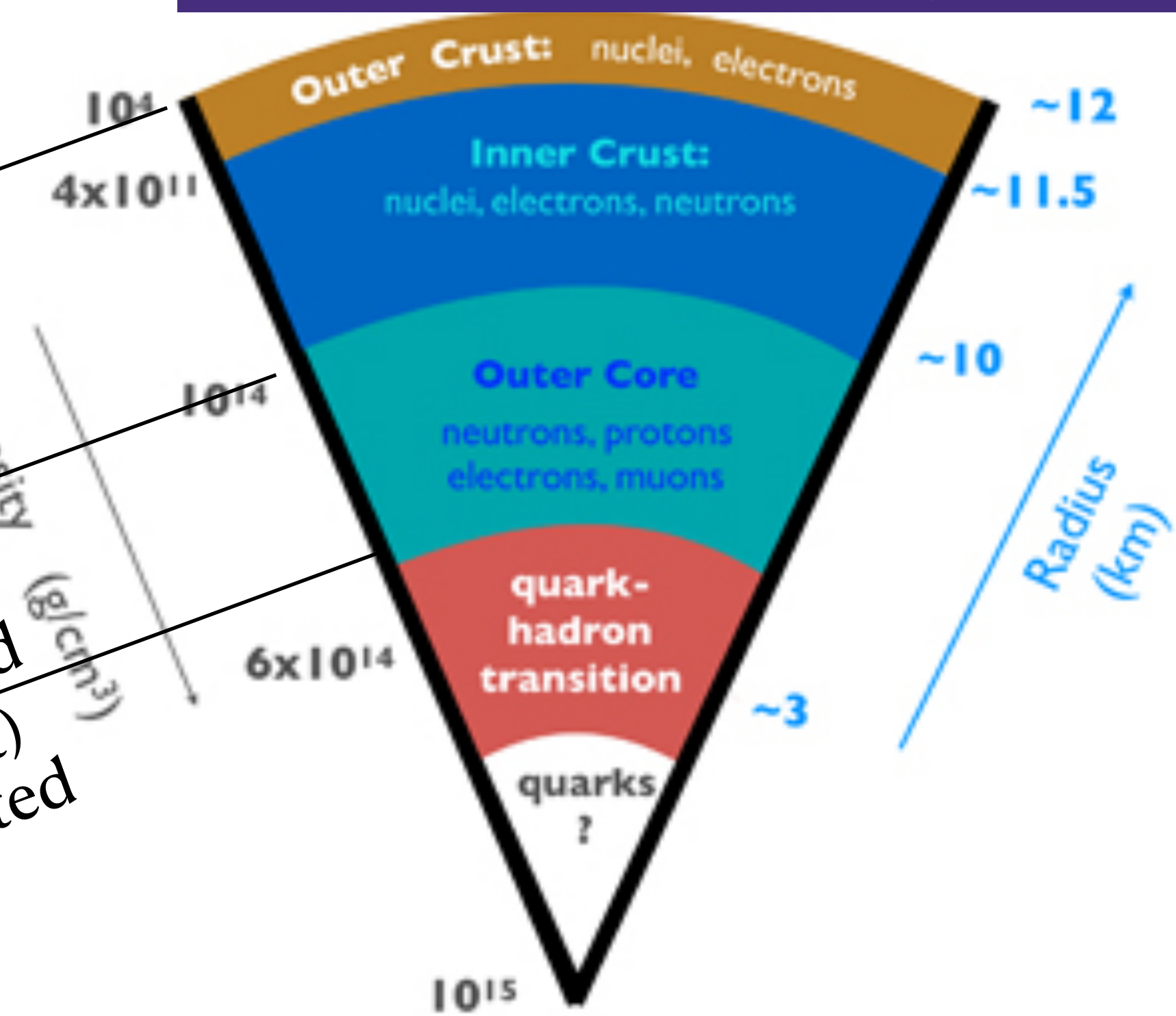
Temperature supported

Degenerate electron supported

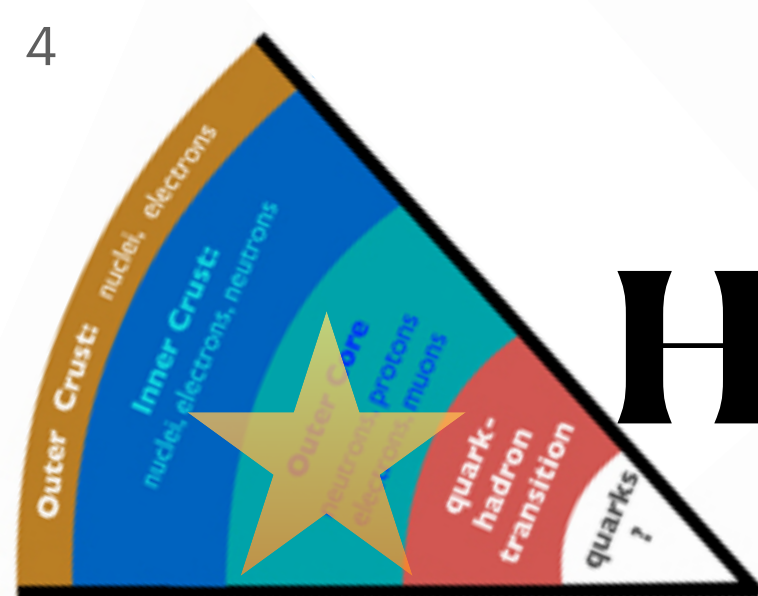
(repulsive nuclear force)
Degenerate neutron supported

(Non perturbative effect)
Degenerate quark supported

Black Hole Limit



Credit: figure from 3G Science White Paper



Hadronic EOS

- Uniform nuclear matter

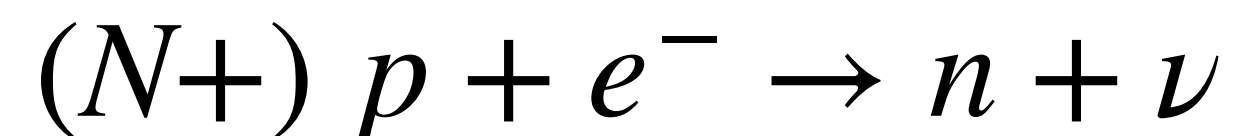
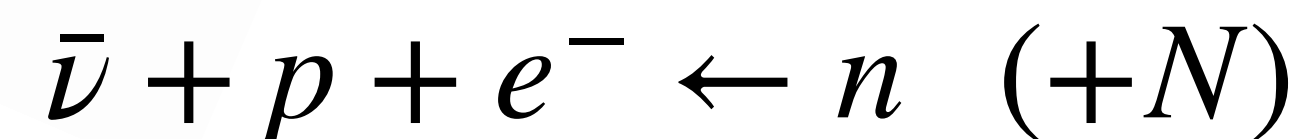
$$\varepsilon(n_p, n_n) \propto n_p^{5/3} + n_n^{5/3} + \dots$$

$$p(n_p, n_n) = n_p \mu_p + n_n \mu_n - \varepsilon$$

$$n, p, e^-$$

$$n_{n,p}, P, \mu$$

- Nucleon weak β -equilibrium



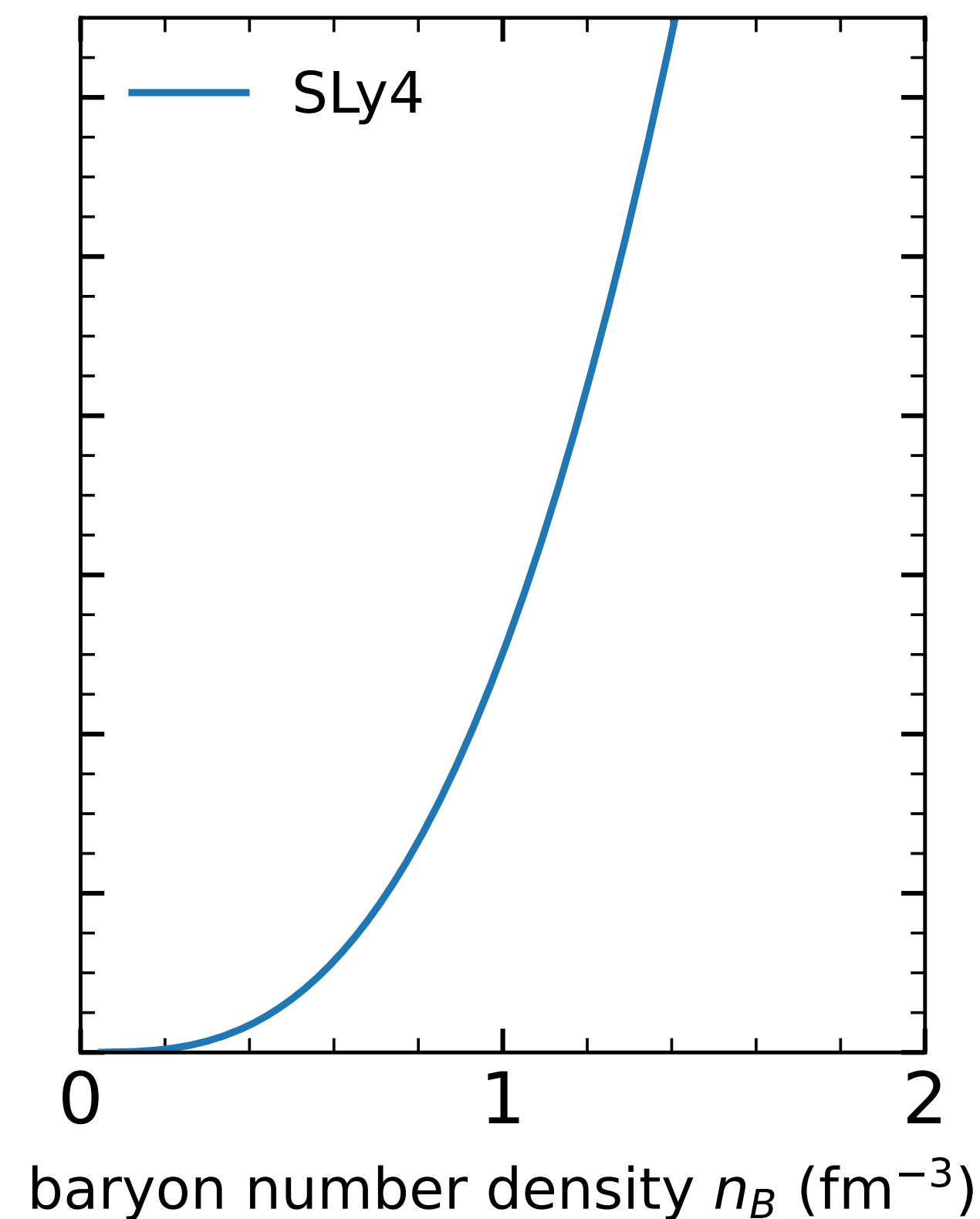
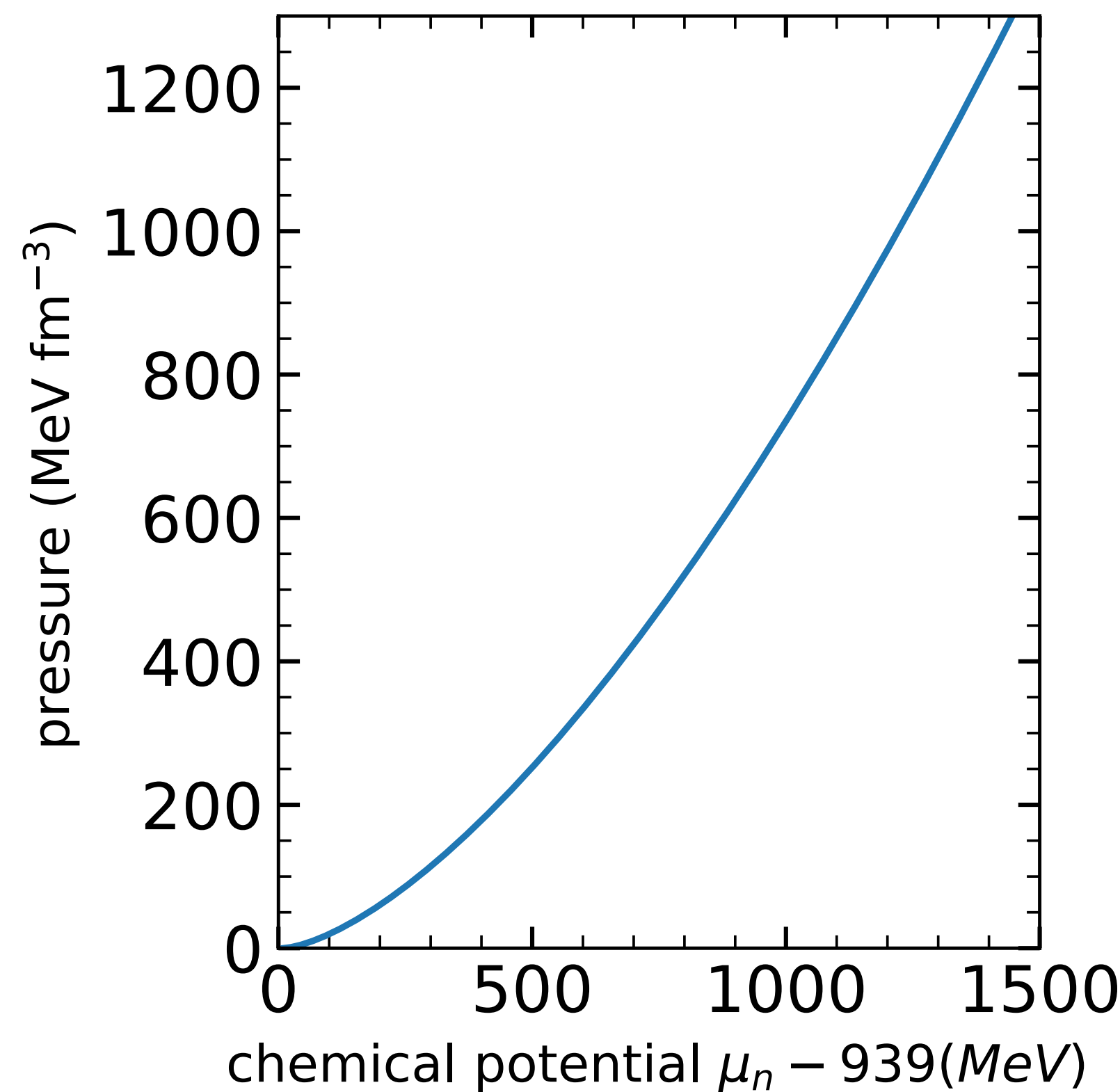
leads to $\mu_p + \mu_e = \mu_n$

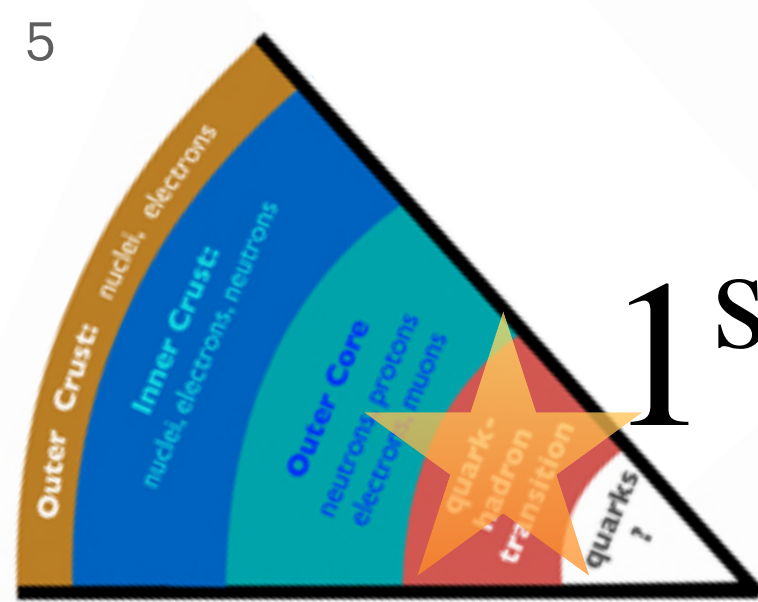
- Baryon number conservation:

$$n_p + n_n = n_B$$

- Local charge neutrality:

$$n_p + n_e = 0$$





1st-order Transition

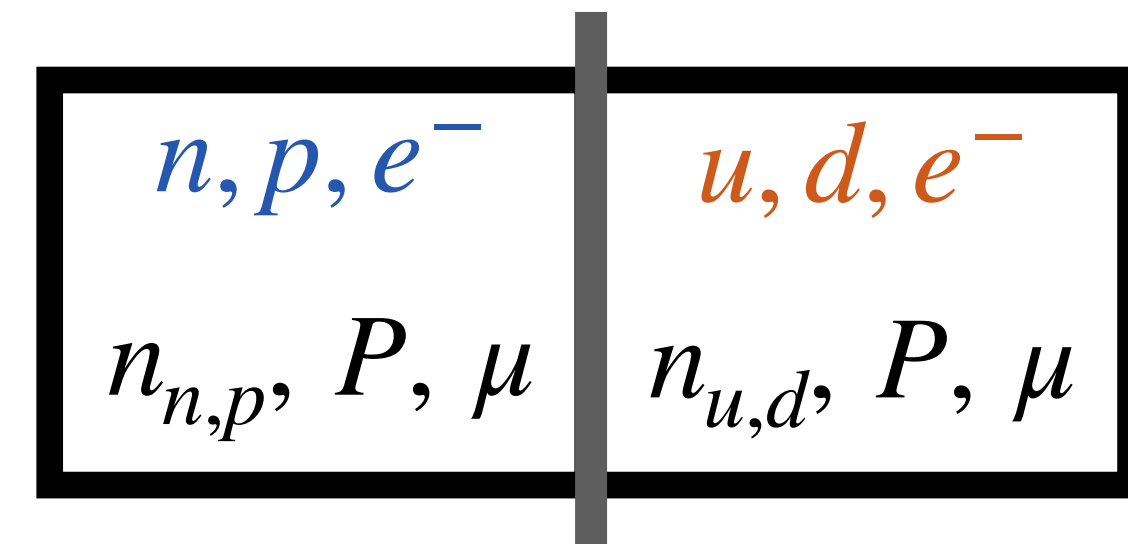
Maxwell construction

- Mechanical equilibrium

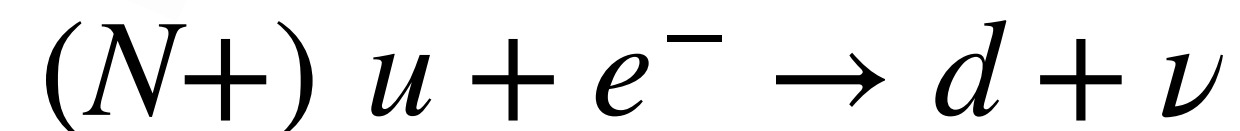
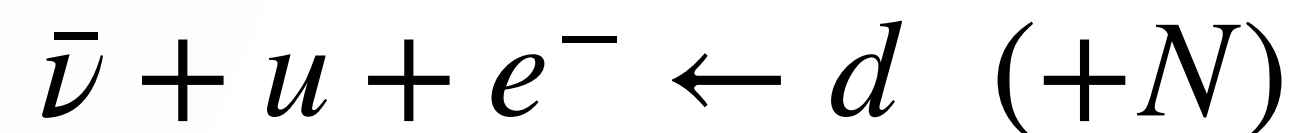
$$P_{npe} = P_{ude} = P$$

