Progress and open questions in core-collapse supernova theory

Hiroki Nagakura
(National Astronomical Observatory of Japan)

CSQCD2024 at YITP, Kyoto, Japan: October 7-11, 2024

Recent reviews of current theory of core-collapse supernova

解説

重力崩壊型超新星の物理――研究の現状と今後の課題

長倉洋樹
国立天文台
hiroki.nagakura@nao.ac.jp

山田章一
早稲田大学
shoichi@waseda.jp

重力崩壊型超新星爆発は，宇宙で起こる**大質量星**の爆発現象である．爆発を駆動している星の中心付近では，高密度（核密度）かつ高温（10 MeV 以上）環境が実現され，強い力・弱い力・電磁気力・重力という自然界で働く 4 つの基本的な力全てが爆発機構に関わっており，理論物理学の観点からも興味深い．爆発によって重元素の生成と宇宙空間への放出が起こるため，宇宙の化学組成を決める重要な天体現象である．また，爆発後には**中性子星**やブラックホールなどの高密度天体を残すことから，宇宙で起こる様々な他の高エネルギー天体現象とも密接に関連する．このように，重力崩壊型超新星爆発の研究は非常に学際的な分野であり，素・核・宇宙・天文学などの幅広い分野の研究者らによって，実験・観測・理論・シミュレーションなどの様々なアプローチにより研究が行われている．

超新星爆発を駆動している中心エンジンは，複数の物理過程が非線形に絡まった系である．その爆発機構は複雑で，理論宇宙物理学の難題の一つとして位置づけられてきた．しかし，ここ 10 年ほどの間に，超新星爆発の理論は著しく進展した．特に，理論計算（数値シミュレーション）においては，それまで爆発の再現に失敗していたのに対し，近年ではこれに成功するモデルが多く報告されている．こうした進展の一つの理由は，計算機能力の向上と数値計算手法の発展のおかげで，より正確に詳細な物理過程を取り込んだ多次元**ニュートリノ輻射流体計算**が実行可能になったことである．例えば，第一原理計算に最も近いとされる，ボルツマン方程式を直接解く多次元輻射流体計算が「富岳」などのスパコンで少数のモデルに対して実行されている一方，近似的なニュートリノ輸送法を用いた多次元計算がより多くのモデルに対して系統的に行われている．また，ニュートリノと物質との弱い相互作用の扱いについても精密化が進み，例えば核子の weak current における形状因子やストレンジネスの寄与，さらには多体効果なども，シミュレーションでは既に取り込まれている．

シミュレーションが，長時間かつ様々なタイプの大質量星に対して系統的に行えるようになり，観測量の定量的な推定が行えるようになってきたことも，近年の重要な進歩である．実際，過去の超新星理論モデルとは違い，最終的な爆発エネルギーの値や形成される中性子星の質量や半径などが定量的に議論できるようになってきた．電磁波・重力波・ニュートリノに関する理論モデルの精度も格段に上がり，**マルチメッセンジャー天文学**の発展にも貢献している．

このように超新星爆発の研究は，近年著しく発展したが，それでも超新星爆発機構が完全に解明されたわけではない．実際，現在考慮されているニュートリノ反応の取り扱いには不定性が大きく，それが爆発可否に影響する可能性がある．また，**ニュートリノ集団振動**に代表される量子運動論的な効果は，現在の最も進んだ超新星爆発計算にも取り込まれておらず，現在の超新星爆発の理論を一変させてしまうかもしれない．ニュートリノ反応計算の精度を上げ，量子運動論的ニュートリノ輻射輸送計算に基づいた超新星モデルの再構築が，今後 10 年の超新星爆発の理論的研究の主要なターゲットになるだろう．

――用語解説――

大質量星：
恒星は核融合反応によって光り輝いている星．大質量星は太陽のおよそ 10 倍以上の質量を持つ恒星を指す．

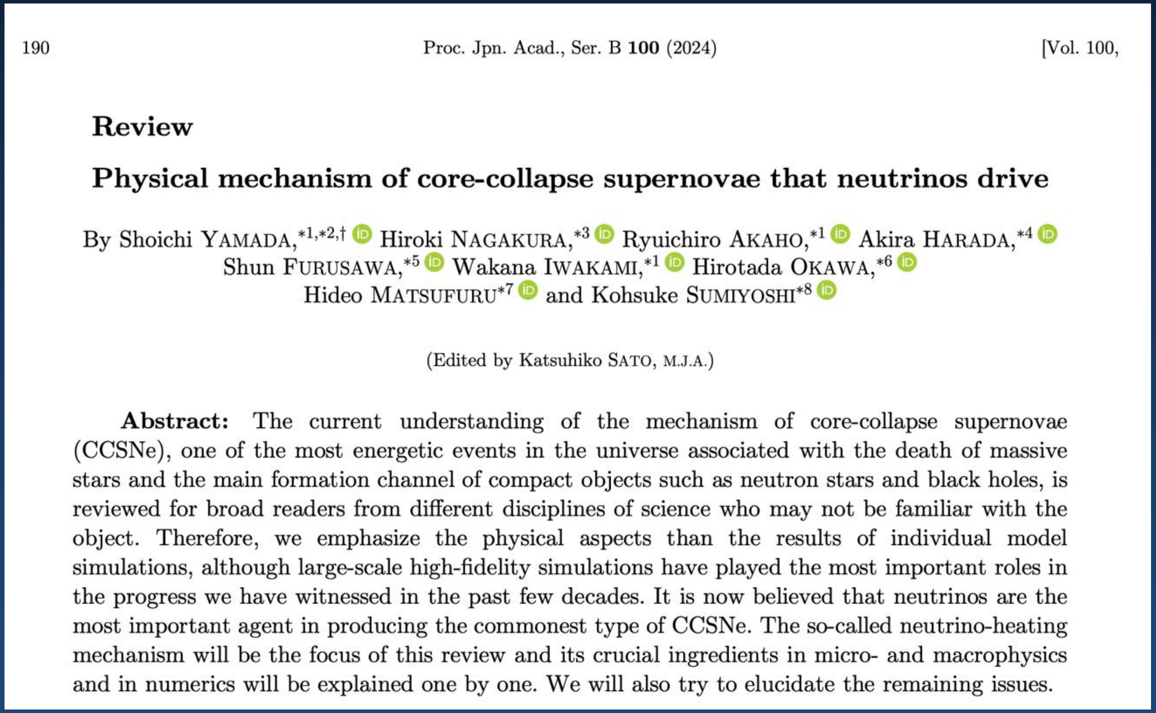
中性子星：
主に中性子から構成されている半径 10 km ほどの星．質量は太陽よりも重く，現在のところ，その 2 倍程度の質量を持つものの存在が明らかになっている．