- Strong equilibrium

$$\mu_n = \mu_u + 2\mu_d = \mu$$



- (Modified) Urca process



leads to $\mu_u + \mu_e = \mu_d = \mu_s$

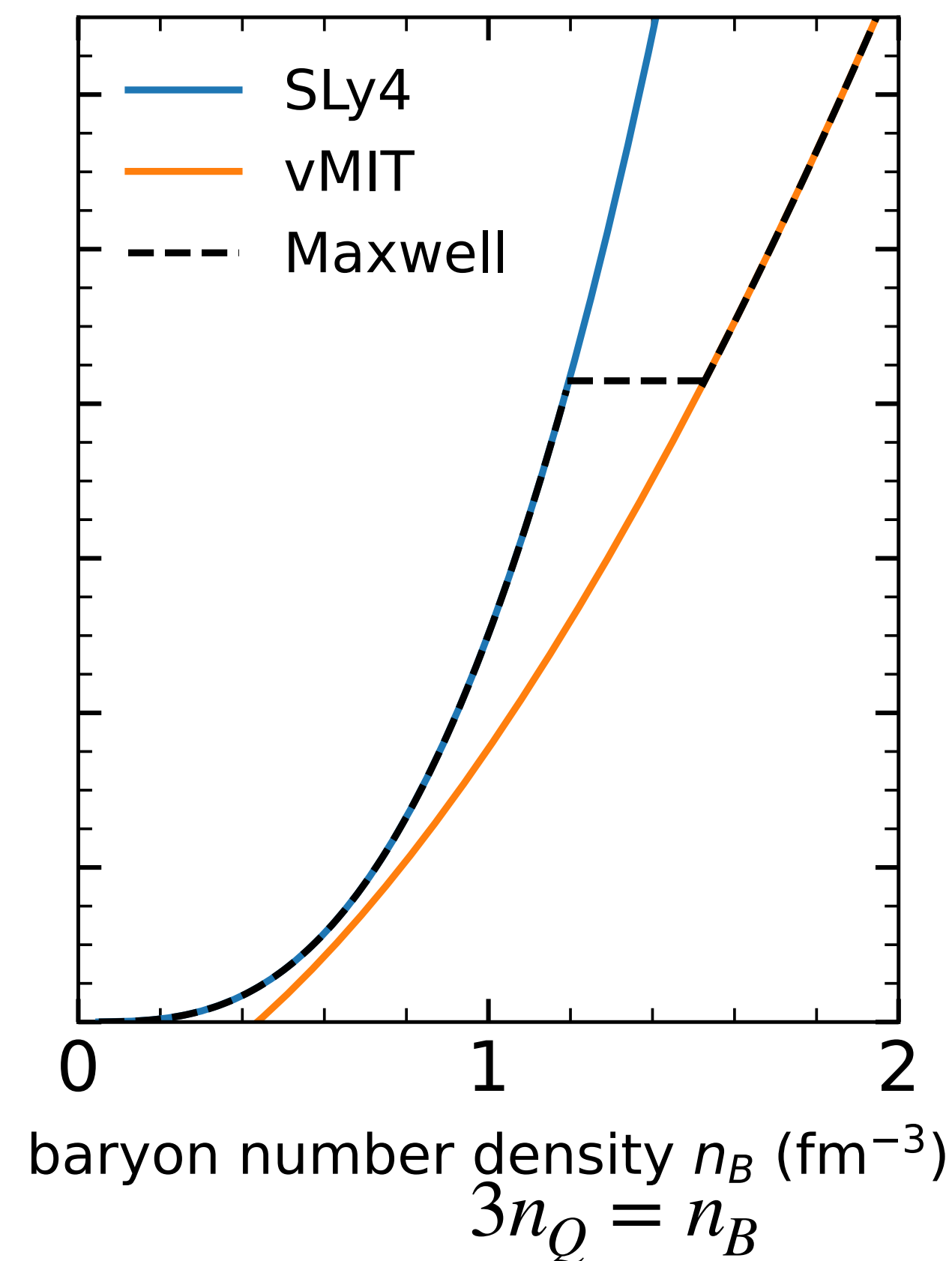
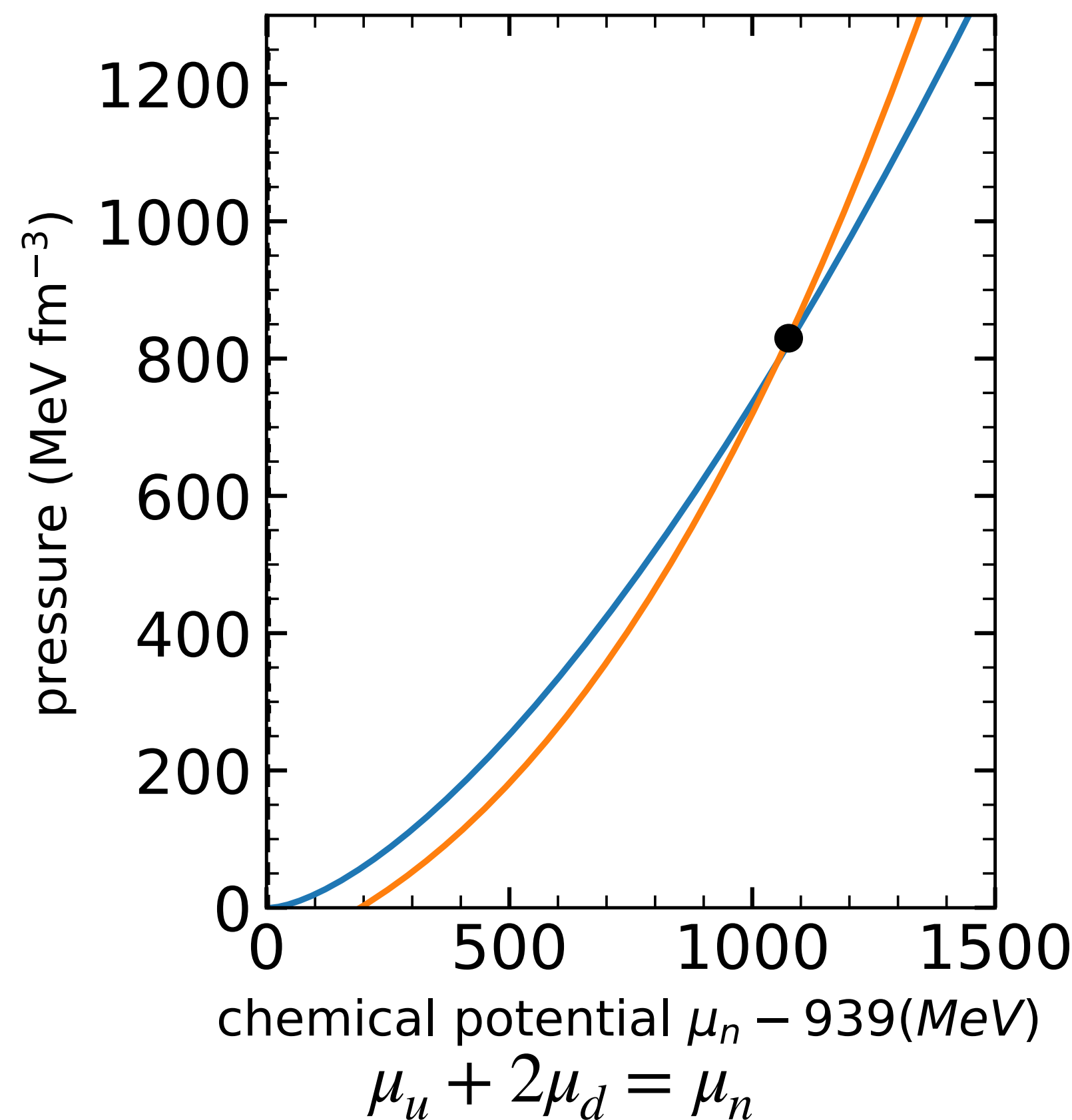
- Baryon number conservation:

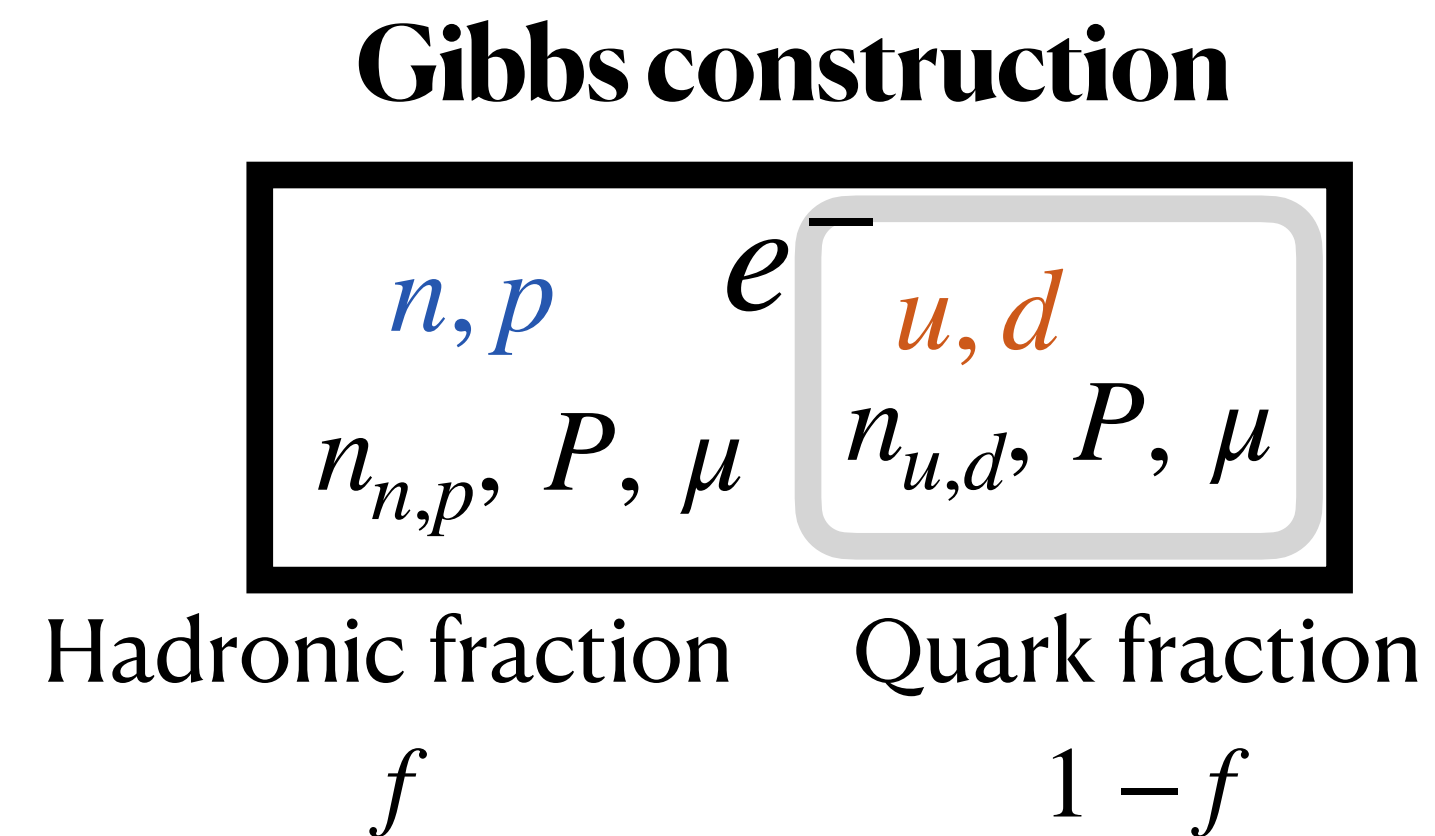
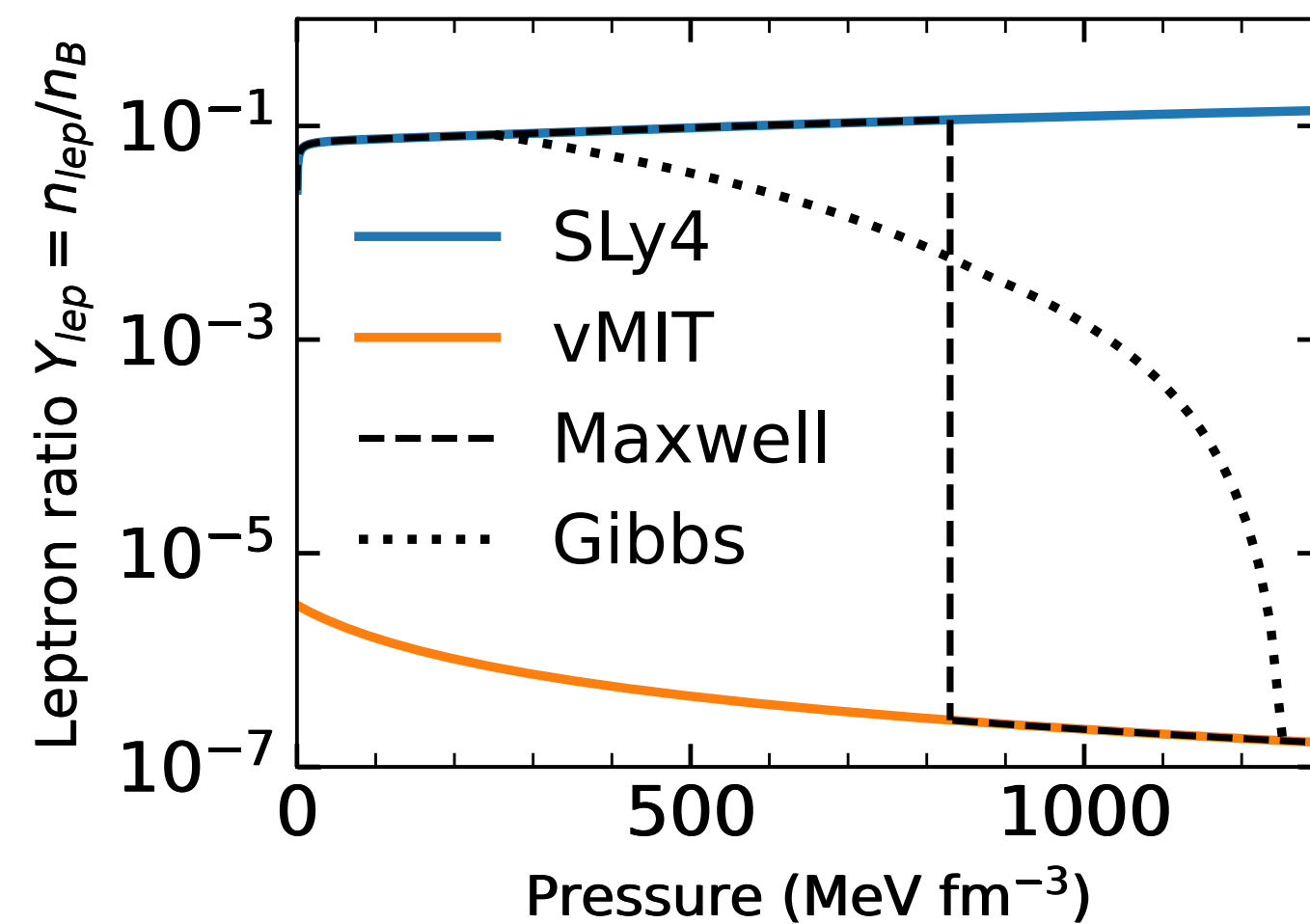
$$n_u + n_d + n_s = n_Q = n_B/3$$

- Local charge neutrality:

$$n_{e,Q} = \frac{2}{3}n_u - \frac{1}{3}n_d - \frac{1}{3}n_s$$

$$n_{e,N} = n_p$$





- Local charge neutrality (Maxwell):

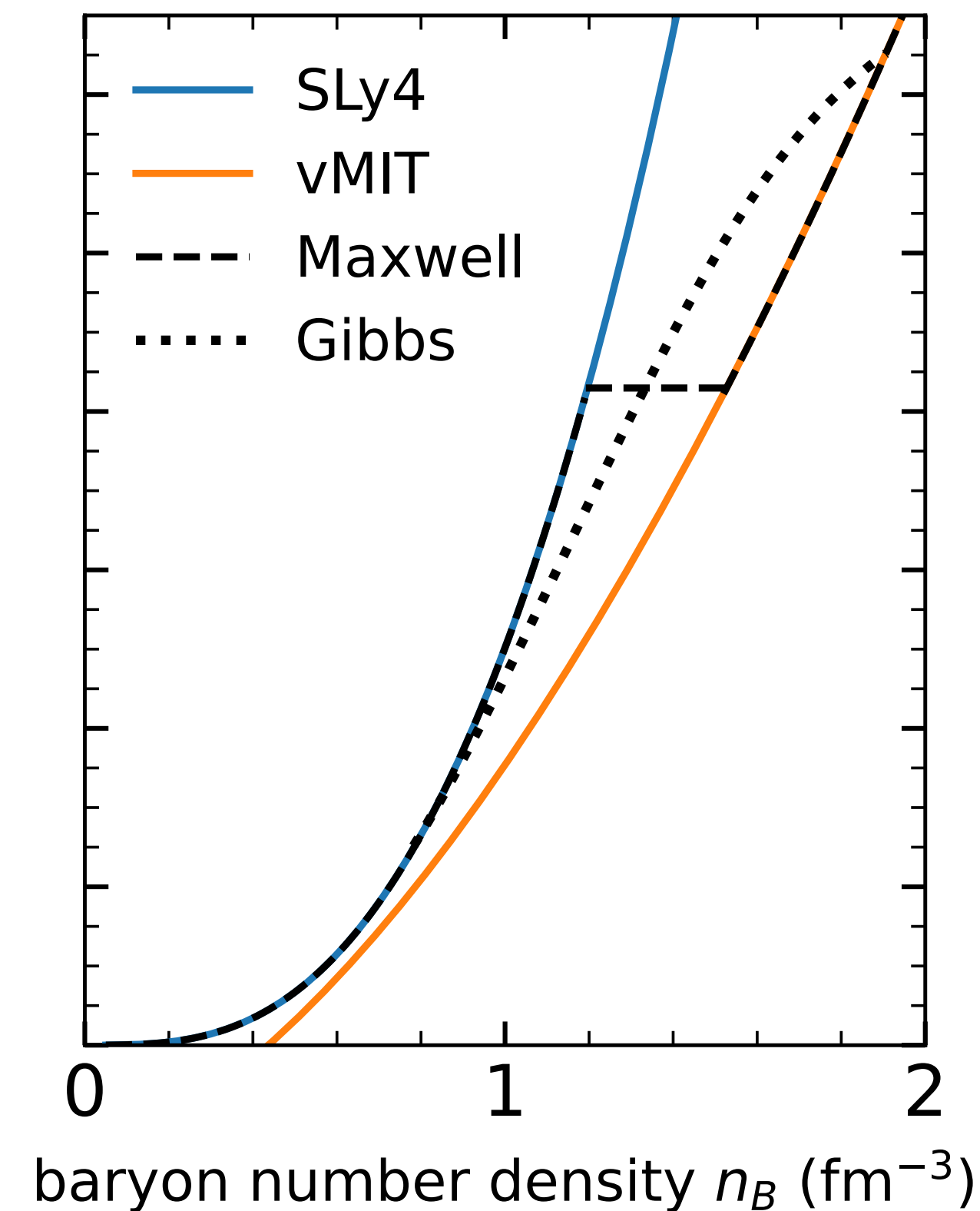
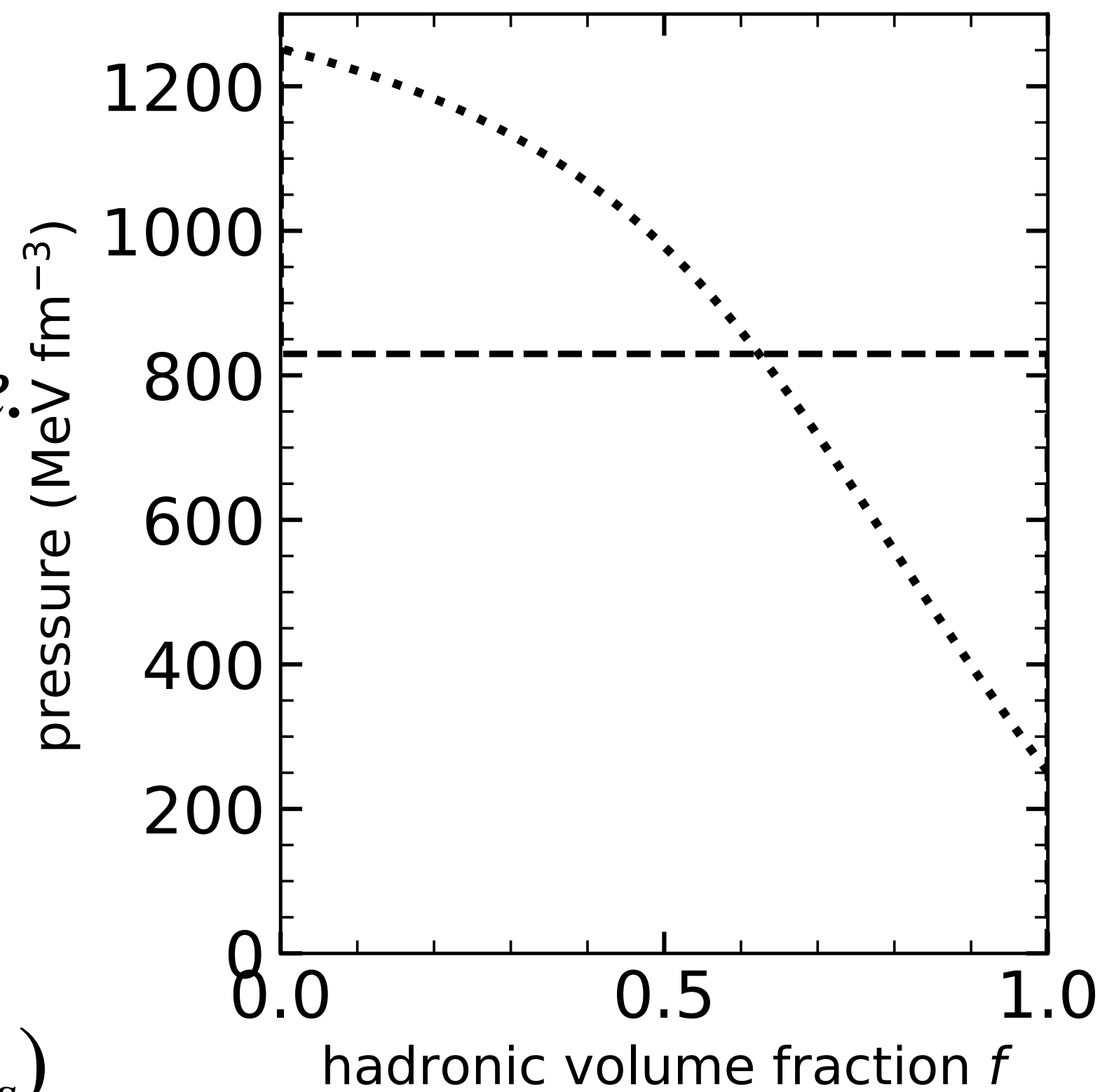
$$n_{e,Q} = \frac{2}{3}n_u - \frac{1}{3}n_d - \frac{1}{3}n_s$$

$$n_{e,N} = n_p$$
- Leptons *aren't* balanced at the interface.
- Energy *isn't* minimized!

- Global charge neutrality (Gibbs):

$$n_e = fn_{e,N} + (1 - f)n_{e,Q}$$

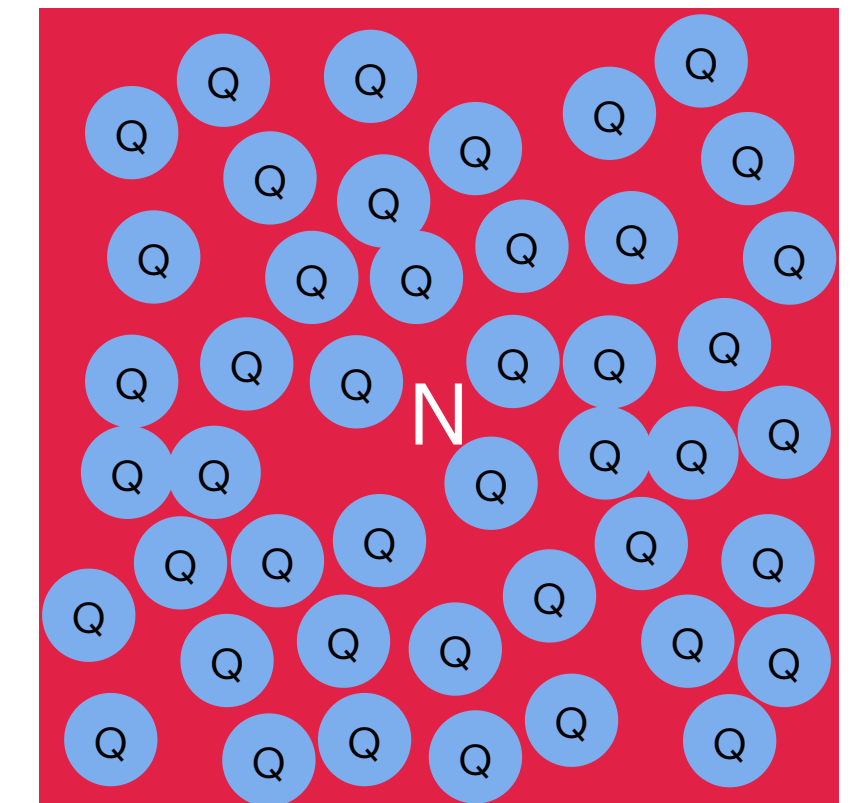
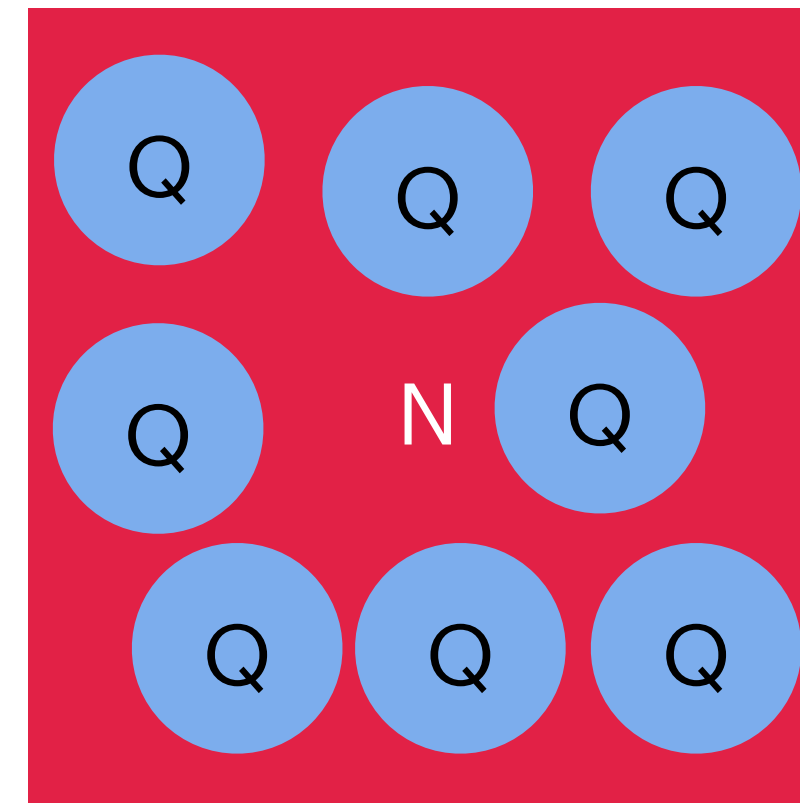
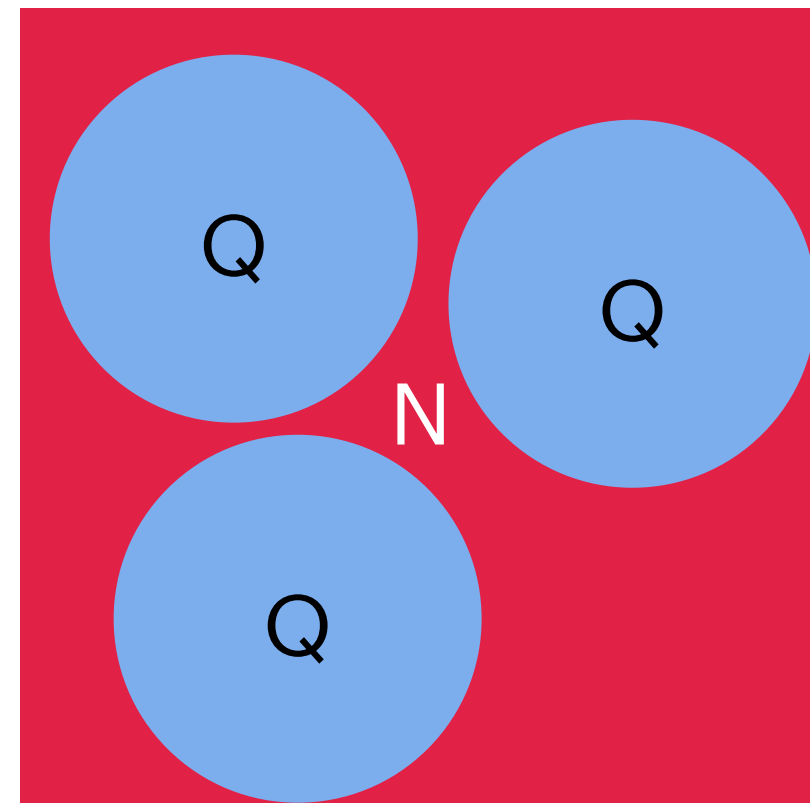
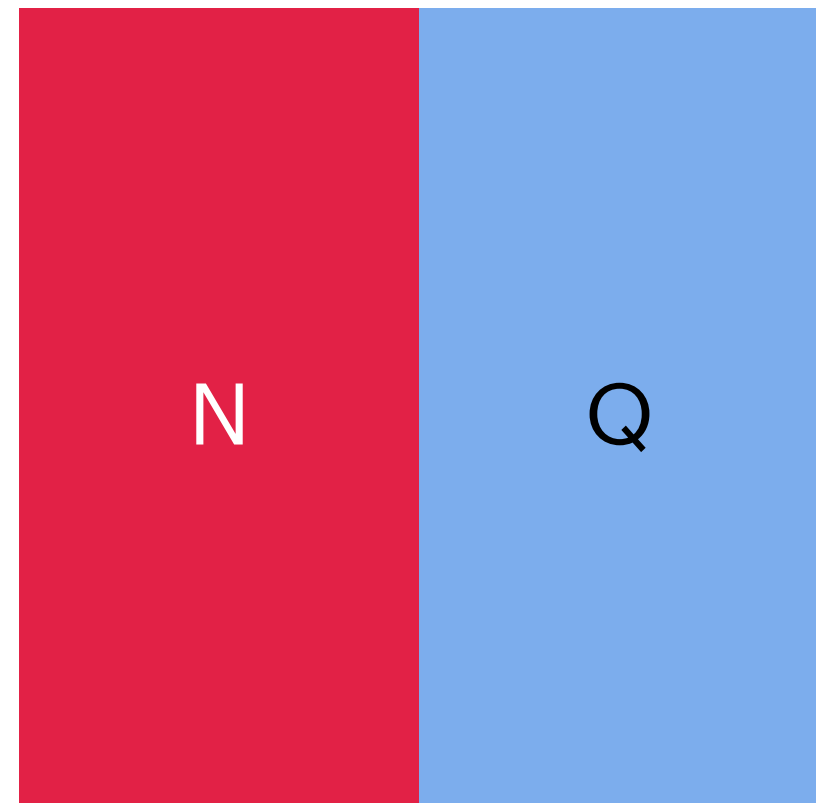
$$n_B = f(n_p + n_n) + \frac{1 - f}{3}(n_u + n_d + n_s)$$



Problem of Gibbs Construction

- e.g. volume fraction $f = 0.5$:

$$n_{e,N} = n_p \quad n_{e,Q} = \frac{2n_u - n_d - n_s}{3}$$



$$n_e = fn_{e,N} + (1 - f)n_{e,Q}$$

Surface energy increases \longrightarrow

\longleftarrow Coulomb energy increases

- Gibbs construction assumes infinite mixing leading to infinite boundary.
- Gibbs construction is realistic only when surface tension is negligibly small.
- The amount of boundary reflected on the charge neutrality condition.

Between Maxwell & Gibbs

Partially local & partially global

- Locally neutral lepton densities:

$$n_{e,N} = n_p, \quad n_{e,Q} = \frac{2}{3}n_u - \frac{1}{3}n_d - \frac{1}{3}n_s$$

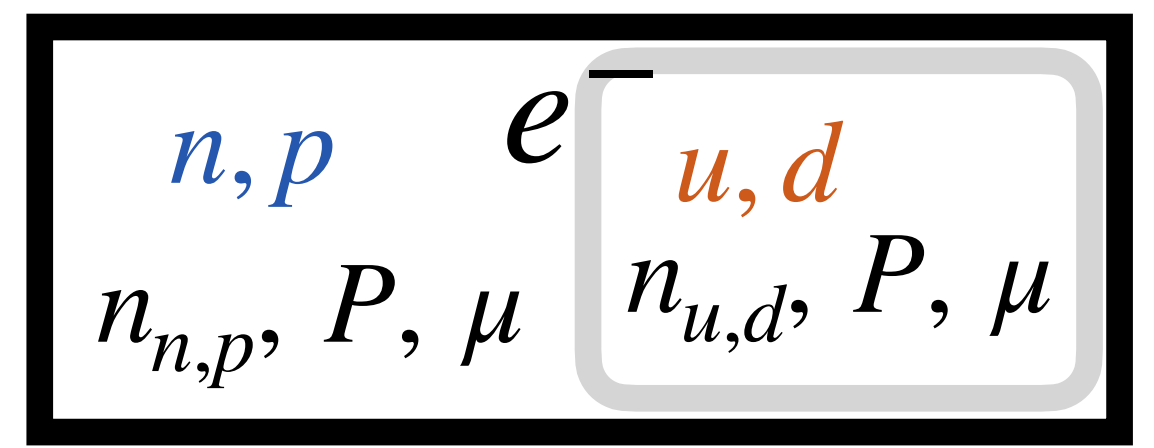
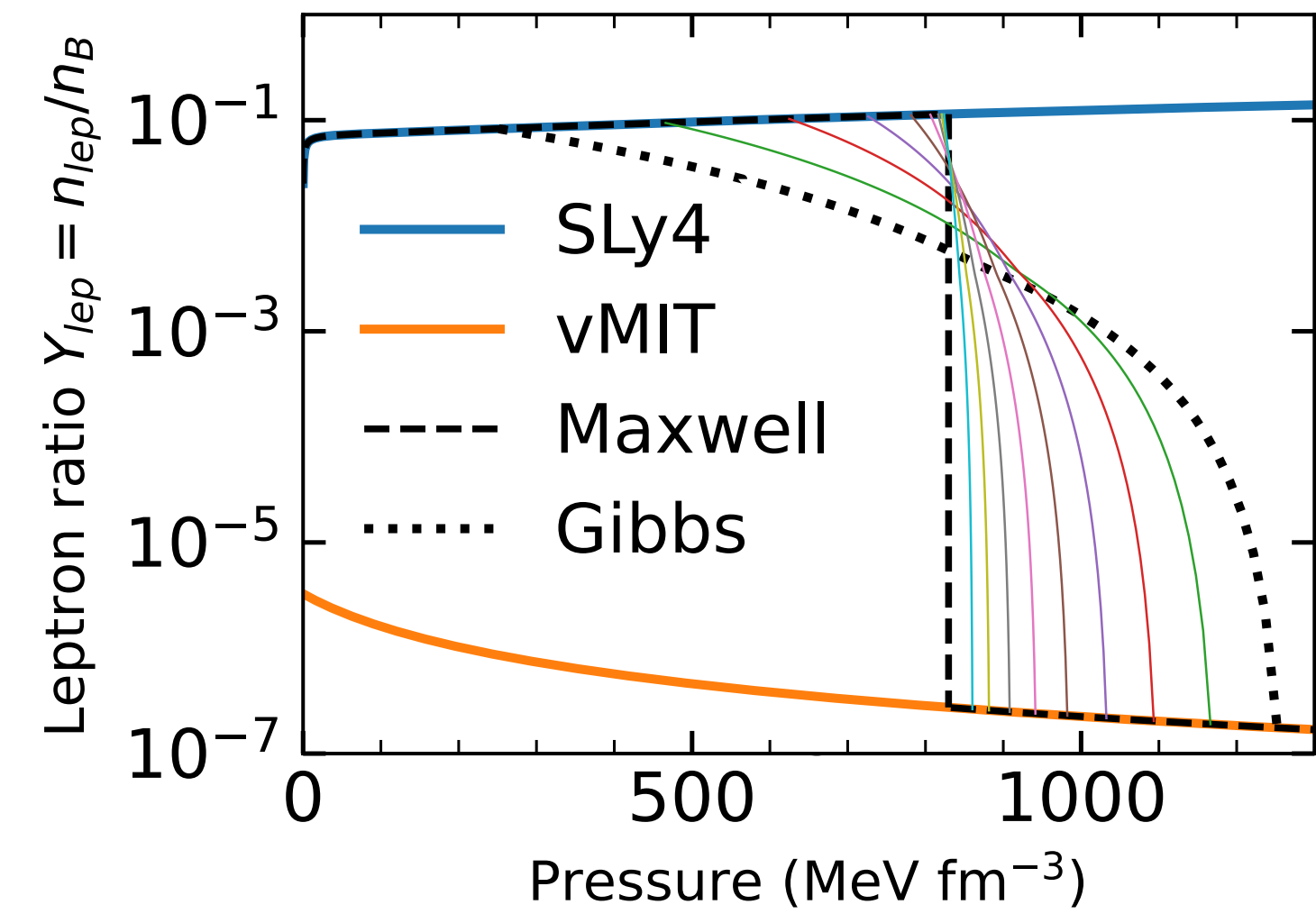
- Global lepton density, $n_{e,G}$

- Total lepton density:

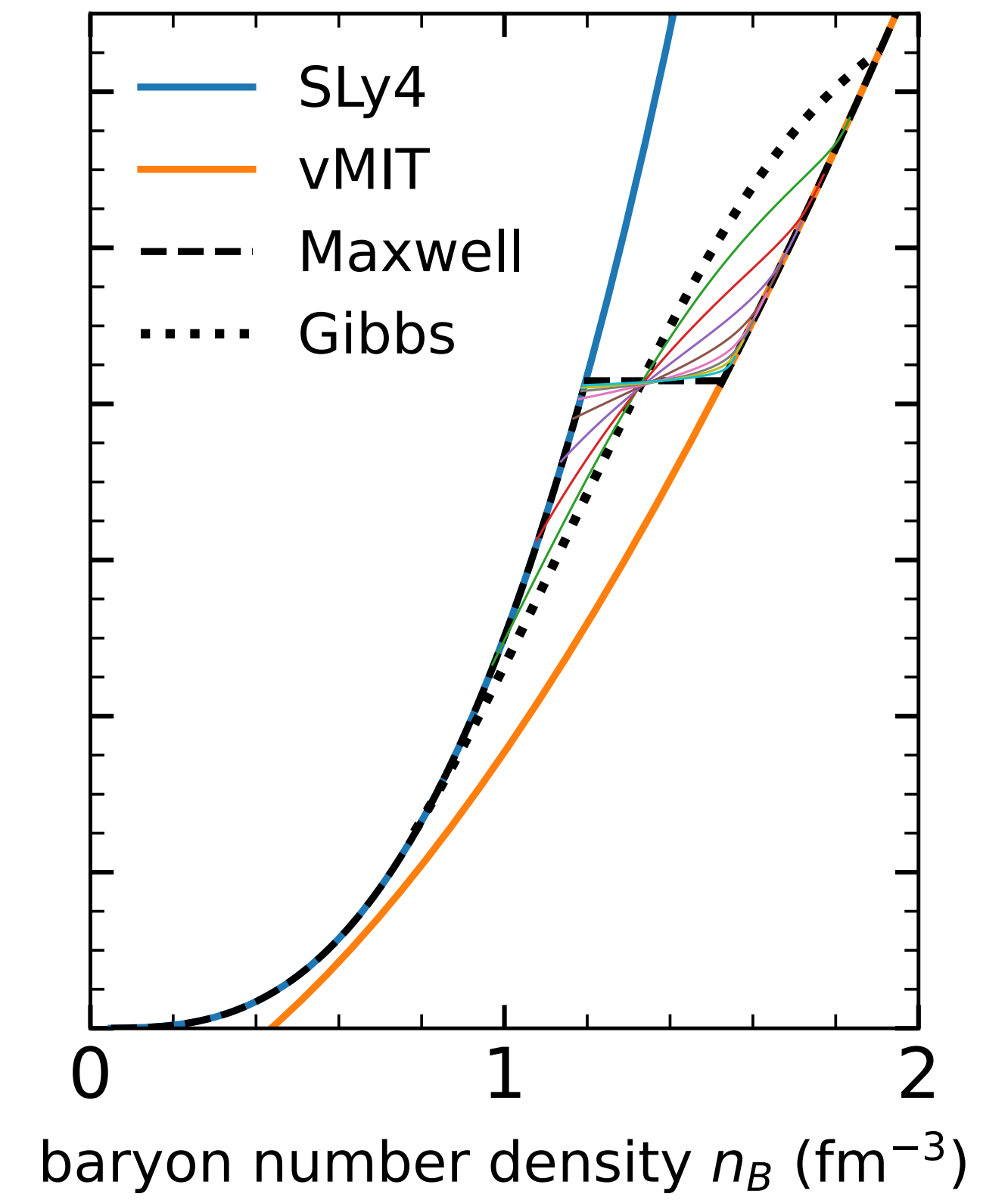
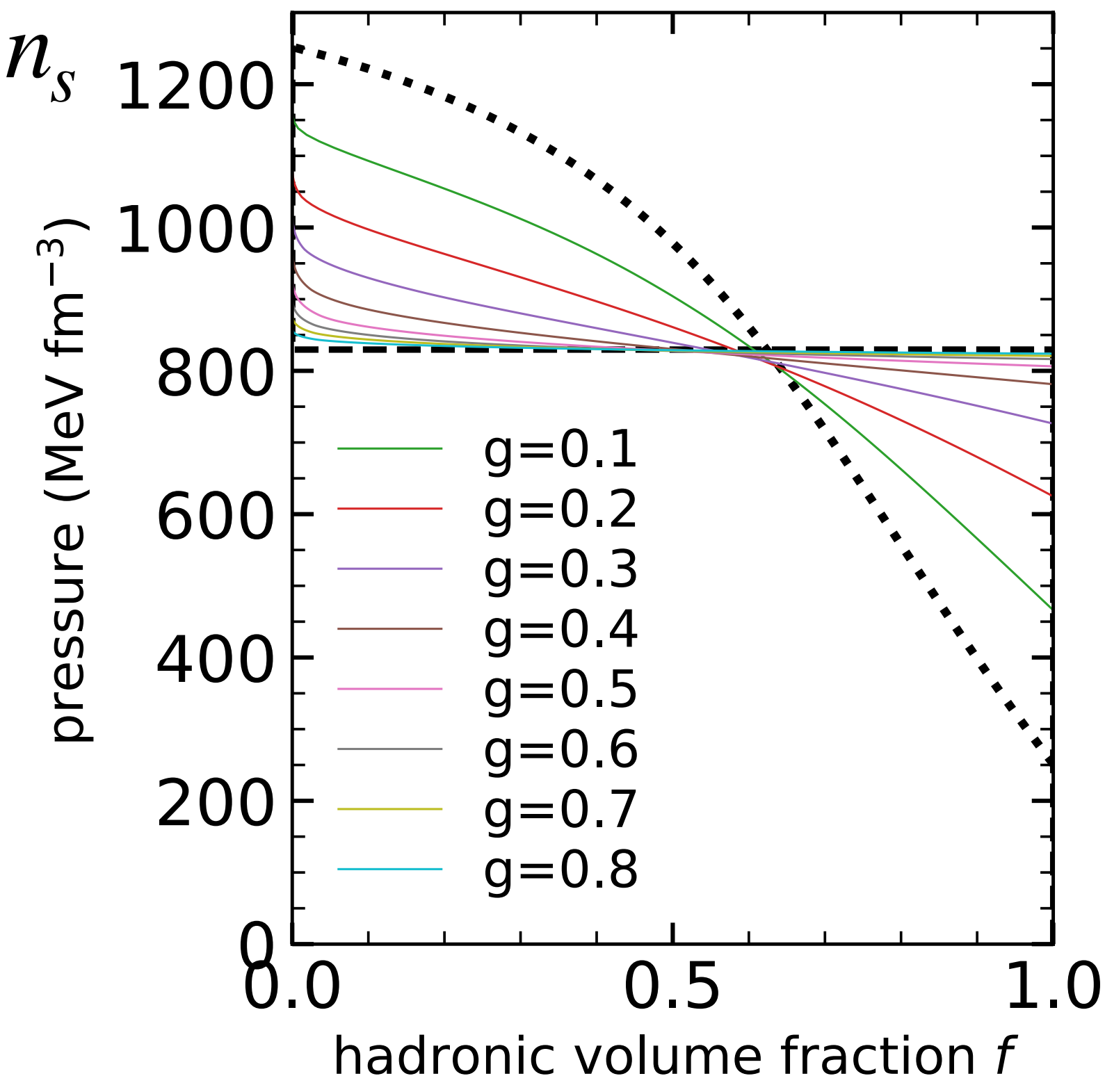
$$n_e = g(fn_{e,N} + (1-f)n_{e,Q}) + (1-g)n_{e,G}$$