ニュートリノ輻射流体シミュレーション：
ニュートリノと物質の相互作用を考慮し，物質の流体力学的運動と，ニュートリノ輸送を同時に解く数値計算．以下に，空間 3 次元超新星爆発シミュレーション結果の例を載せる（岩上わかな・長倉洋樹によって作成）．

マルチメッセンジャー天文学：
ある天体から発せられる様々なシグナル（電磁波，ニュートリノ・重力波・宇宙線など）を同時期に観測し，これら複数の観測量を協調させて，天体現象の起源を探ること．

ニュートリノ集団振動：
ニュートリノの自己相互作用によって駆動されるニュートリノ振動（フレーバー混合）現象．

物理学会誌（第1号 2023年）



190 Proc. Jpn. Acad., Ser. B **100** (2024) [Vol. 100,

Review

Physical mechanism of core-collapse supernovae that neutrinos drive

By Shoichi YAMADA,[*1,*2,†] Hiroki NAGAKURA,[*3] Ryuichiro AKAHO,[*1] Akira HARADA,[*4] Shun FURUSAWA,[*5] Wakana IWAKAMI,[*1] Hirotada OKAWA,[*6] Hideo MATSUFURU[*7] and Kohsuke SUMIYOSHI[*8]

(Edited by Katsuhiko SATO, M.J.A.)

Abstract: The current understanding of the mechanism of core-collapse supernovae (CCSNe), one of the most energetic events in the universe associated with the death of massive stars and the main formation channel of compact objects such as neutron stars and black holes, is reviewed for broad readers from different disciplines of science who may not be familiar with the object. Therefore, we emphasize the physical aspects than the results of individual model simulations, although large-scale high-fidelity simulations have played the most important roles in the progress we have witnessed in the past few decades. It is now believed that neutrinos are the most important agent in producing the commonest type of CCSNe. The so-called neutrino-heating mechanism will be the focus of this review and its crucial ingredients in micro- and macrophysics and in numerics will be explained one by one. We will also try to elucidate the remaining issues.

Proceedings of the Japan Academy, Ser. B, Physical and Biological Sciences

Future

A Chronological table: progress of SN (and NS) research

2015-: Multi-dimensional CCSN models with high-fidelity of input physics
Successful CCSN explosions on big iron → Connecting observations

2019-: Diversity
(SL-SNe, FBOT etc..)

2001-: Establishing 1D-Boltzmann CCSN models
(Liebendörfer et al., Sumiyoshi et al.....)

2015: Dawn of GW-astronomy

1990-: Recognizing importance of fluid instabilities
on CCSNe (Mezacappa, Janka, and Burrows......)

2010: Discovery of 2 Msun NS

1998: GRB-CCSN connection

1985-: Bruenn edited "Core" of CCSN theory

1987: IMB, Kamiokande-II made the first direct
detections of supernova neutrinos

1985: Neutrino-heating explosion
was proposed by Bethe and Wilson

1974: Observation of Hulse-Taylor binary

1966: Colgate and White
Neutrino emission from
stellar implosion

1967-: Discovery of the first radio pulsar (Hewish et al. and Gold)

1938-: Observations of extragalactic supernova and their remnants
(See e.g., Baade 1938)

1933-: Baade and Zwicky
Hypothesized Connection between neutron star and "super-nova"

Past

As of 2024, more than 50 thousands supernovae have been detected (CBAT)
(ZTF and Pan-STARRS observe dozens of supernovae per day!)
There are more than 50 thousands papers regarding “supernova” (ADS)