- $g = 0 \rightarrow$ Gibbs transition
- $g = 1 \rightarrow$ Maxwell transition

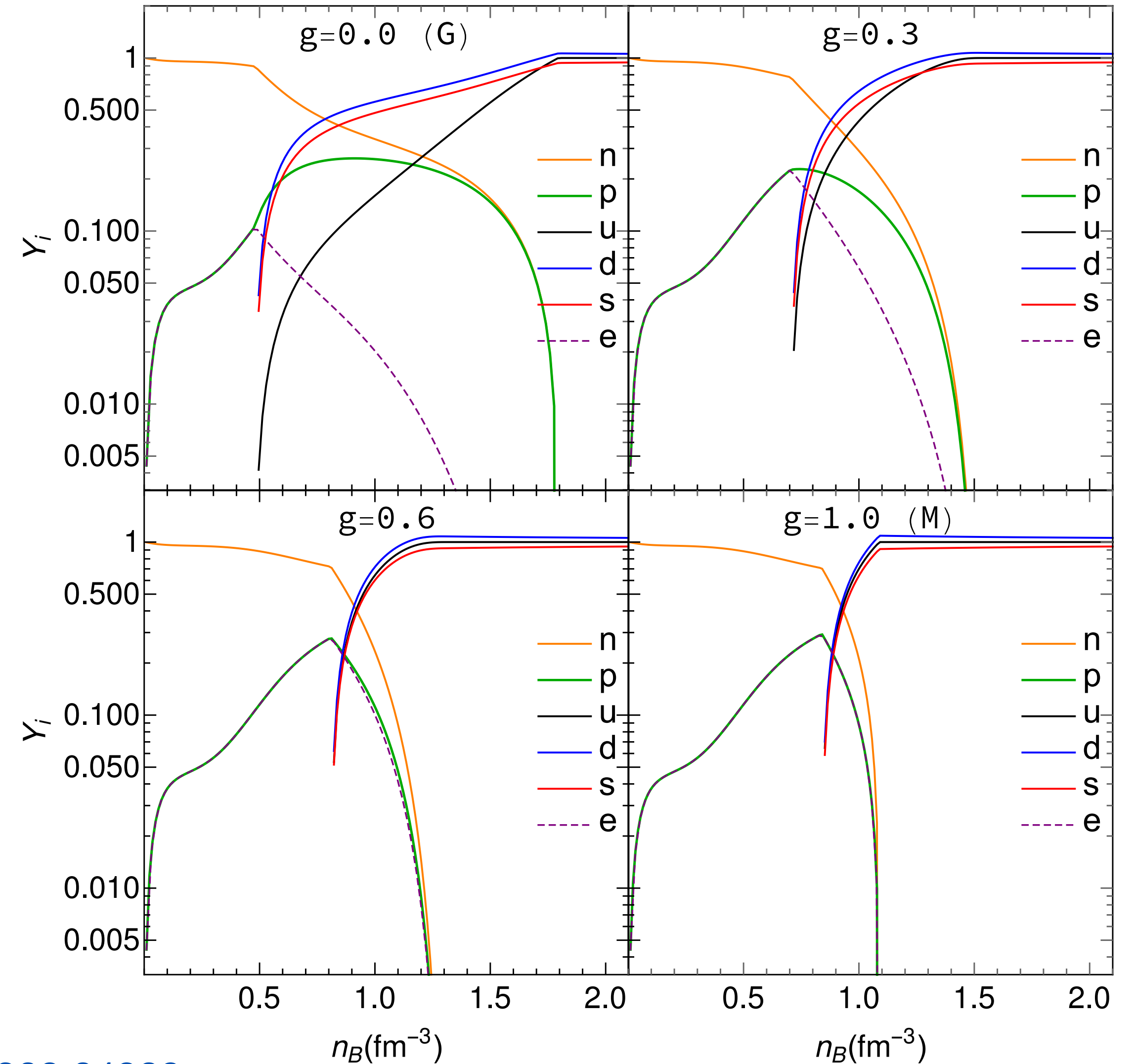
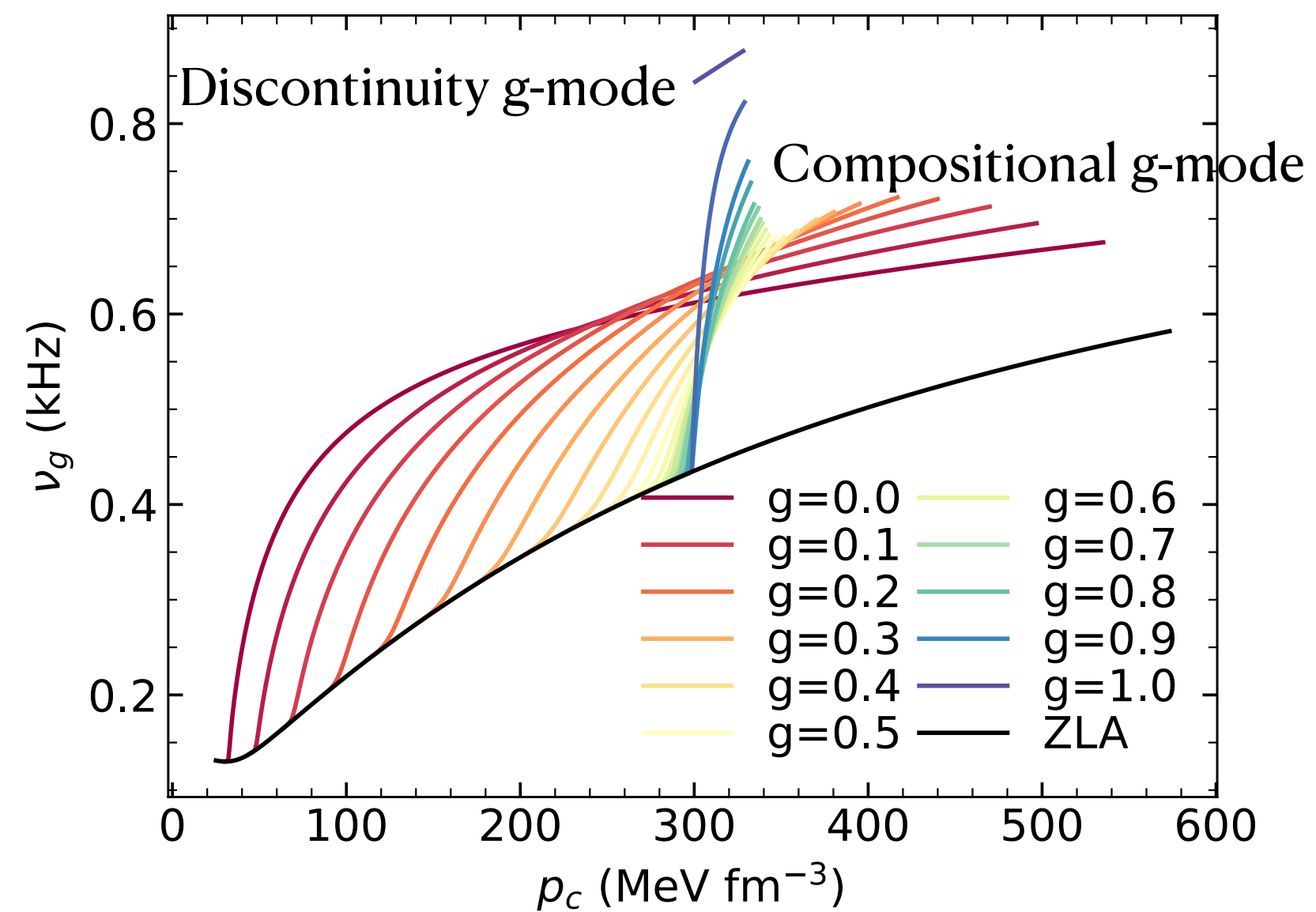
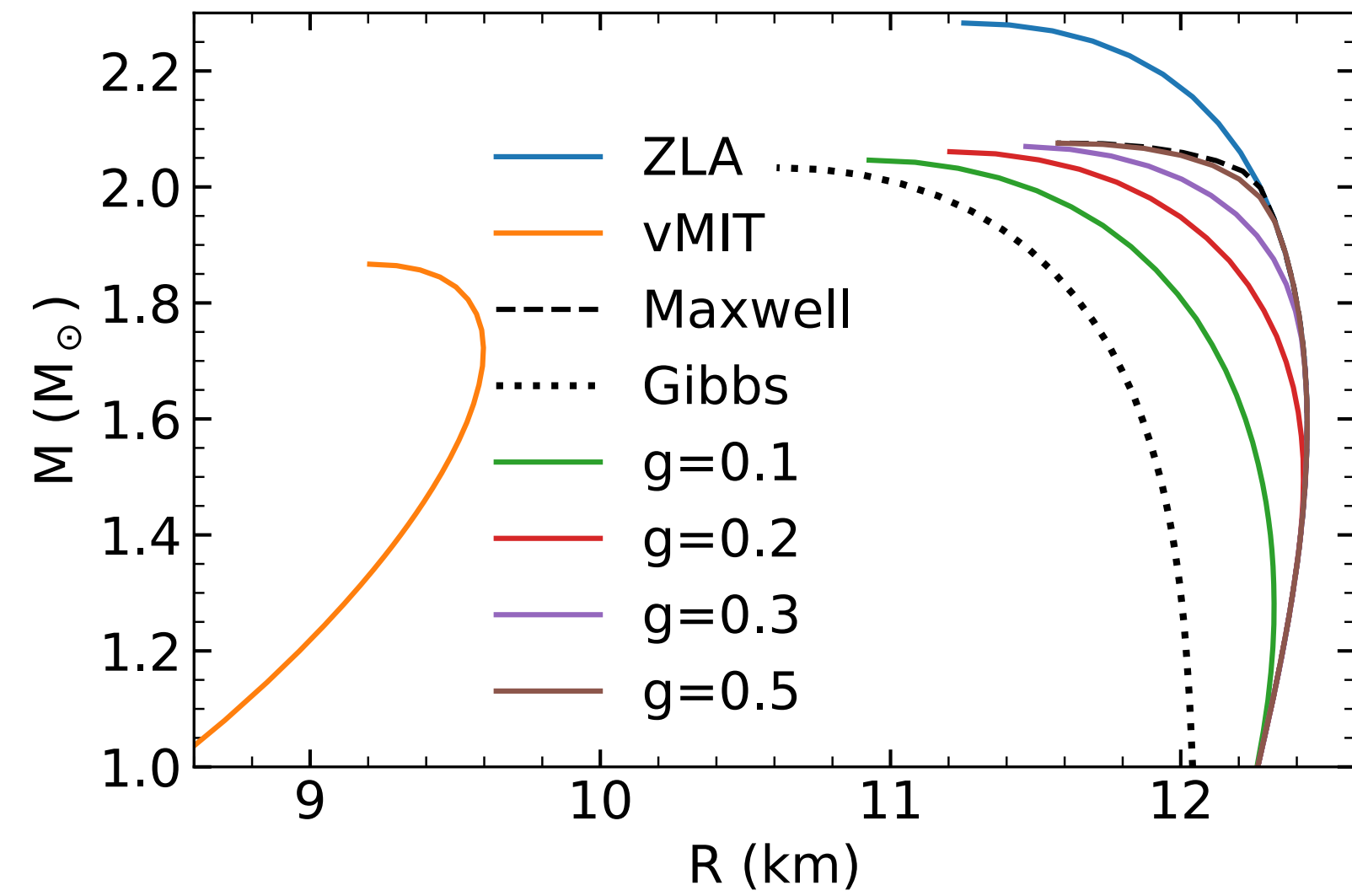
- g could be determined by Surface & Coulomb energy.



Global lepton fraction
 $1 - g$



Between Maxwell & Gibbs



Between Maxwell & Gibbs

Extend to finite temperature:

- Relativistic Fermi integrals, JEL polynomials.

- Introduce anti-particles as,

$$\mu_{e^-} = -\mu_{e^+}$$

$$\mu_{\mu^-} = -\mu_{\mu^+}$$

$$\mu_{u^-} = -\mu_{u^+}$$

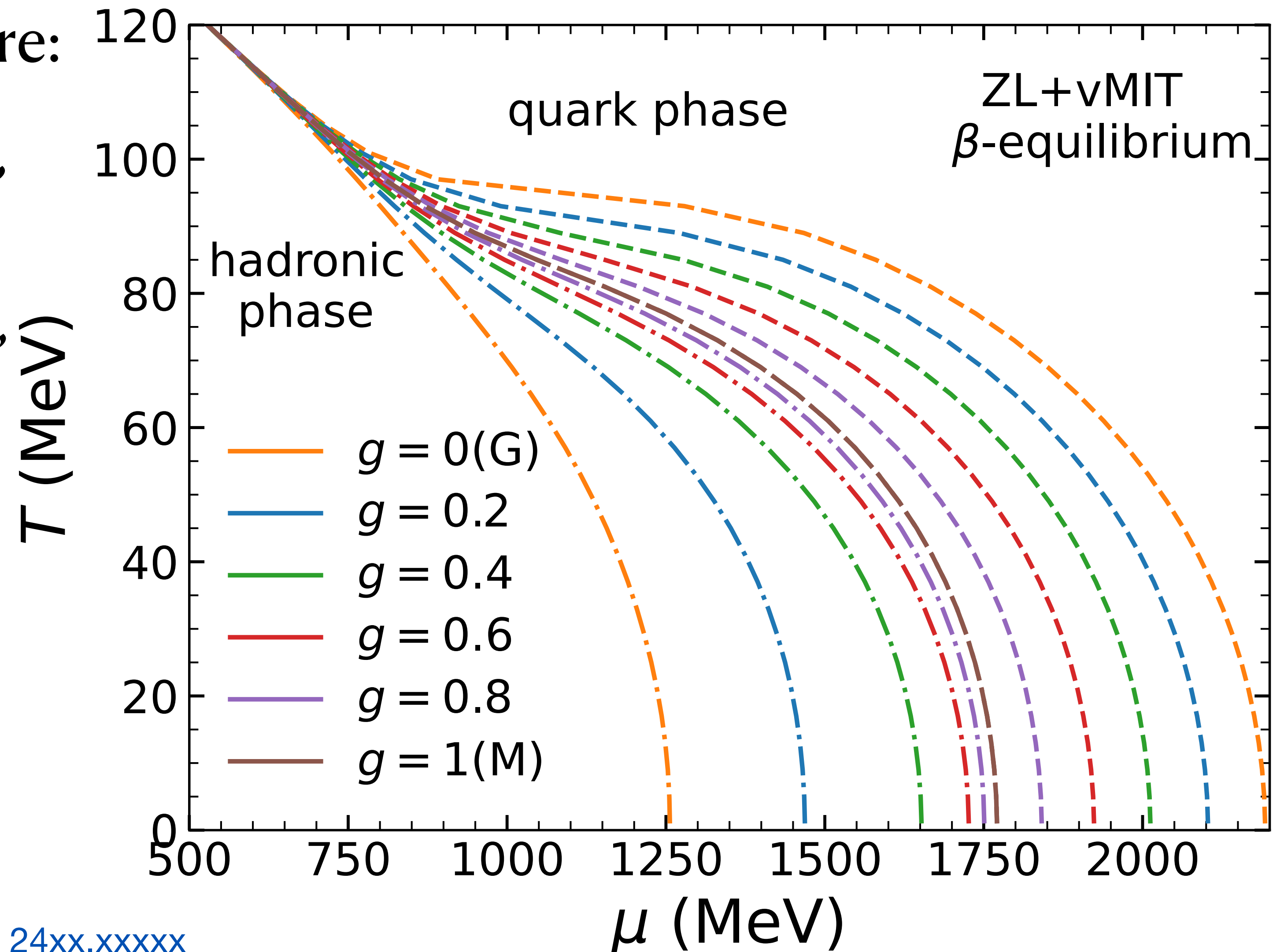
$$\mu_{d^-} = -\mu_{d^+}$$

$$\mu_{s^-} = -\mu_{s^+}$$

- Add photon contribution,

$$\varepsilon_{\text{photon}} \propto T^4$$

[arXiv: 24xx.xxxxx](https://arxiv.org/abs/24xx.xxxxx)



Between Maxwell & Gibbs

Extend to off- β -equilibrium:

- Ignore β -equilibrium condition,

$$\mu_d = \mu_u + g\mu_{e,Q} + (1 - g)\mu_{e,G}$$

$$\mu_n = \mu_p + g\mu_{e,N} + (1 - g)\mu_{e,G}$$

- And replace it with,

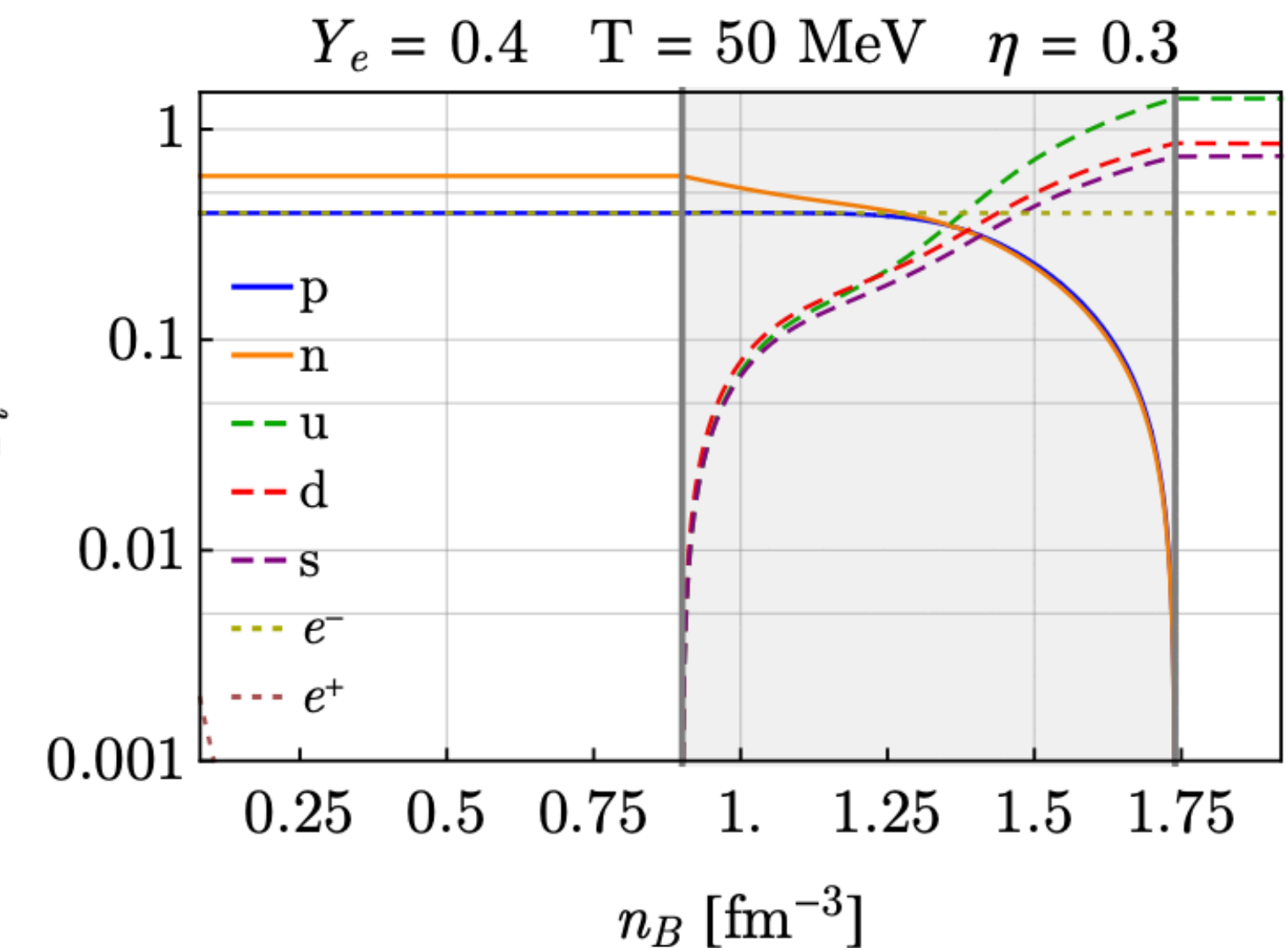
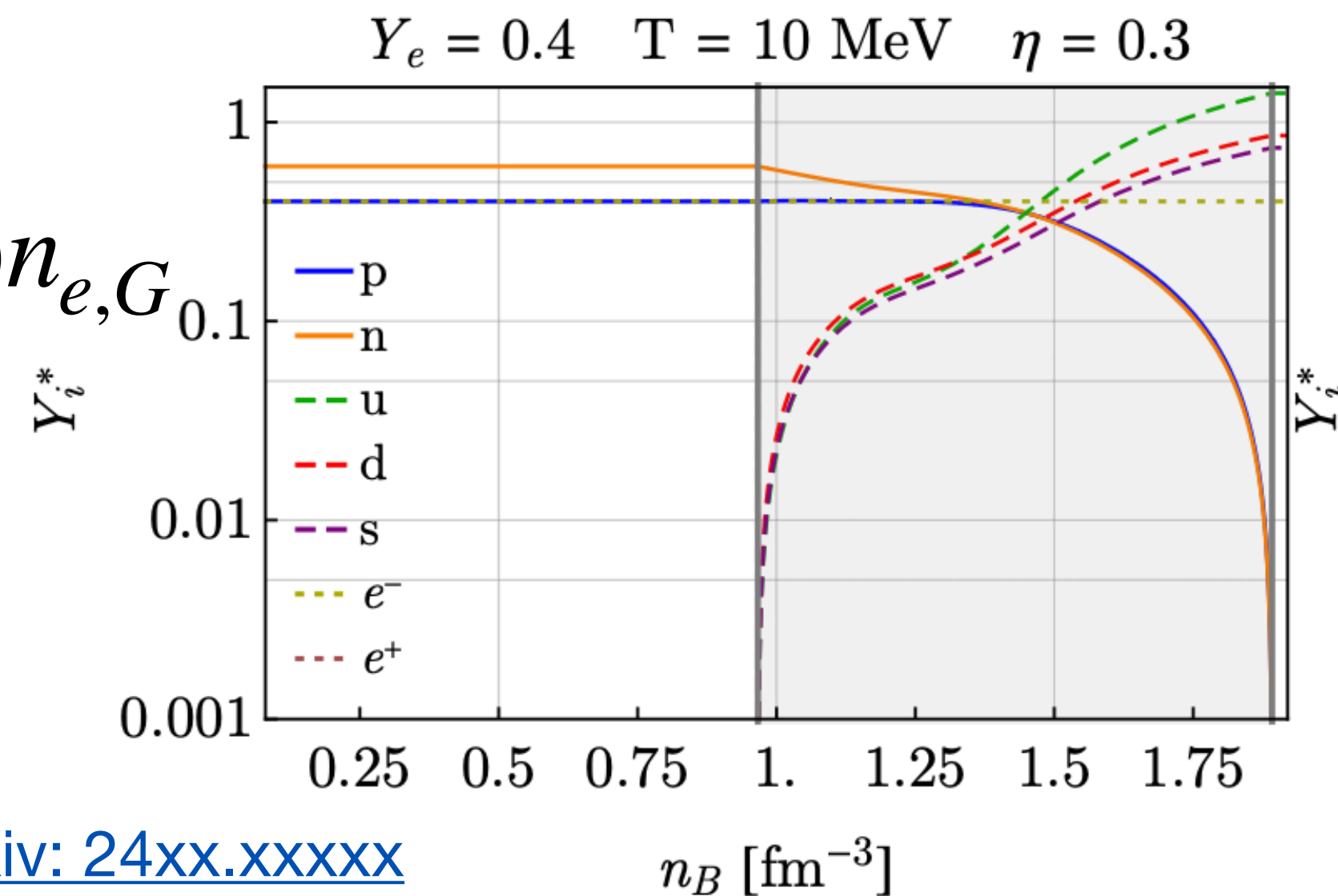
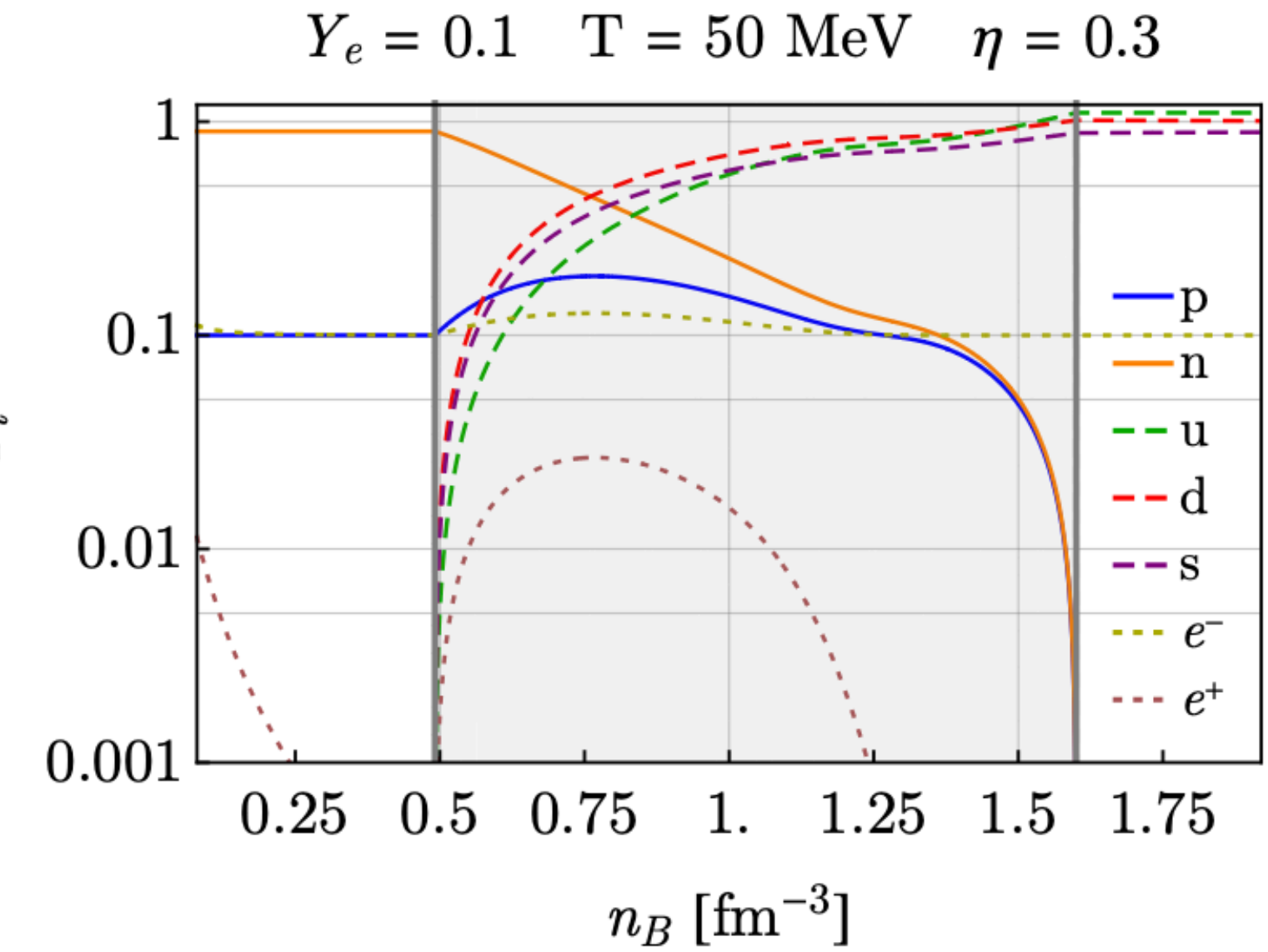
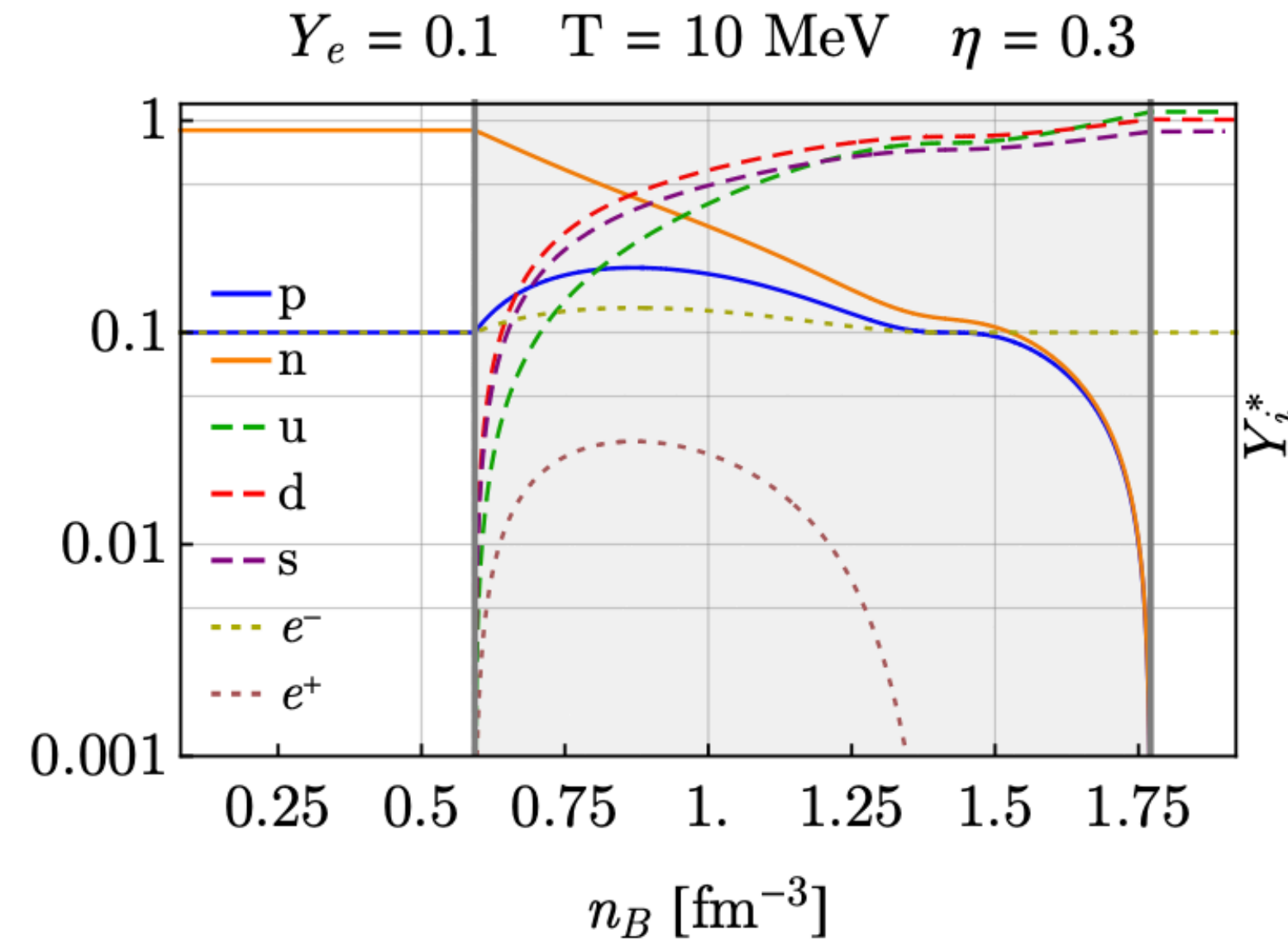
$$n_e = n_B Y_e =$$

$$g(fn_{e,N} + (1 - f)n_{e,Q}) + (1 - g)n_{e,G}$$

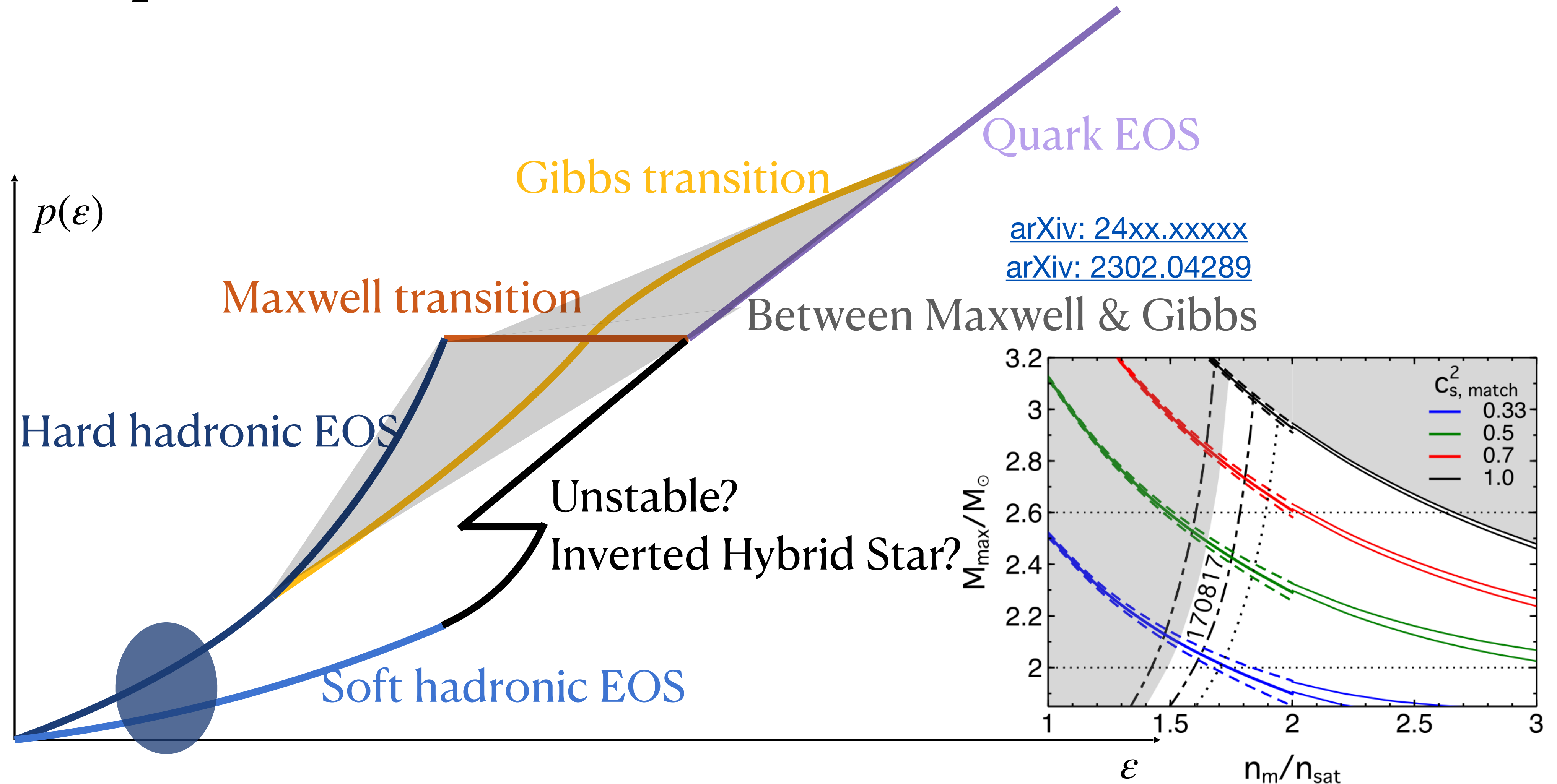
- The final EOS is,

$$\varepsilon = \varepsilon(n_B, Y_e, T)$$

$$p = p(n_B, Y_e, T)$$



Hadron-quark Transition in Neutron Star Core



Soft hadronic EOSs is flavored by ab-initio calculation, nuclear experiments & neutron star merger observation.

[arXiv: 2406.05267](https://arxiv.org/abs/2406.05267)

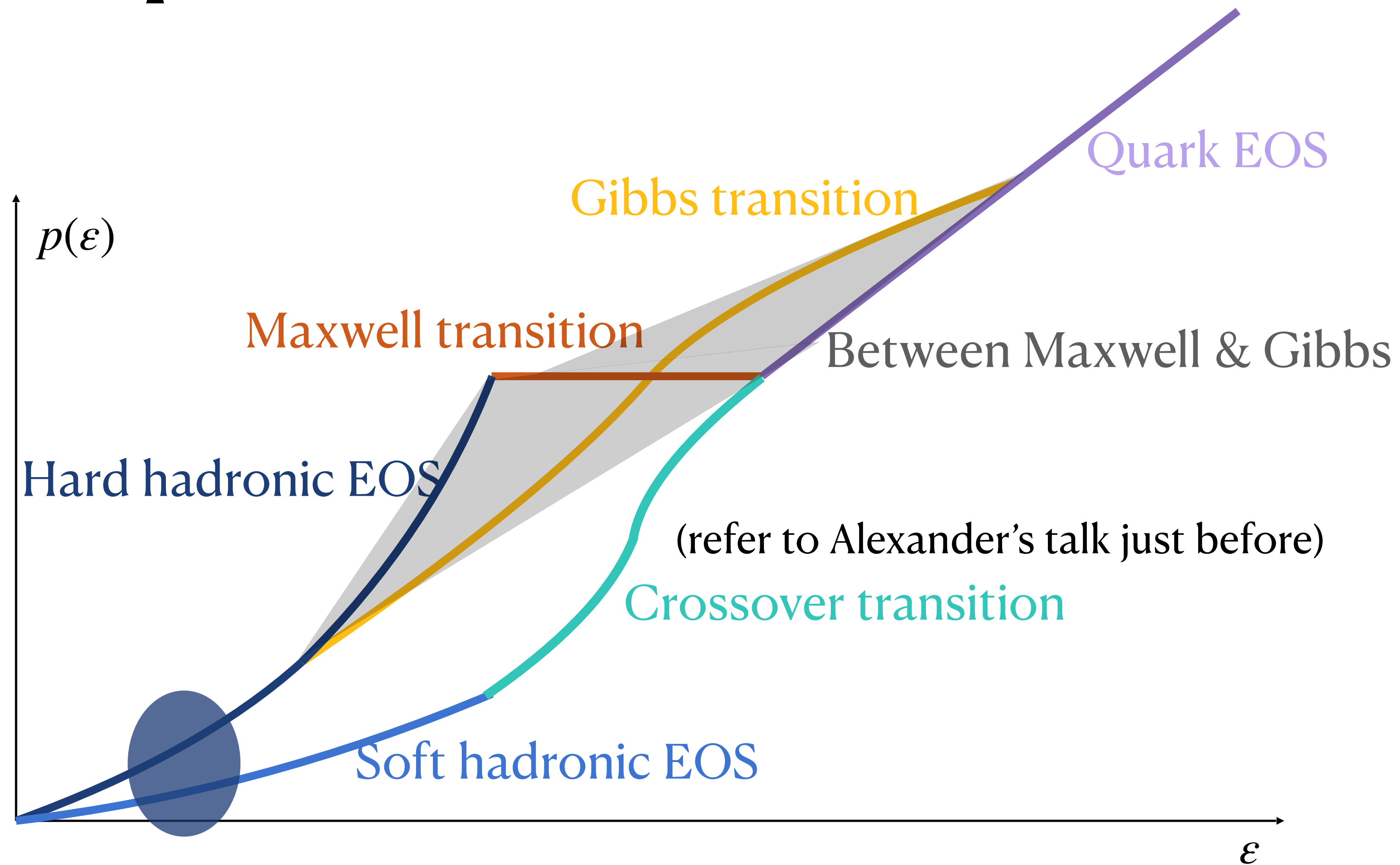
[arXiv: 1808.02858](https://arxiv.org/abs/1808.02858)

[arXiv: 2009.06441](https://arxiv.org/abs/2009.06441)

[arXiv: 24xx.xxxxx](https://arxiv.org/abs/24xx.xxxxx)

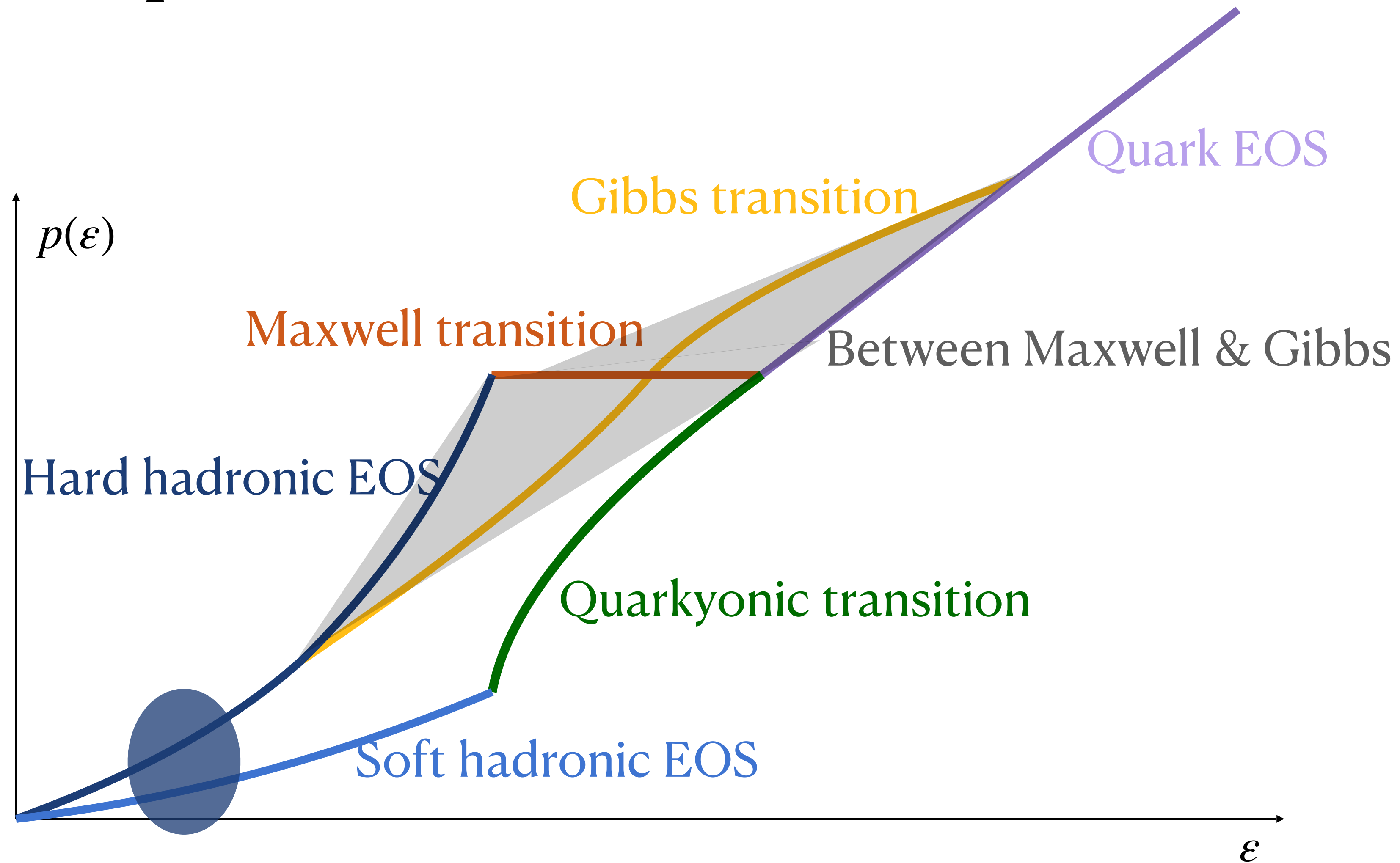
[arXiv: 2302.04289](https://arxiv.org/abs/2302.04289)

Hadron-quark Transition in Neutron Star Core



Soft hadronic EOSs is flavored by ab-initio calculation, nuclear experiments & neutron star merger observation.