- Observational Phenomenology

V Progenitor mass: > 10 Msun

V 1 event /galaxy/ 100 yrs

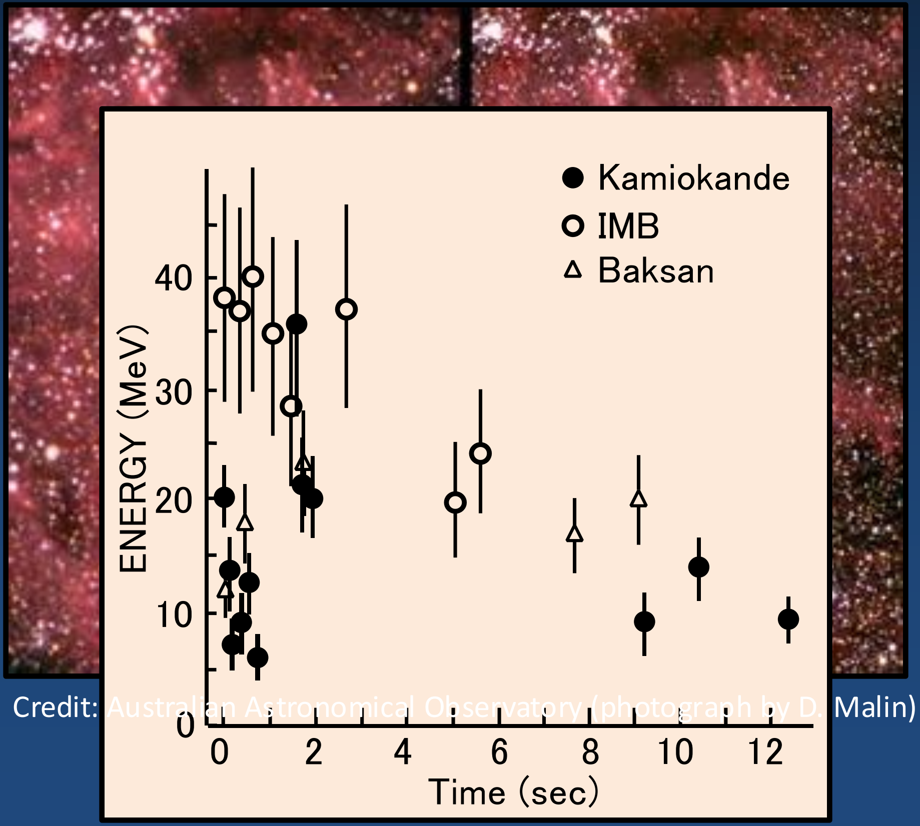
V Explosion energy: 10^{51} erg

V Nickel mass: 0.1Msun

V Neutron star remnant

V Neutrino emission: 10^{53} erg

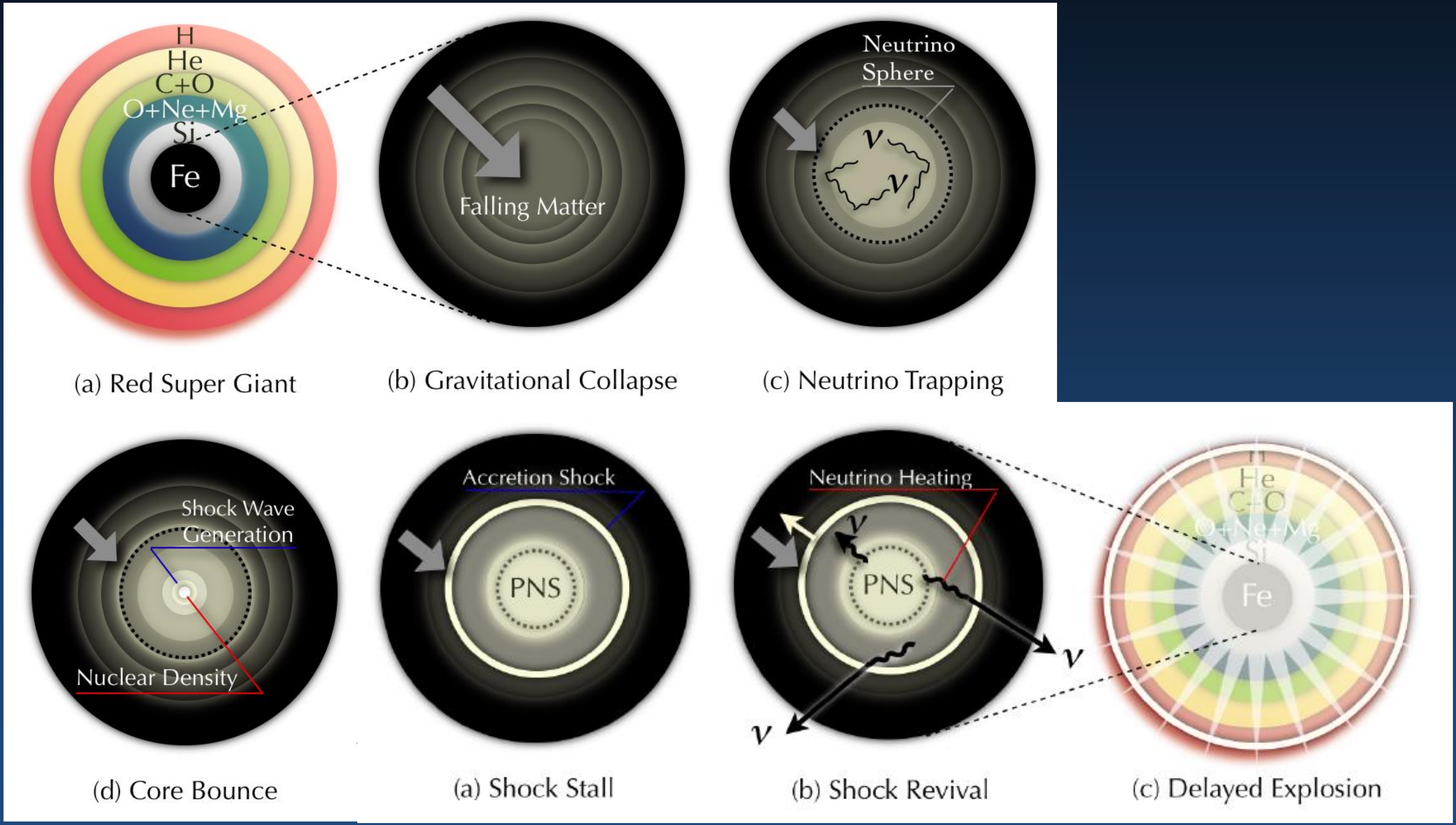
1987A: Optical image



Credit: Australian Astronomical Observatory (photograph by D. Malin)

Adapted from J-Features (M. Nakahata)