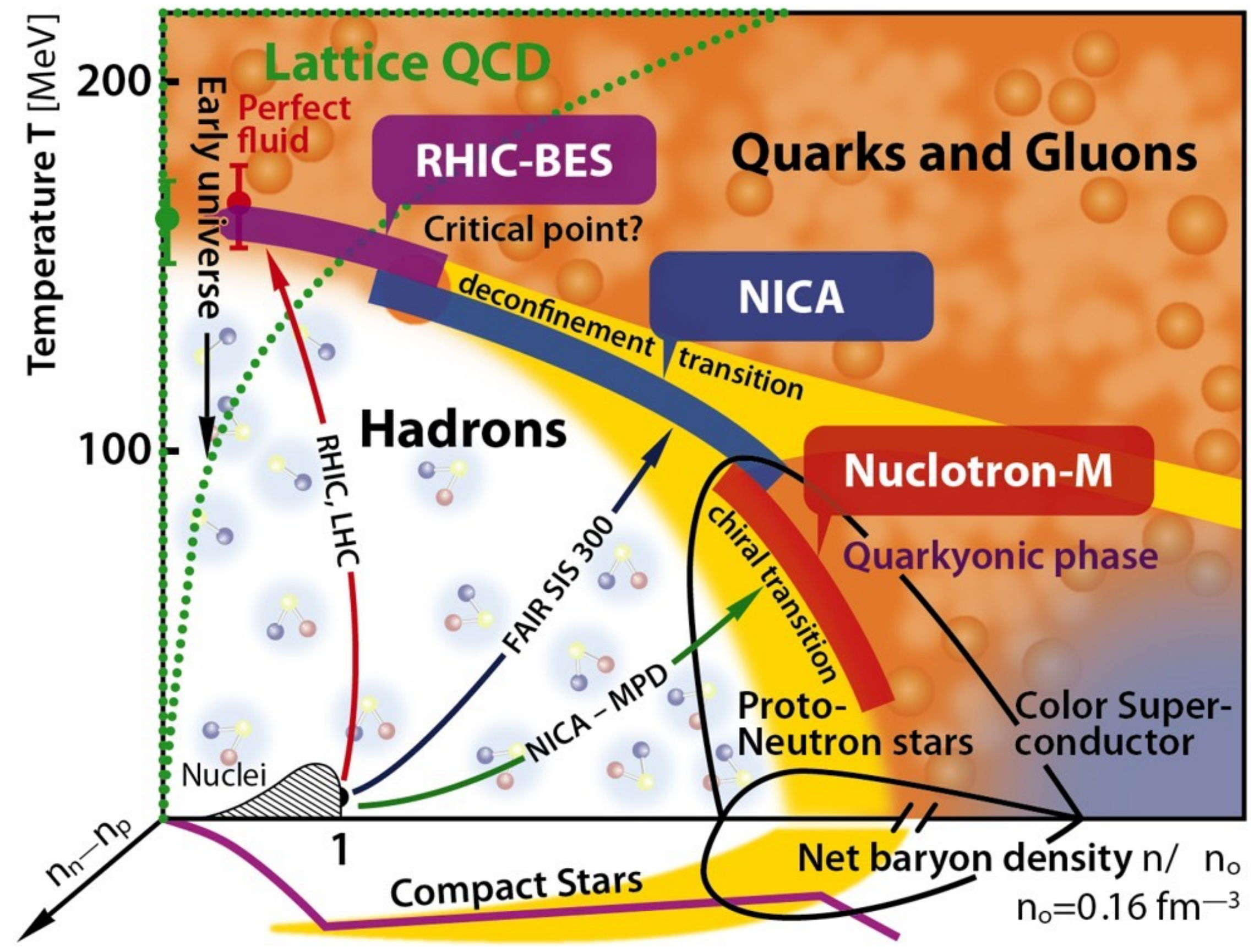
Hadron-quark Transition in Neutron Star Core



Soft hadronic EOSs is flavored by ab-initio calculation, nuclear experiments & neutron star merger observation.

Quarkyonic Matter

- The hypothetical phase between hadronic matter and deconfined quark matter (David Blaschke 2008).



<https://nica.jinr.ru/physics.php>

Sanjay and McLerran 2018

Dynamical realization:

K. Jeong et. al. 2020

T. Kojo et. al. 2021

Y. Fujimoto et. al. 2023

Extend isospin, flavor, finite T:

Zhao & Lattimer 2020

S. Sen et. al. 2021

D. Duarte et. al. 2021

J. Margueron et. al. 2021

Include better hadronic EOS:

G. Cao et. al. 2021

A. Kumar et. al. 2022

C. Xia et. al. 2023

Asymptotic Free

Gross, Wilczek and Politzer 1973

- QCD beta function:

$$\beta(\alpha_s) = q^2 \frac{\partial \alpha_s}{\partial q^2} = -\beta_0 \alpha_s^2 - \beta_1 \alpha_s^3 - \dots$$

$$\text{where } \alpha_s = \frac{g^2}{4\pi}, \beta_0 = \frac{33 - 2N_f}{12\pi} > 0, \beta_1 = \frac{153 - 19N_f}{24\pi^2} > 0$$

- Keep only the first term on the right-hand side,

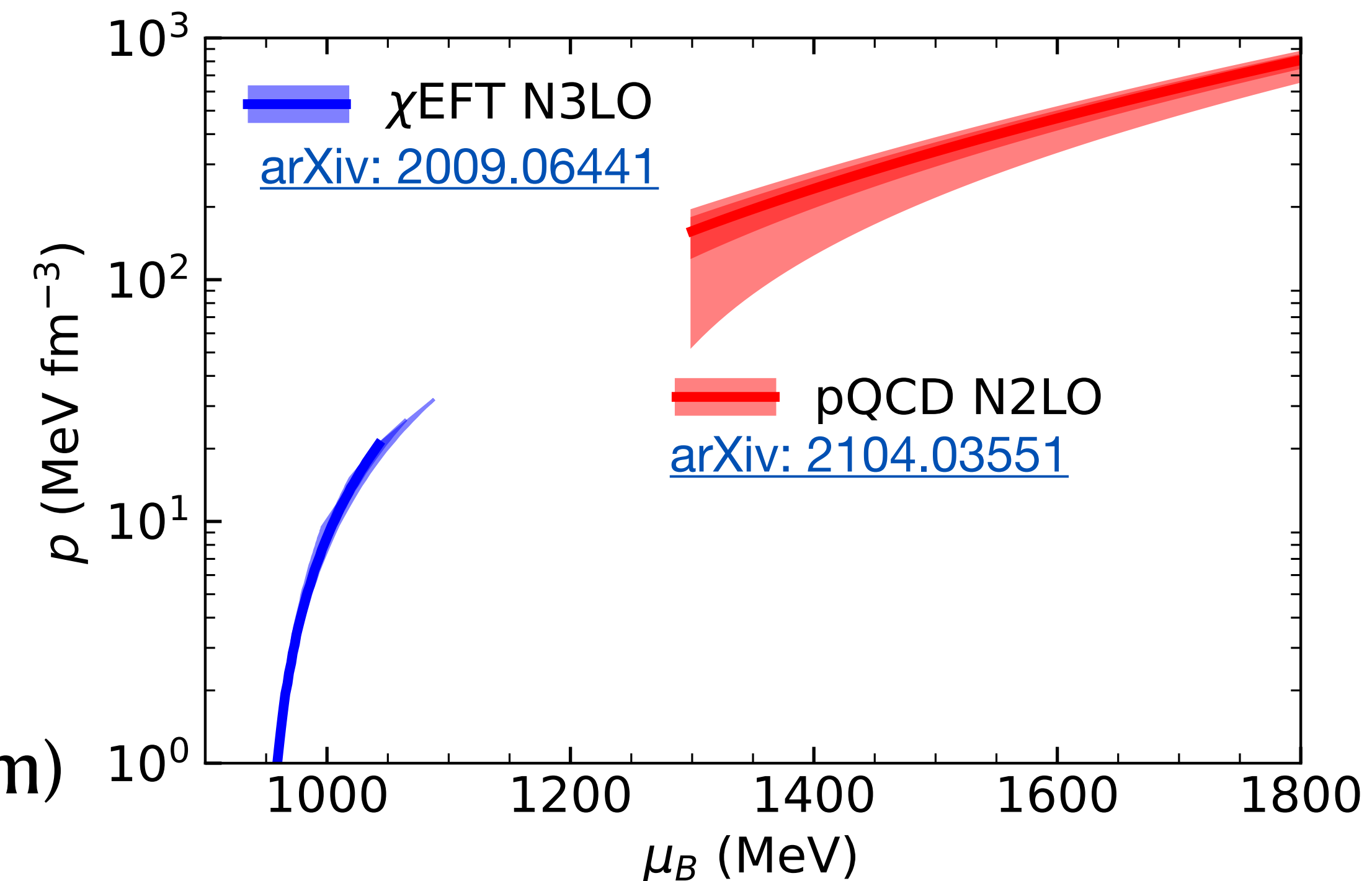
$$\alpha_s \approx \frac{1}{\beta_0 \log q^2 / \Lambda_{QCD}^2}$$

$$\text{therefore } \lim_{q \gg \Lambda_{QCD}} \alpha_s(q) \rightarrow 0$$

- Perturbative QCD:

QCD Lagrangian (quark-gluon coupling)

+ Analytical method (vacuum and ring diagram)

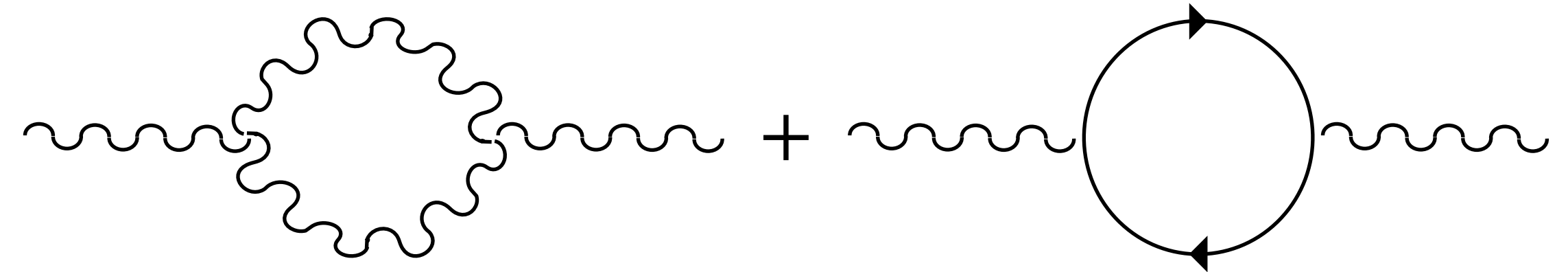


Speculation from large N_c

McLerran & Pisarski 2007

- Confinement due to screening of gluons

$$m_{Debye}^2 \approx \Pi = g^2 \left[\left(N_c + \frac{N_f}{2} \right) \frac{T^2}{3} + \frac{N_f \mu^2}{2\pi^2} + \dots \right]$$



- Large $m_{Debye} \longrightarrow$ stronger screening \longrightarrow weaker long-range interactions \longrightarrow deconfinement

- Large N_c limit: $N_c \rightarrow \infty$ while fixing $\lambda_{tHooft} = g^2 N_c$ and N_f :

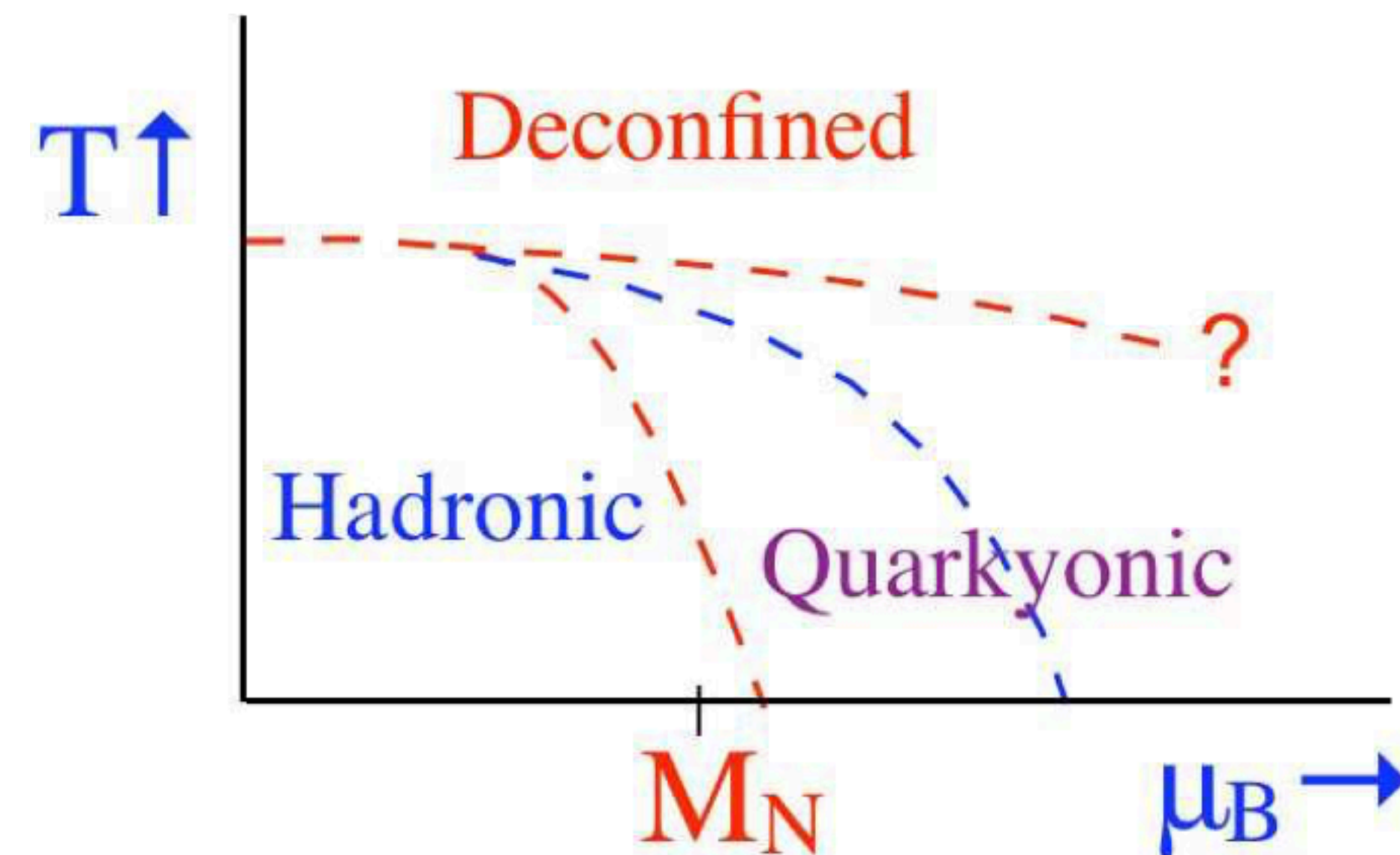
$$m_{Debye}^2 \propto T^2 \text{ for high temperature;}$$

$$m_{Debye}^2 \propto \frac{\mu^2}{N_c} \rightarrow 0 \text{ for high chemical potential.}$$

- Asymptotic free + Confinement (at the same time) ????

Quark + Baryon = Quarkyonic matter

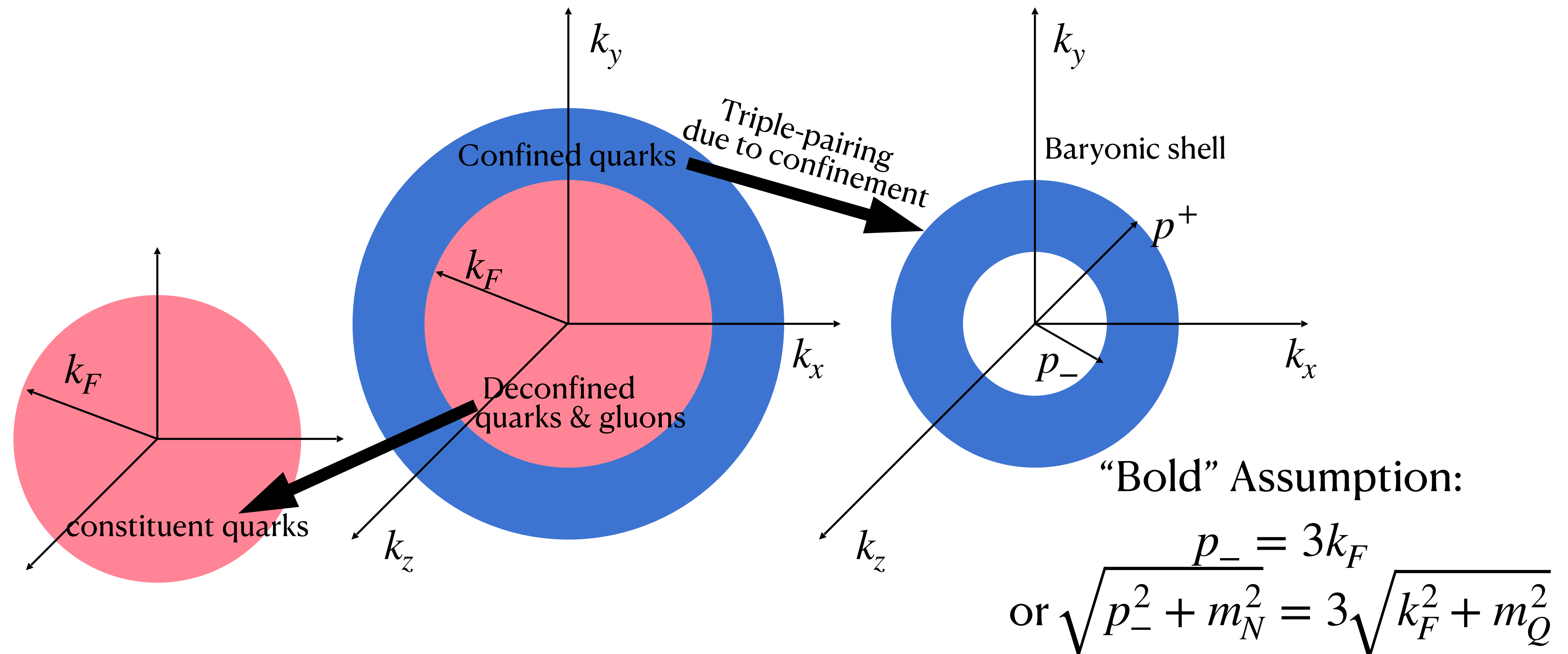
(refer to Larry's talk tomorrow for more details)

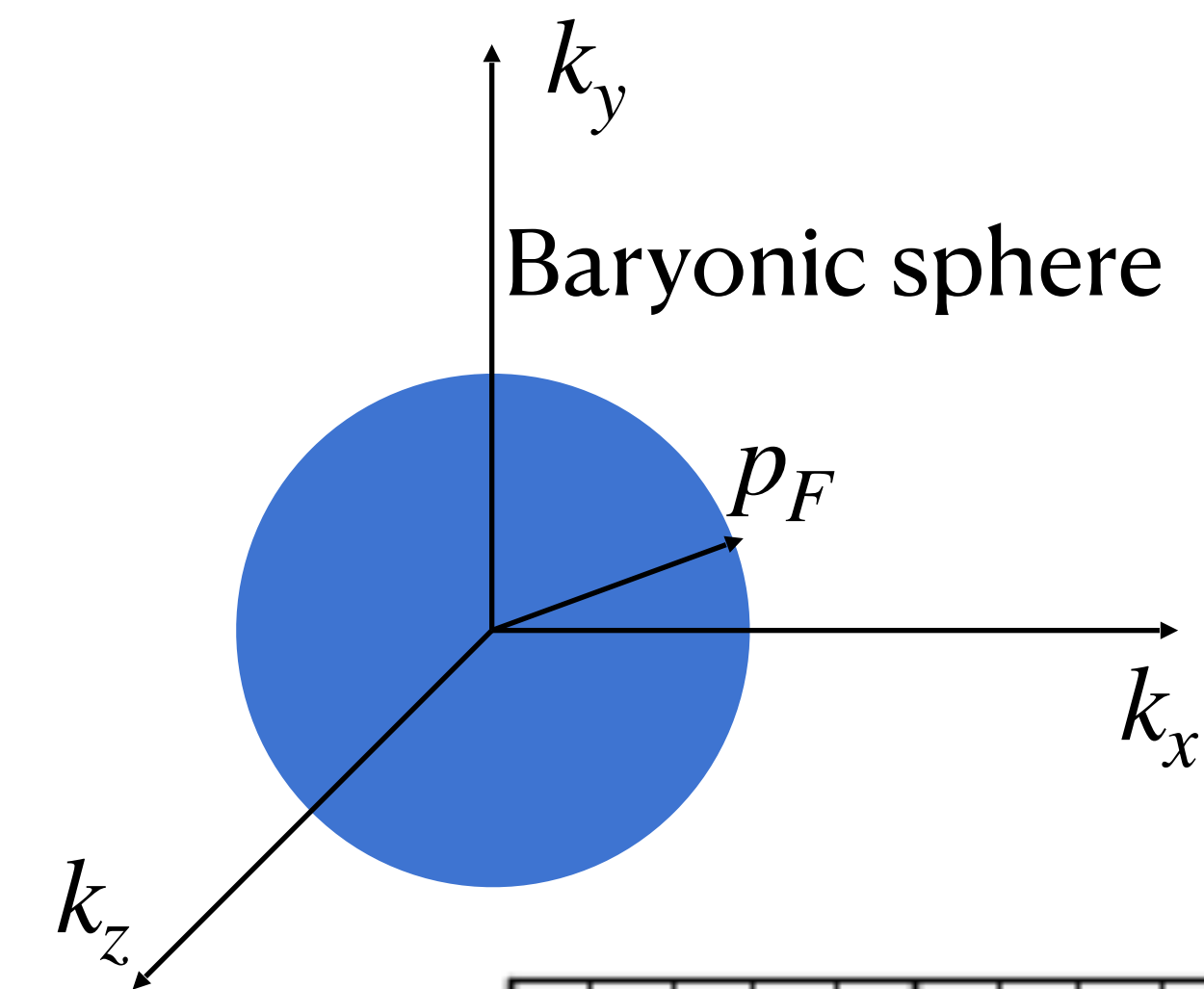


Quarkyonic Matter Momentum Space

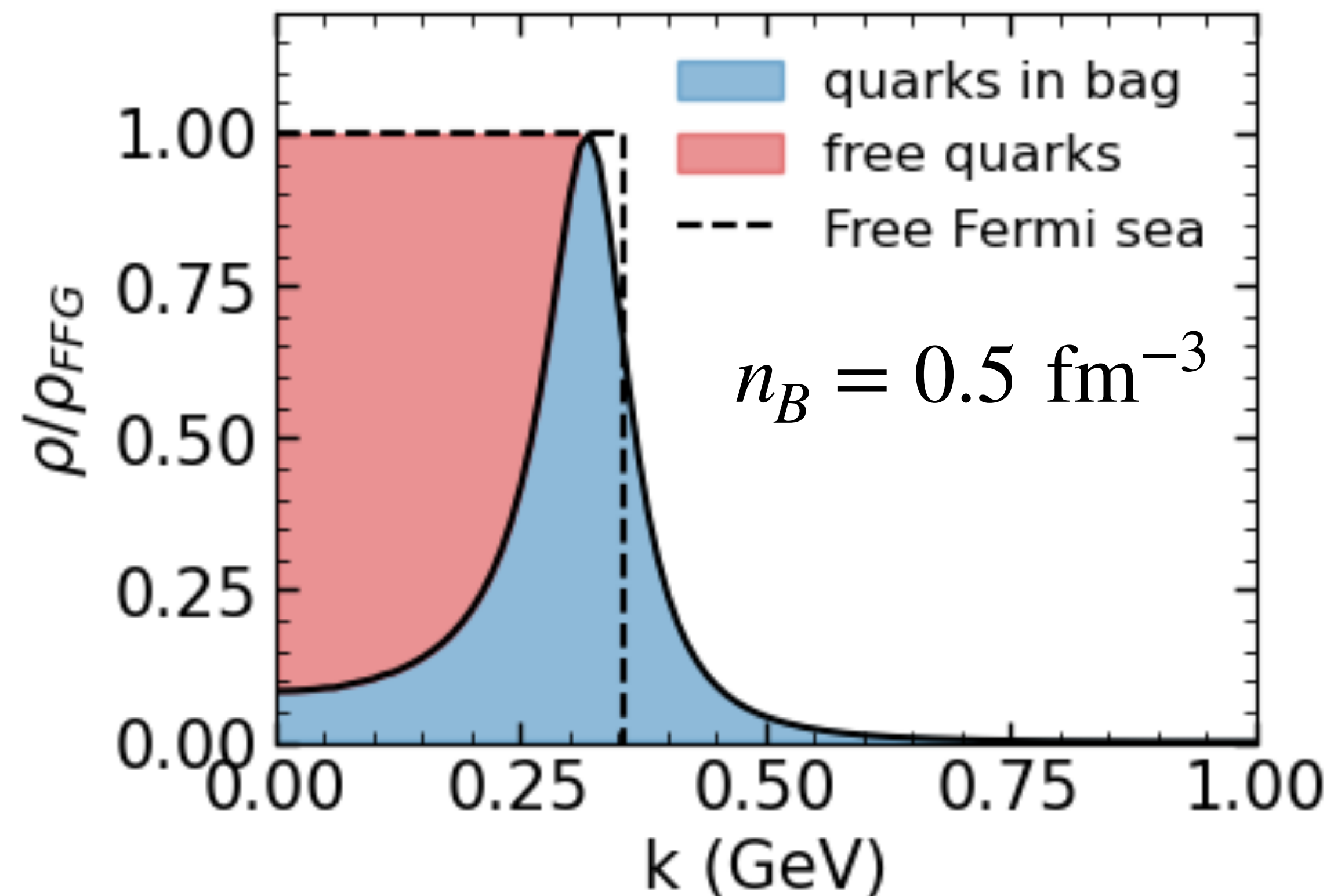
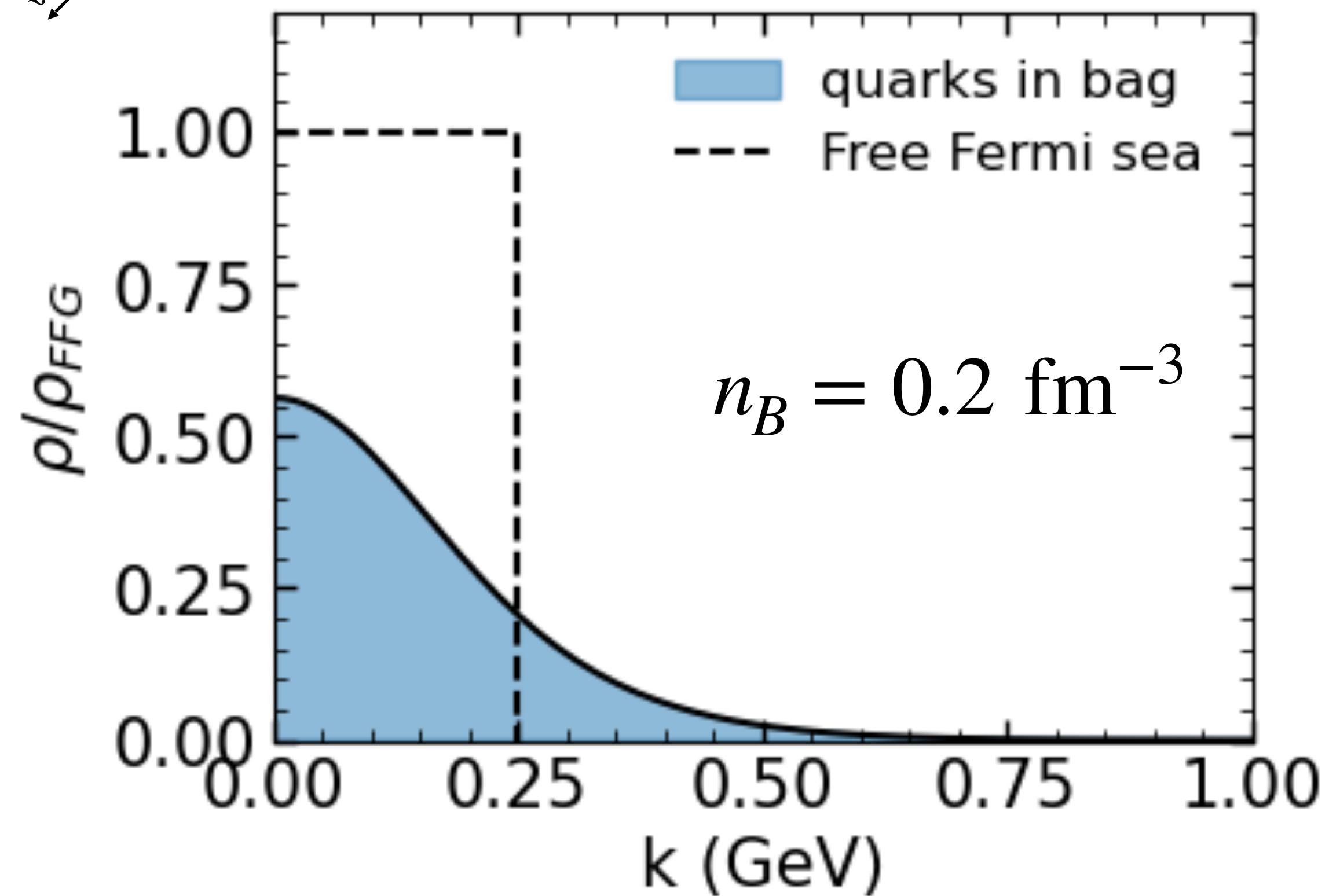
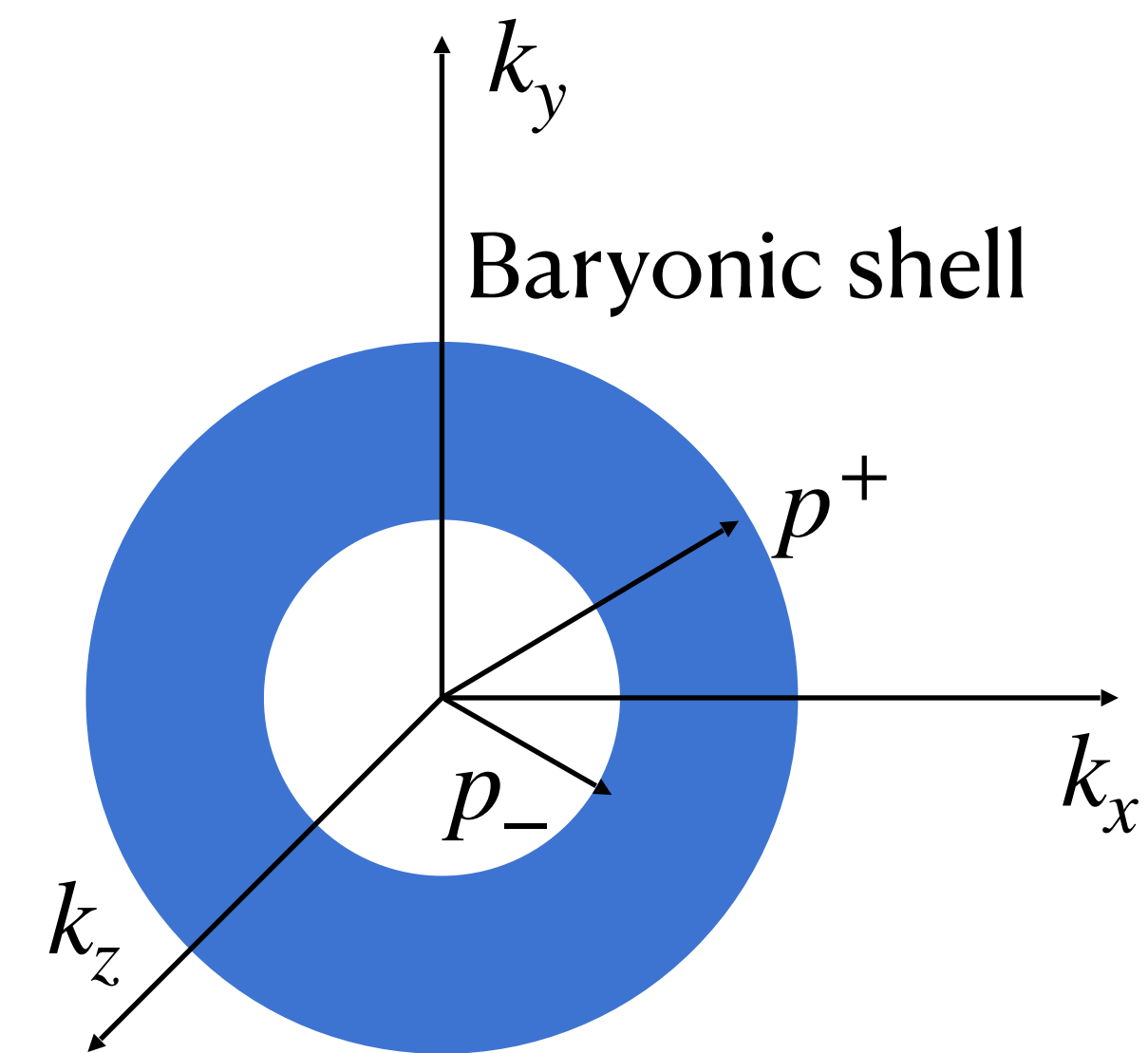
Nucleons are degenerate with quarks (quark-hadron duality)

- Perturbative quarks = quarks deep inside Fermi sphere
- Baryons = triple-pair of quarks near Fermi surface