Standard Model of Core-Collapse Supernovae



(a) Red Super Giant (b) Gravitational Collapse (c) Neutrino Trapping

(d) Core Bounce (a) Shock Stall (b) Shock Revival (c) Delayed Explosion

Figure credit: Iwakami

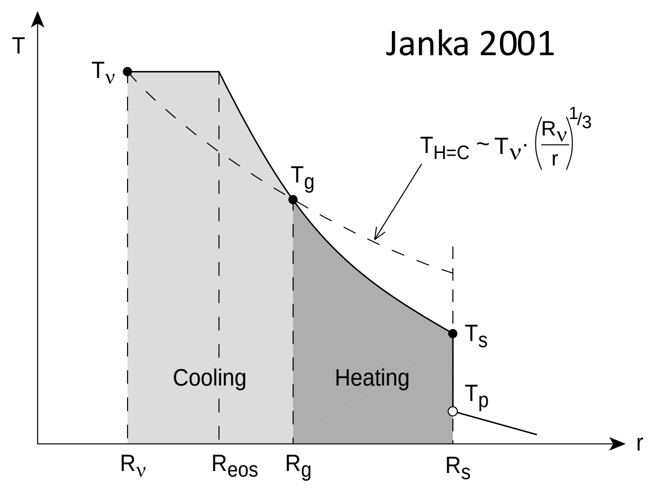
Neutrino-heating mechanism

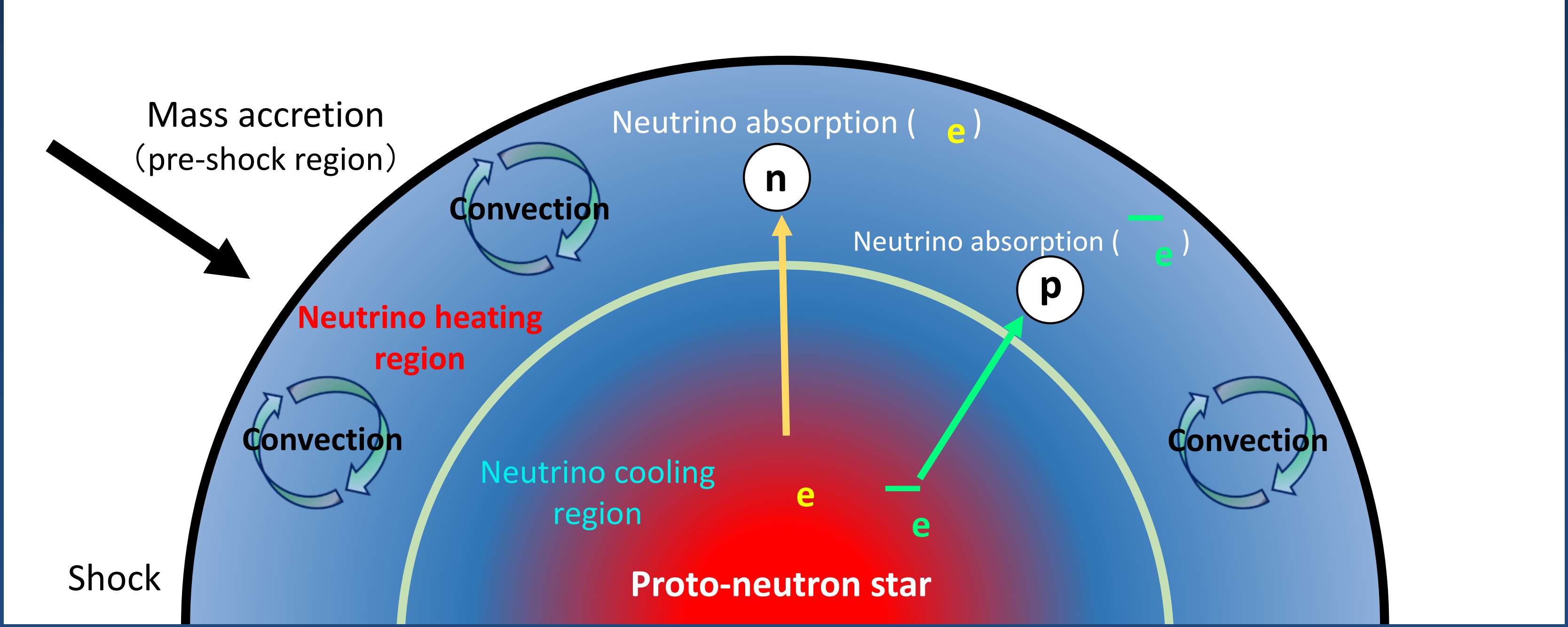
✓ Neutrino Heating Rate

$$Q_\nu^+ \approx 160\,\mathrm{Mev/s}\frac{\rho}{m_a}\frac{L_{\nu_e,52}}{r_7^2\langle\mu_\nu\rangle}\underline{\left(\frac{T_{\nu_e}}{4\,\mathrm{MeV}}\right)}^2$$

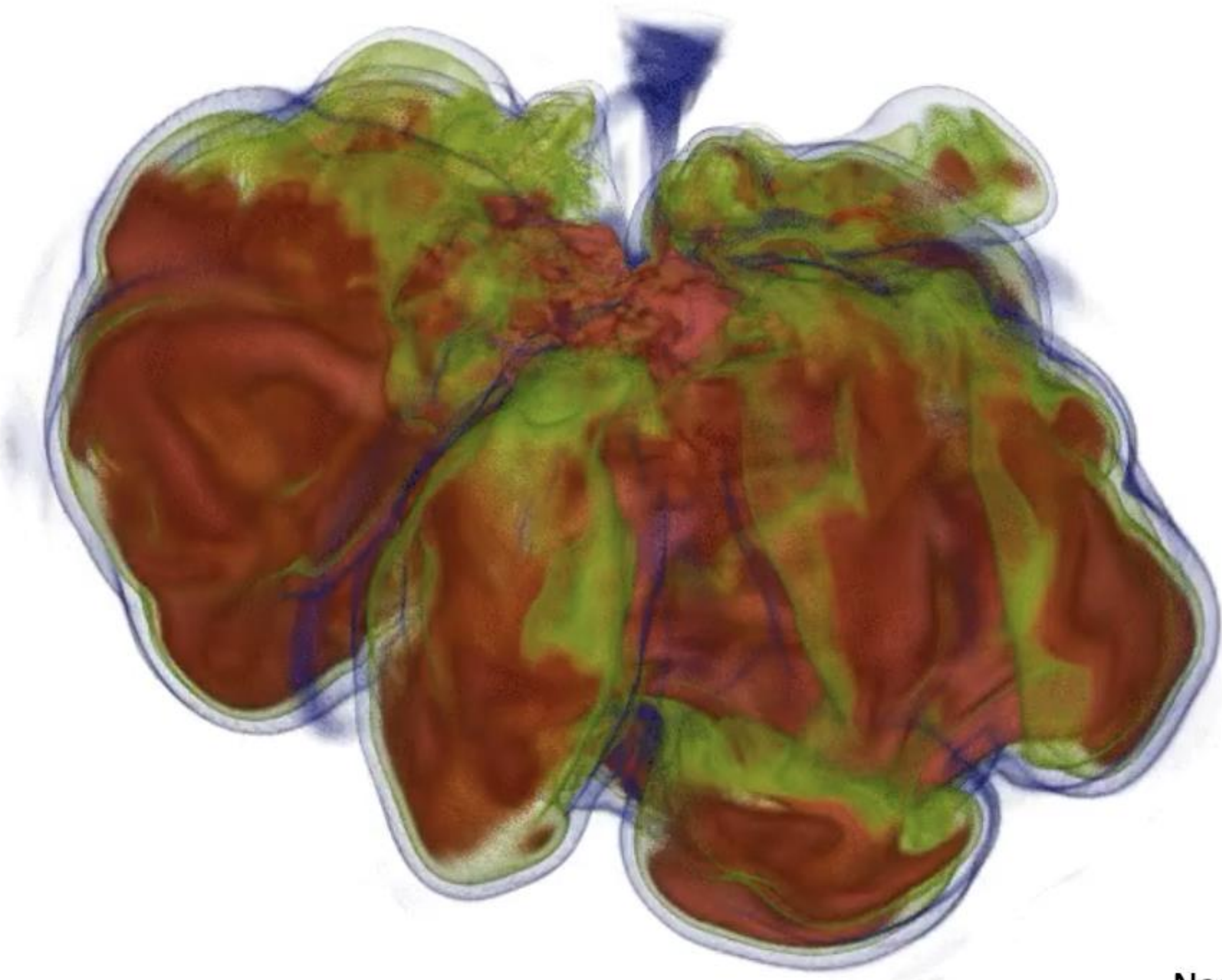
✓ Neutrino Cooling Rate

$$Q_\nu^- \approx 145\,\mathrm{Mev/s}\frac{\rho}{m_a}\underline{\left(\frac{T}{2\,\mathrm{MeV}}\right)}^6$$