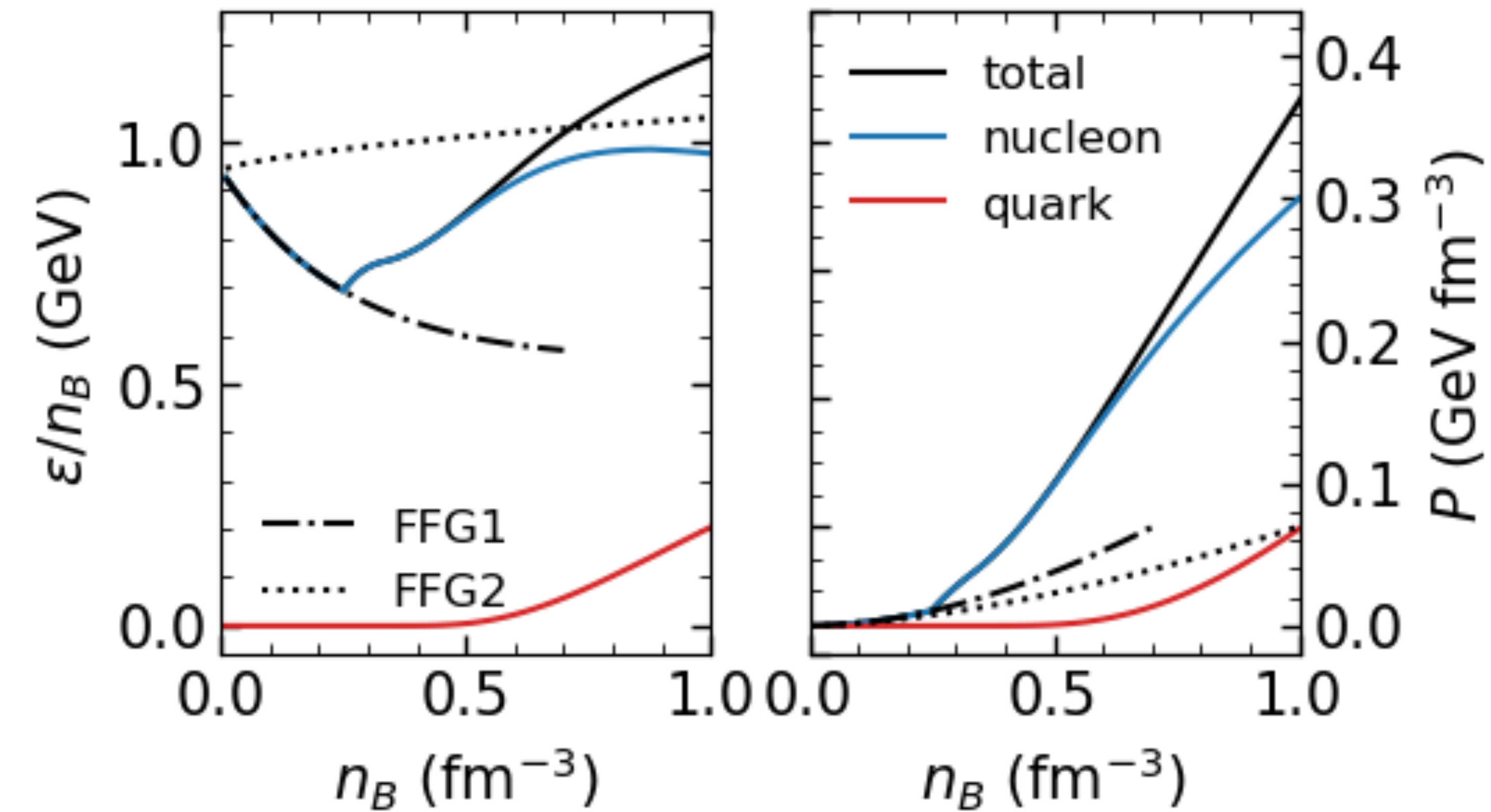




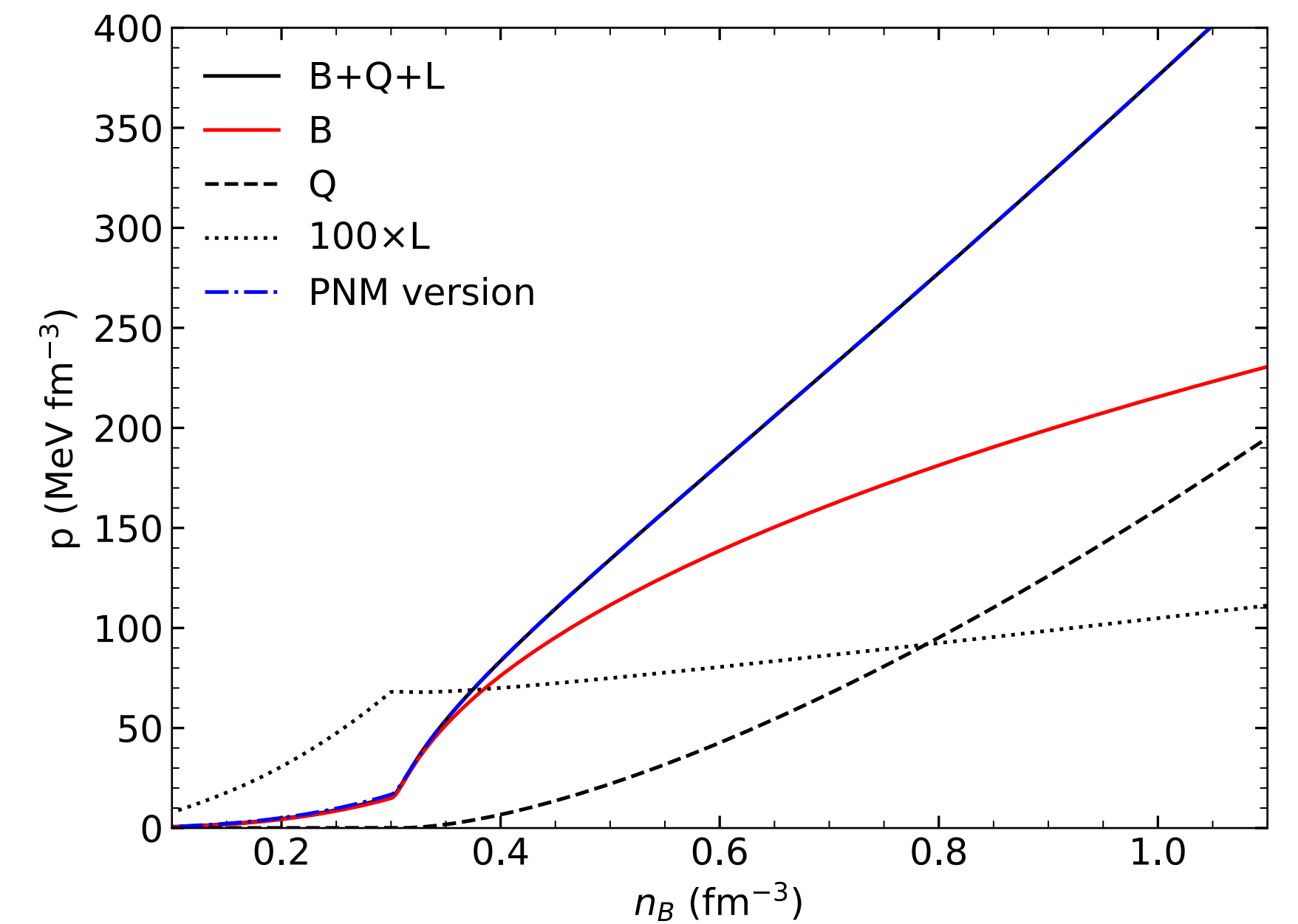
Hadronic to quarkyonic



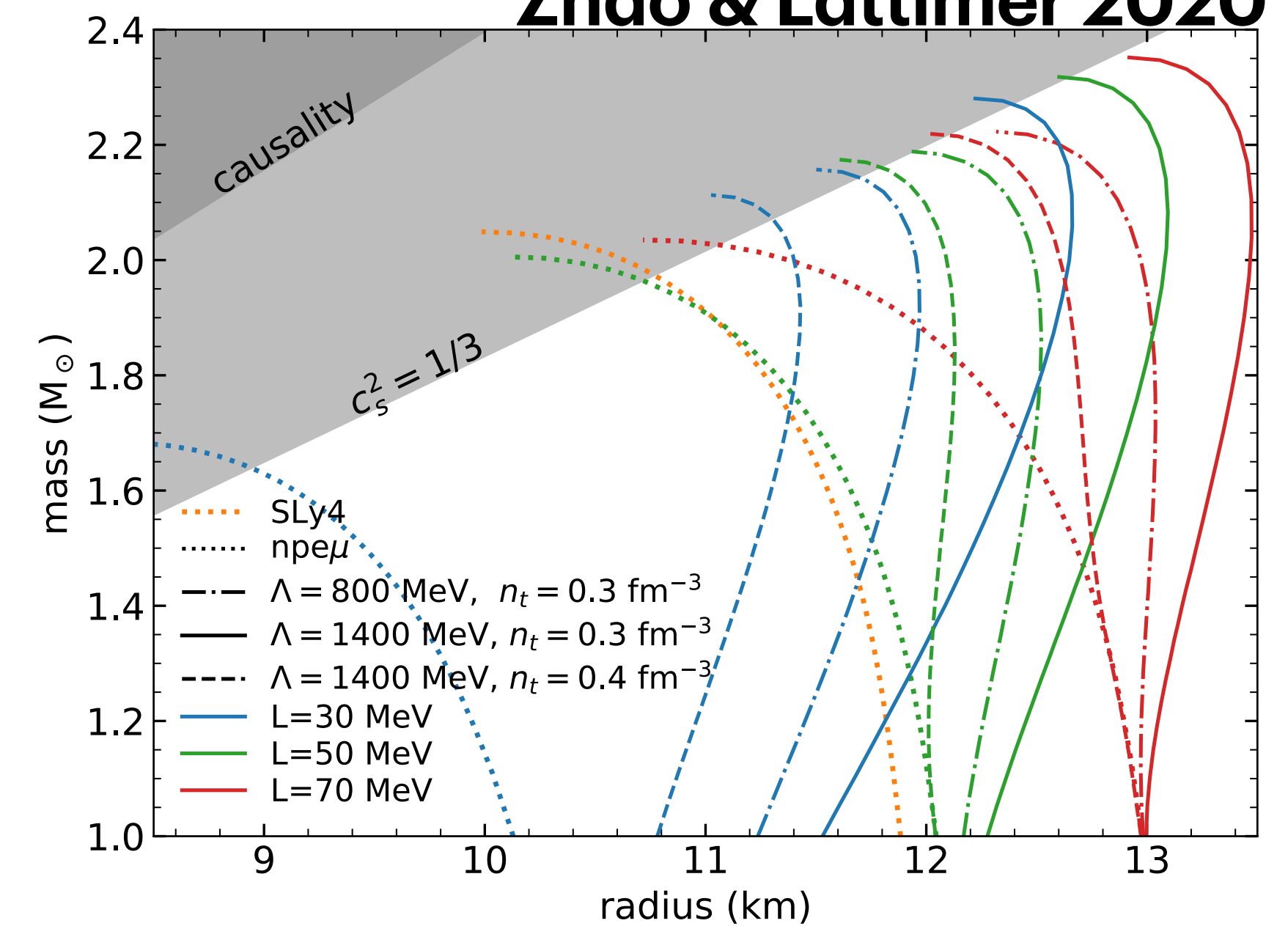
Quarkyonic EOS



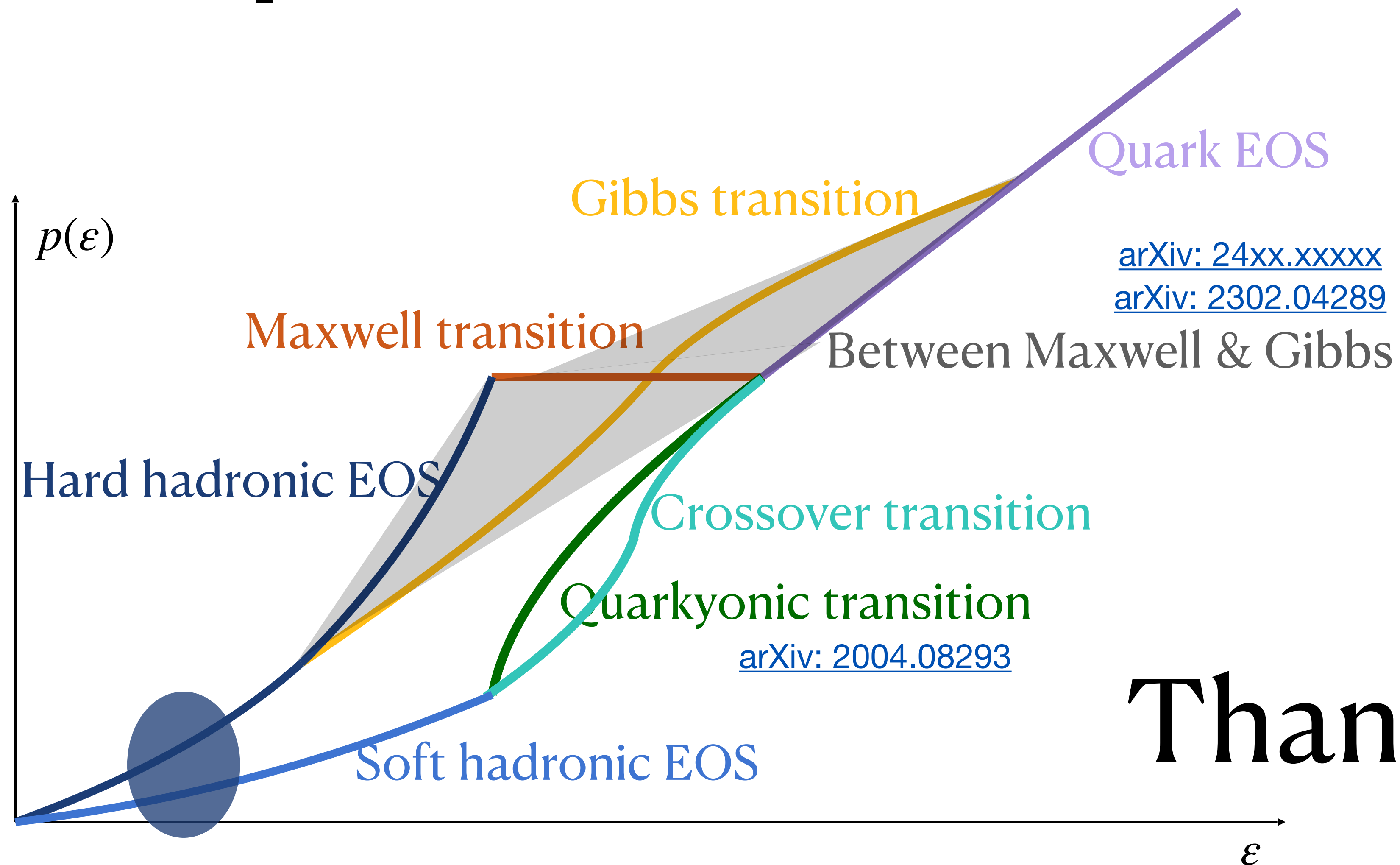
[arXiv: 24xx.xxxxx](https://arxiv.org/abs/24xx.xxxxx)



Zhao & Lattimer 2020



Hadron-quark Transition in Neutron Star Core



[arXiv: 24xx.xxxxx](#)

[arXiv: 2302.04289](#)

[arXiv: 2004.08293](#)

Thank you!

Soft hadronic EOSs is flavored by ab-initio calculation,
nuclear experiments & neutron star merger observation.

[arXiv: 2009.06441](#)

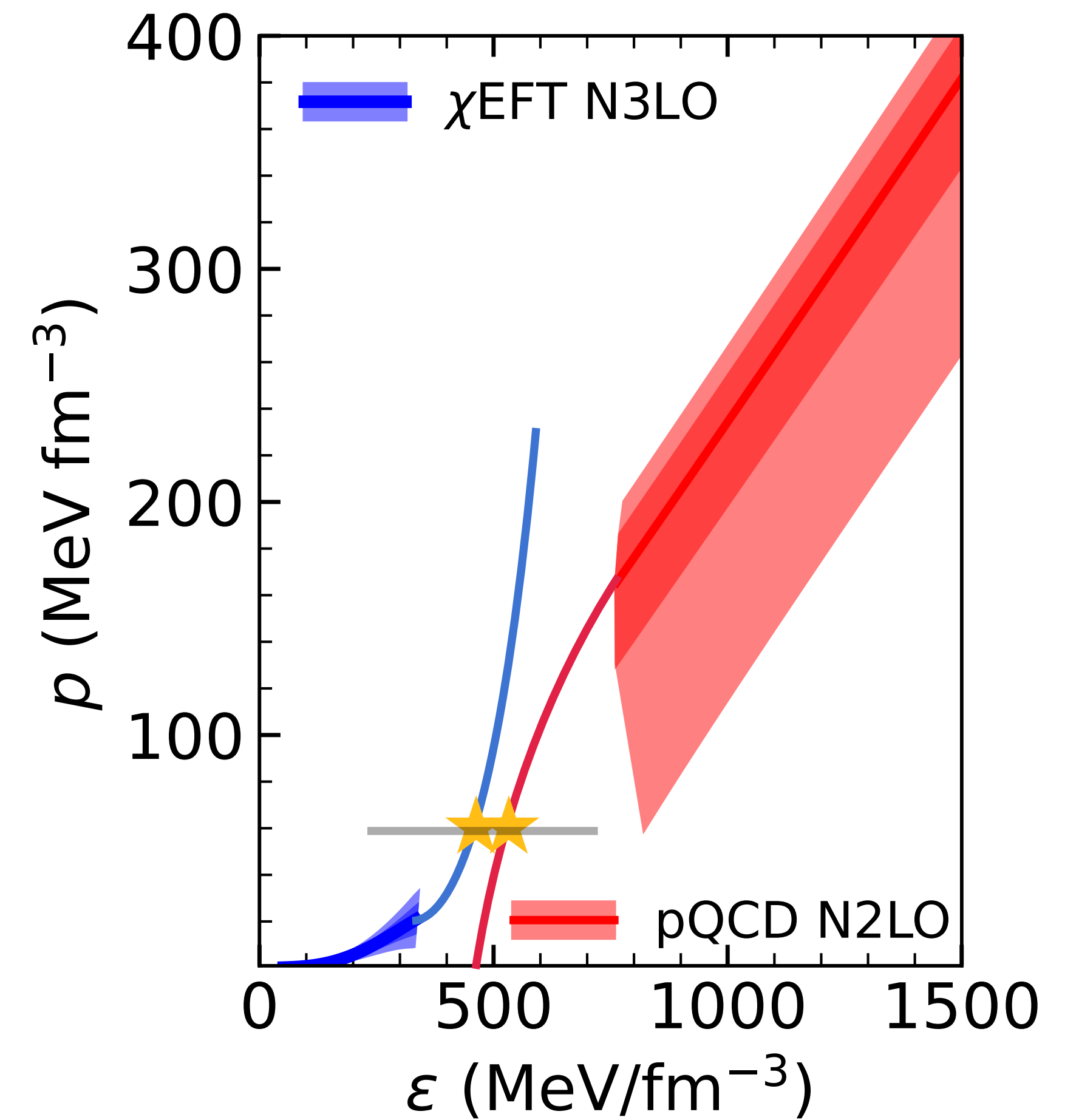
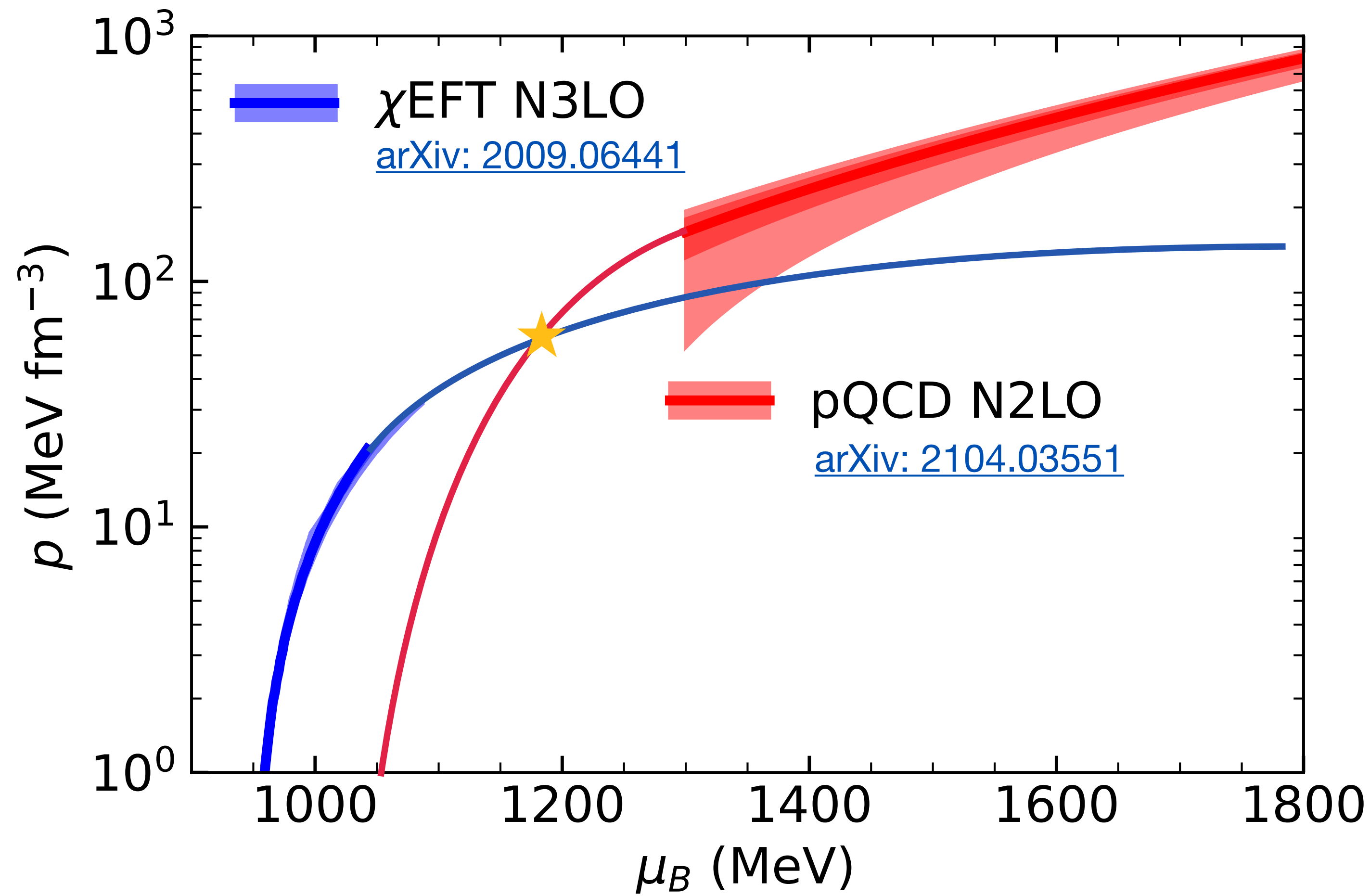
[arXiv: 2406.05267](#)

[arXiv: 1808.02858](#)

Back up slides

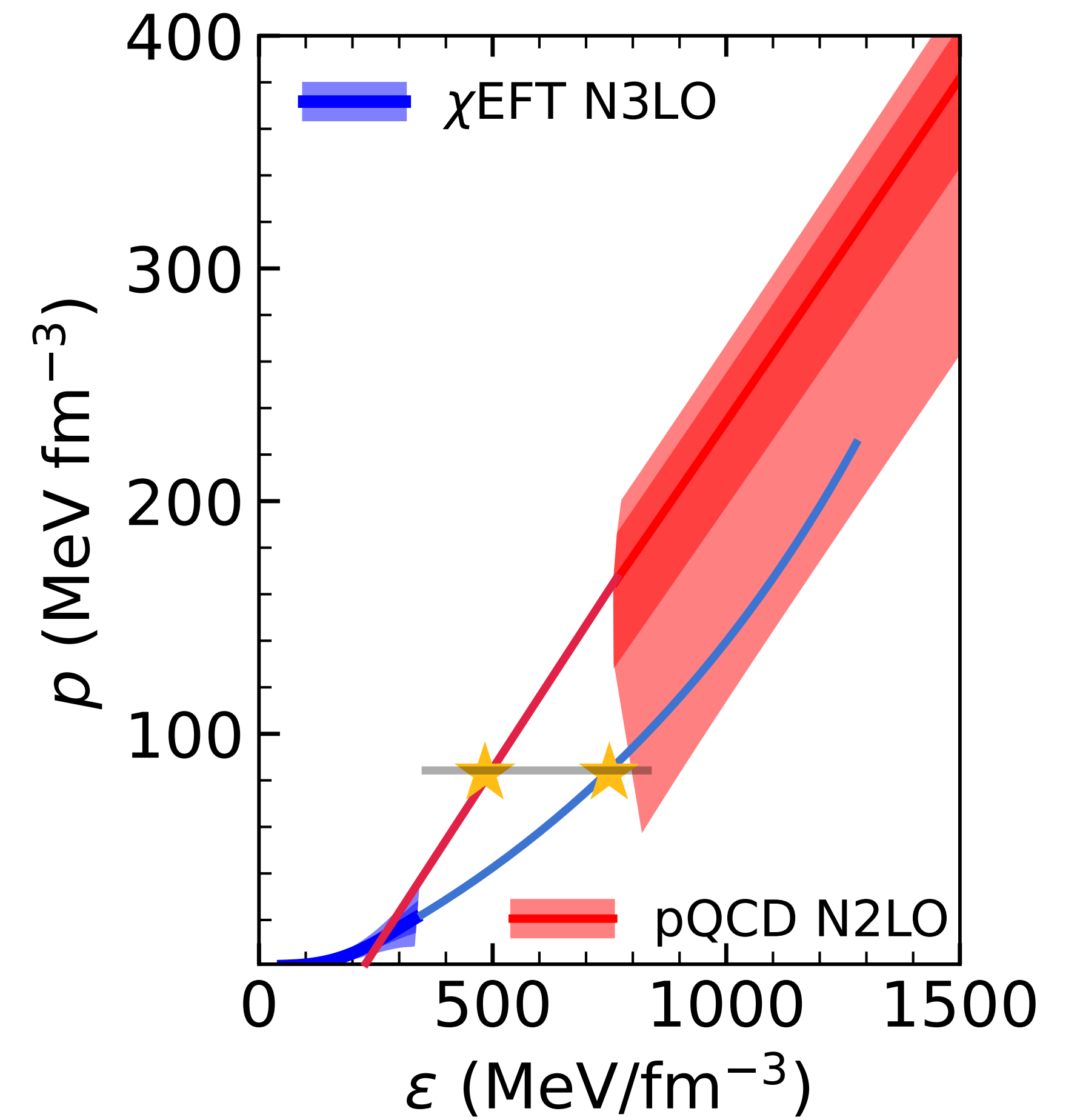
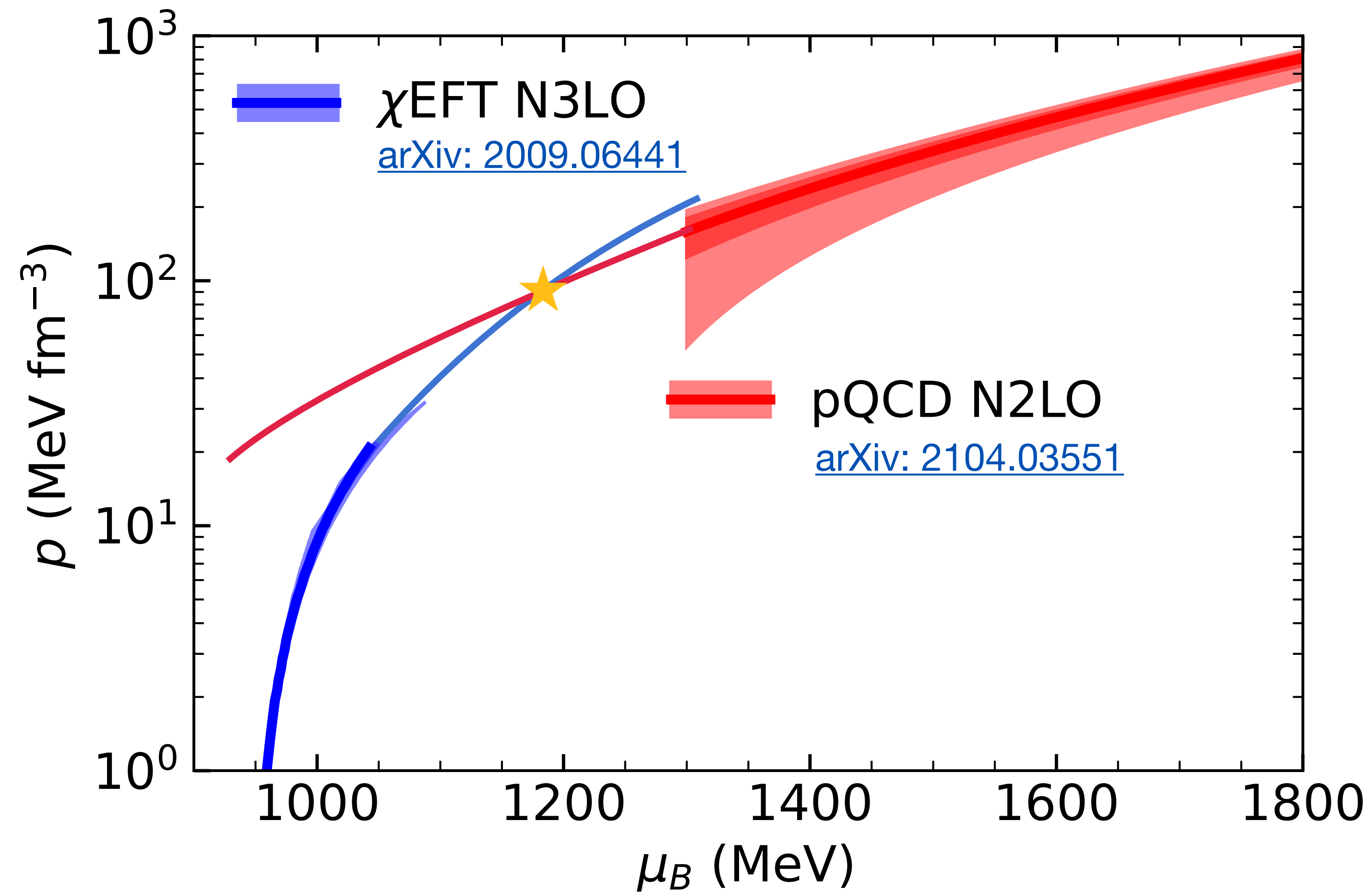
Maxwell Construction

Hybrid Neutron Stars



Maxwell Construction

Inverted Hybrid Star



Crossover Construction

Smooth interpolation