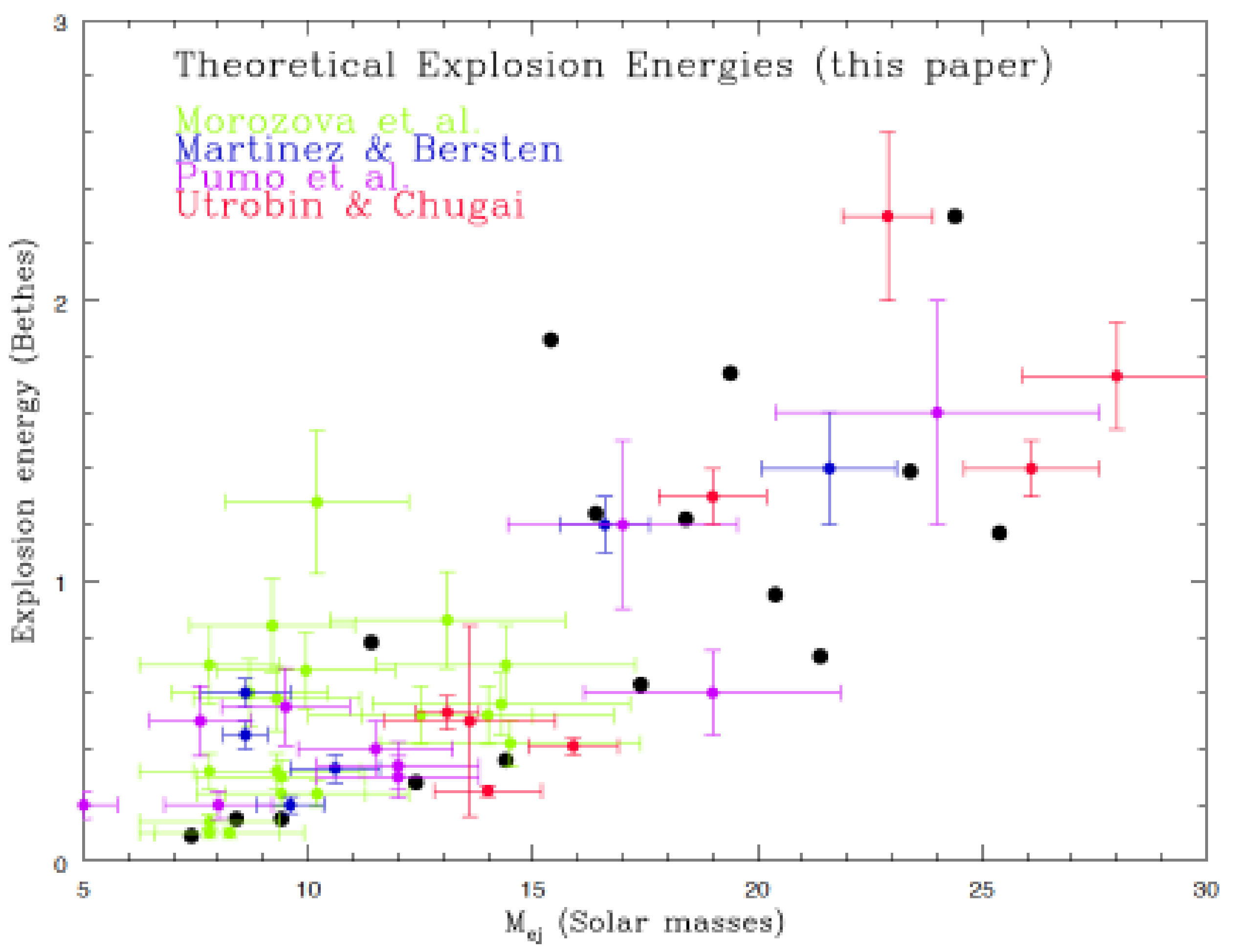
(Nowadays) CCSN explosions can be reproduced in numerical simulations



Nagakura et al. 2019

Comparison between theory (CCSN simulation) and observation

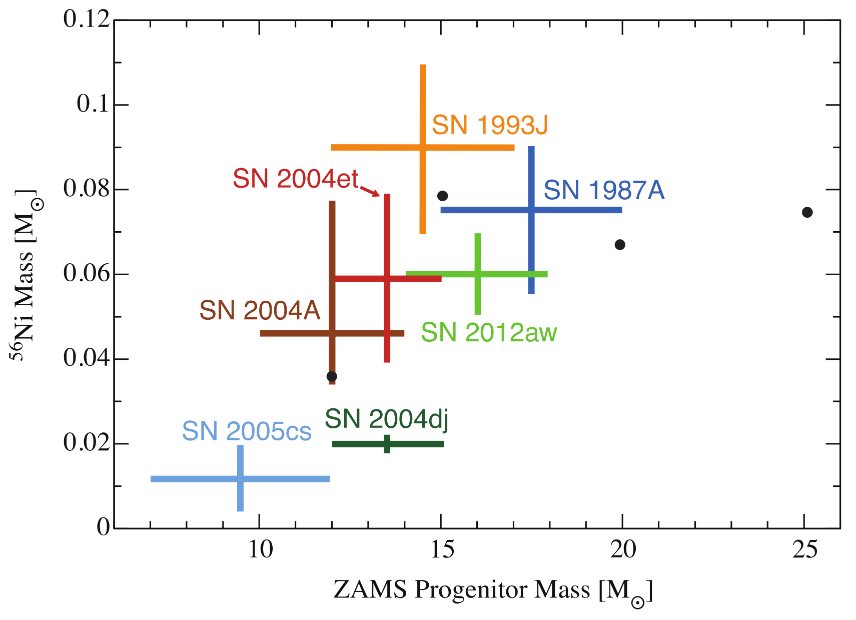
Explosion energy:



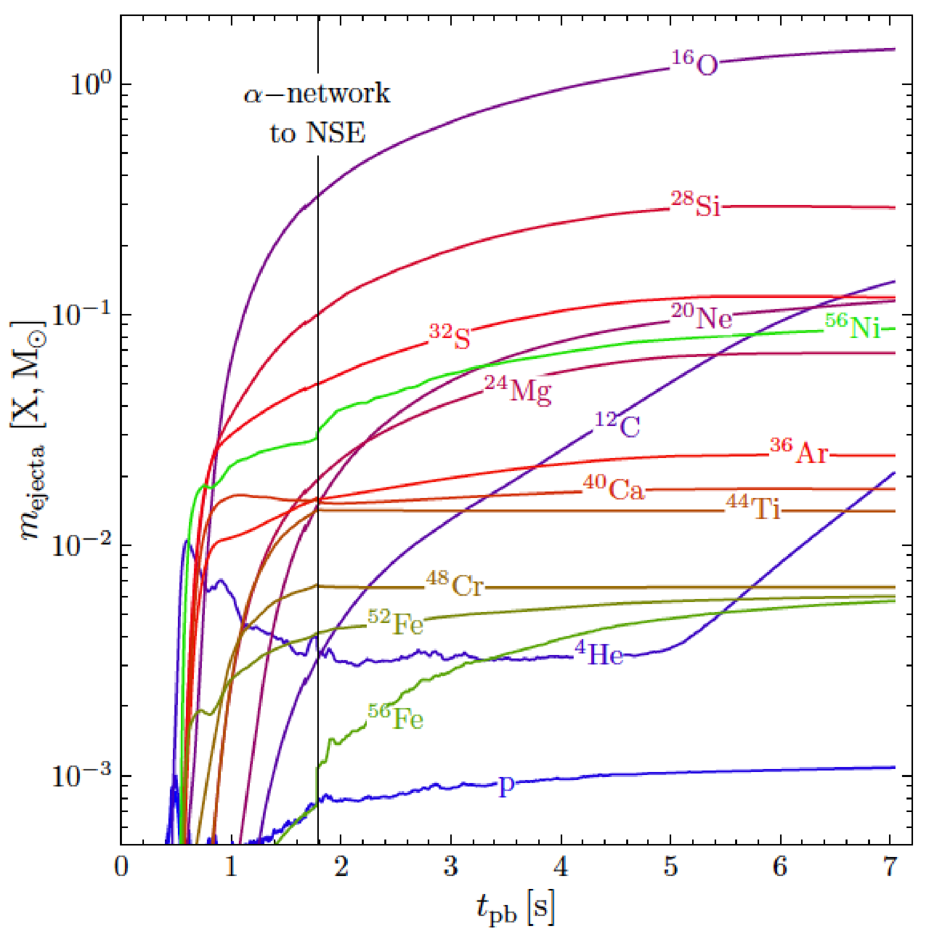
Burrows and Vartanyan 2021

Comparison between theory (CCSN simulation) and observation

Nucleosynthesis:

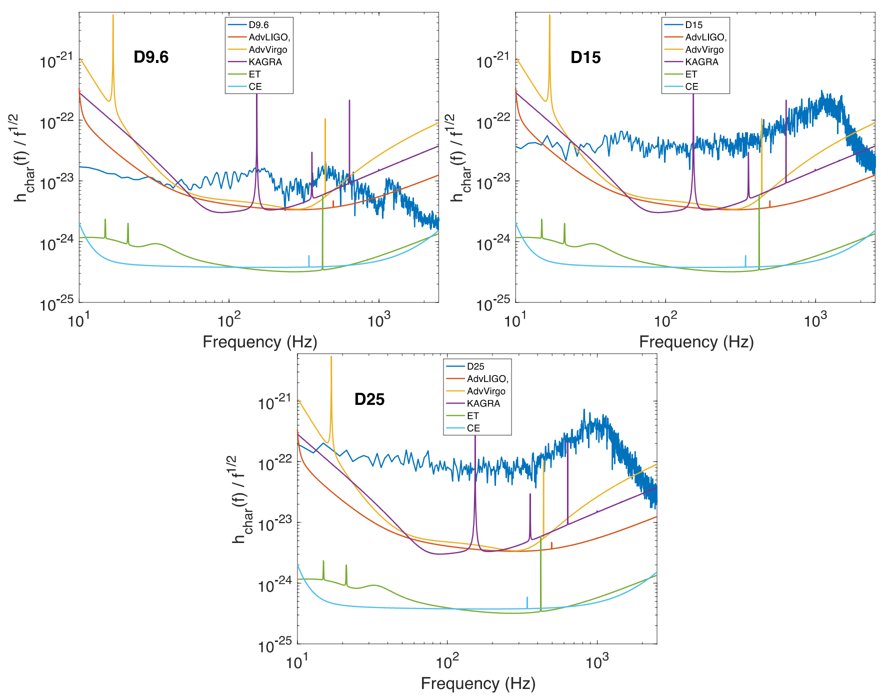


Bruenn et al. 2016

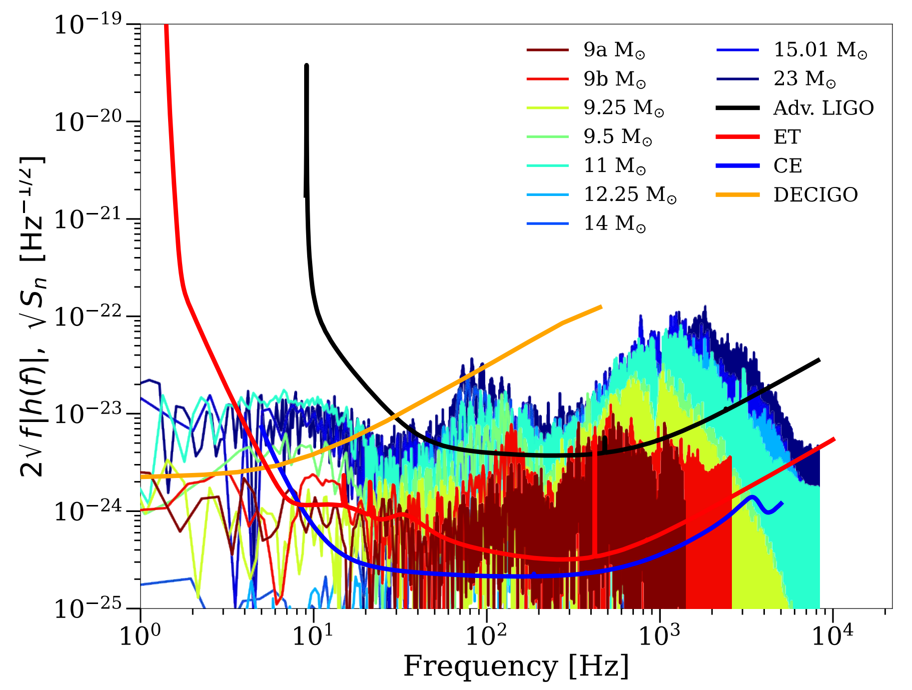


Bollig et al. 2021

Gravitational waves:



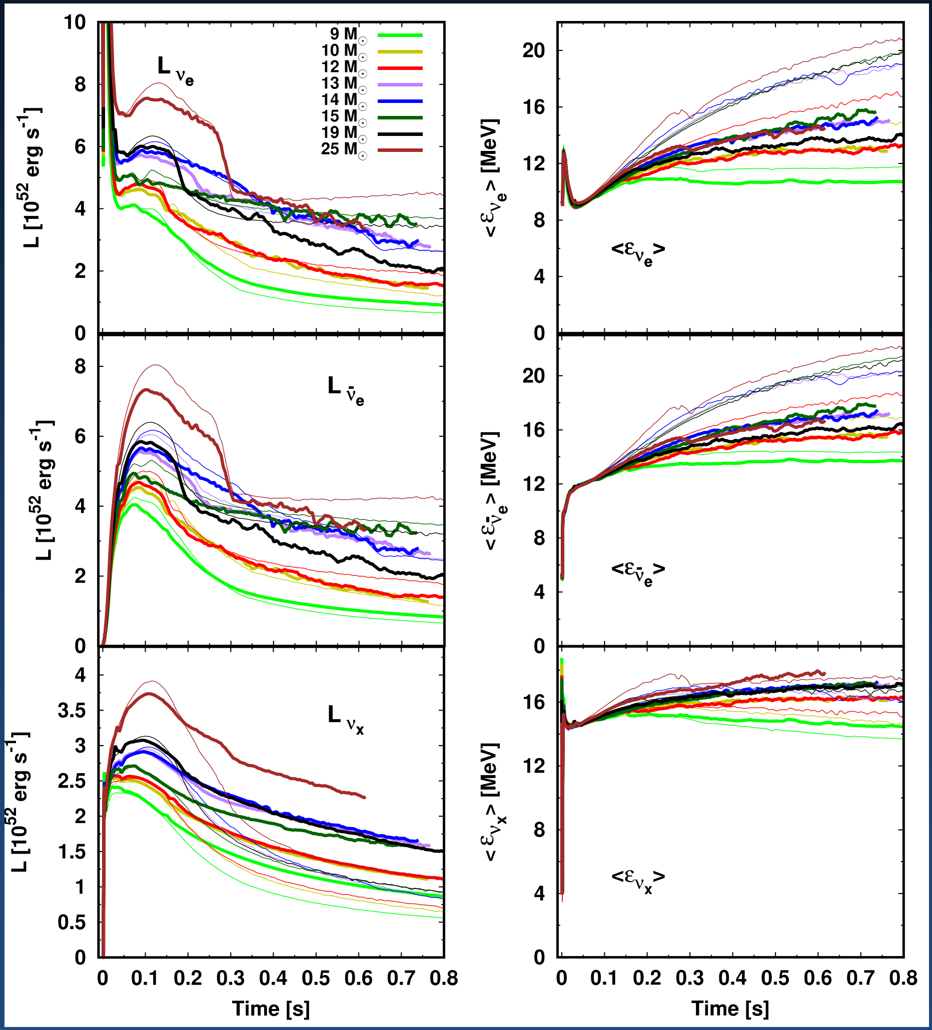
Mezzacappa et al. 2022



Vartanyan et al. 2023

→ For more detailed discussions, see Sotani's talk.

Neutrino signal:



Nagakura et al. 2021

1. Explosion models have low neutrino luminosity than those with non-explosions

(due to less accretion components)

2. The average energy of electro-type neutrinos and their anti-partners are lower in 3D than 1D.

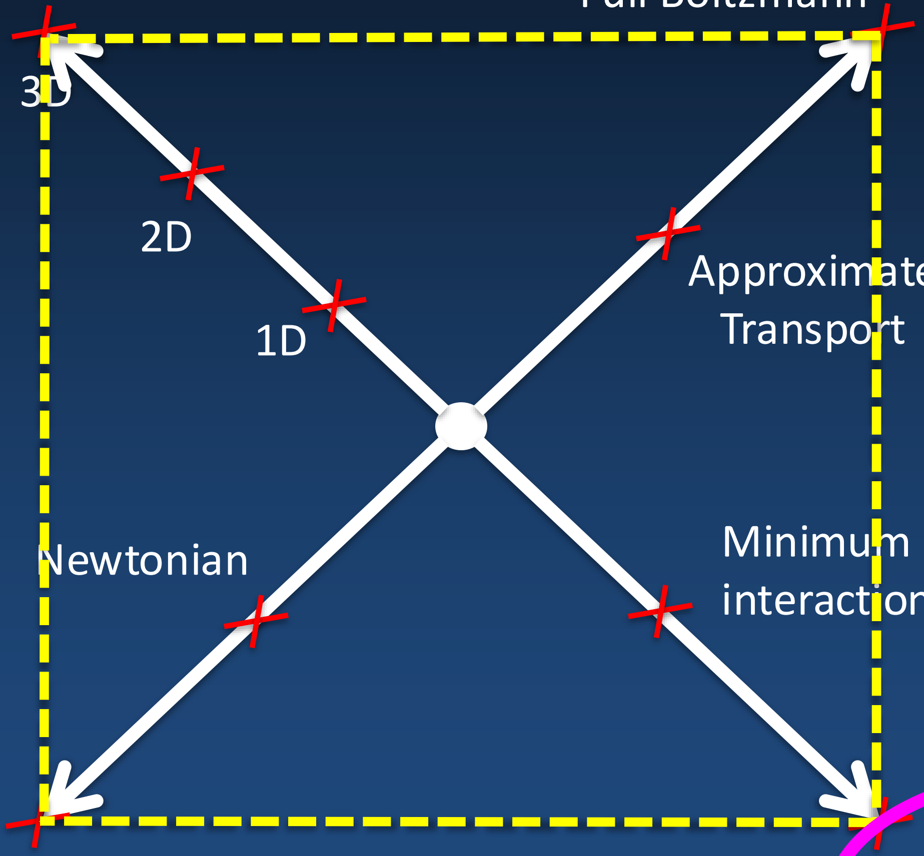
3. Neutrino luminosity of heavy-leptonic neutrinos are higher in 3D than 1D.

(due to PNS convection)

- There remain many ”holes” in CCSN simulations

Dimensionality
(for Hydro)

Neutrino
Transport



Full Boltzmann

3D

2D

1D

Approximate
Transport

Newtonian

Minimum weak
interaction sets

Full GR

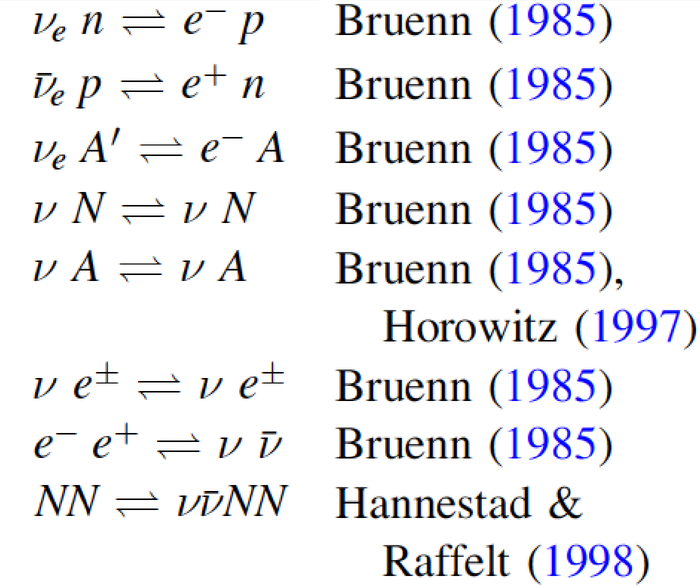
Most advanced
Weak interaction

Gravity

EOS
Weak Interactions

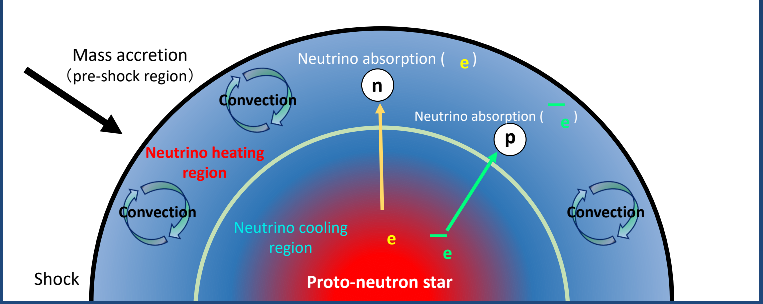
Weak Interactions

Basic Sets:



$\nu_e\, n \rightleftharpoons e^-\, p$	Bruenn (1985)
$\bar{\nu}_e\, p \rightleftharpoons e^+\, n$	Bruenn (1985)
$\nu_e\, A' \rightleftharpoons e^-\, A$	Bruenn (1985)
$\nu\, N \rightleftharpoons \nu\, N$	Bruenn (1985)
$\nu\, A \rightleftharpoons \nu\, A$	Bruenn (1985), Horowitz (1997)
$\nu\, e^\pm \rightleftharpoons \nu\, e^\pm$	Bruenn (1985)
$e^-\, e^+ \rightleftharpoons \nu\, \bar{\nu}$	Bruenn (1985)
$NN \rightleftharpoons \nu\bar{\nu}NN$	Hannestad & Raffelt (1998)

Lentz et al. 2011, Kotake et al. 2018
See also Burrows et al. 2004



Updates

Hadron Sectors (Nucleon scattering):

Nucleon Neutral Weak Current

$$J_\mu = \langle N(p')|F_1(Q^2)\gamma_\mu + \underline{F_2(Q^2)\sigma_{\mu\nu}q^\nu} + G_A(Q^2)\gamma_\mu\gamma_5|N(p)\rangle$$

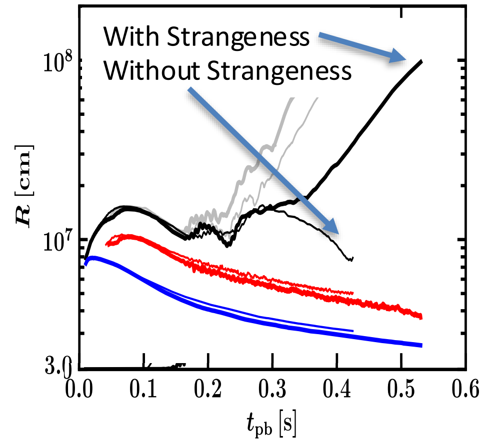
Weak magnetism

$$G_A(Q^2) = \frac{1}{2}\frac{G_A(0)}{(1+Q^2/M_A^2)}\tau_3 + \underline{G_A^s(Q^2)}$$

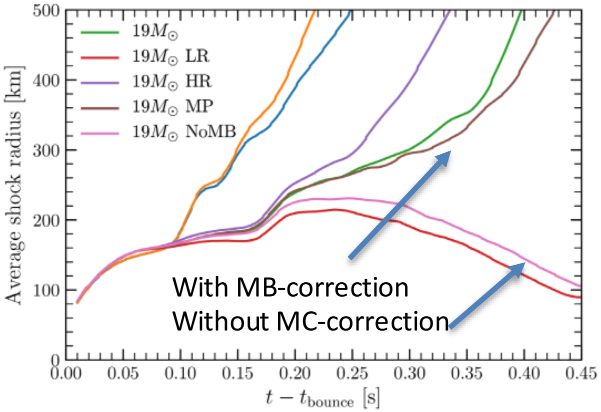
Strangeness contribution

$$\frac{1}{V}\frac{d\sigma}{d\Omega} = \frac{G_F^2E_\nu^2}{16\pi^2}\left[g_a^2(3-\cos\theta)(n_n+n_p)\underline{S_A} + (1+\cos\theta)n_n\underline{S_V}\right].$$

Many-body corrections



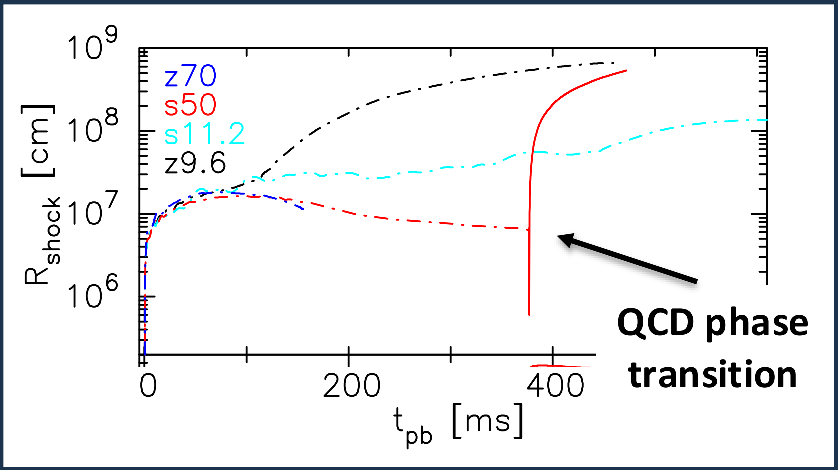
Melson et al. 2015



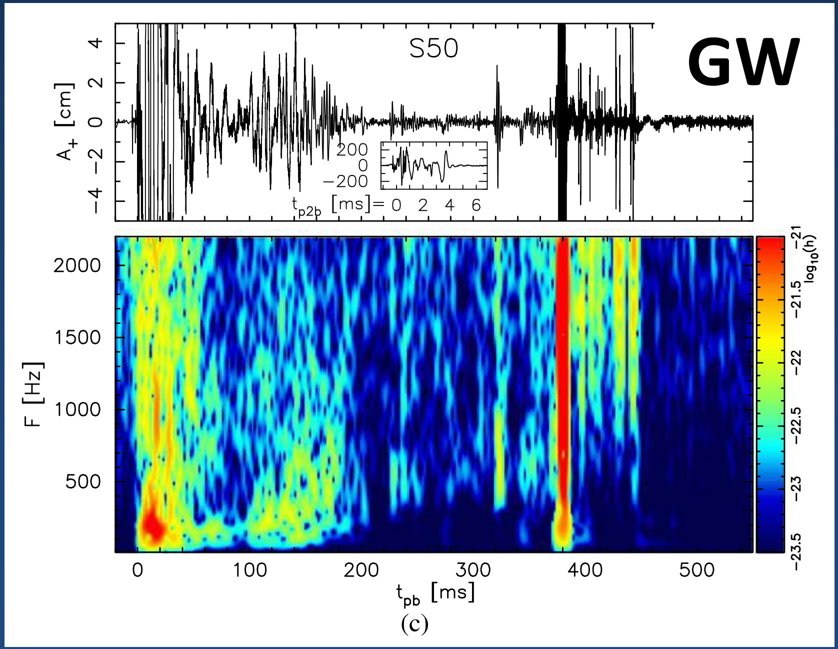
Burrows et al. 2020

CCSN driven by a first-order QCD phase transition

Kuroda et al. 2022

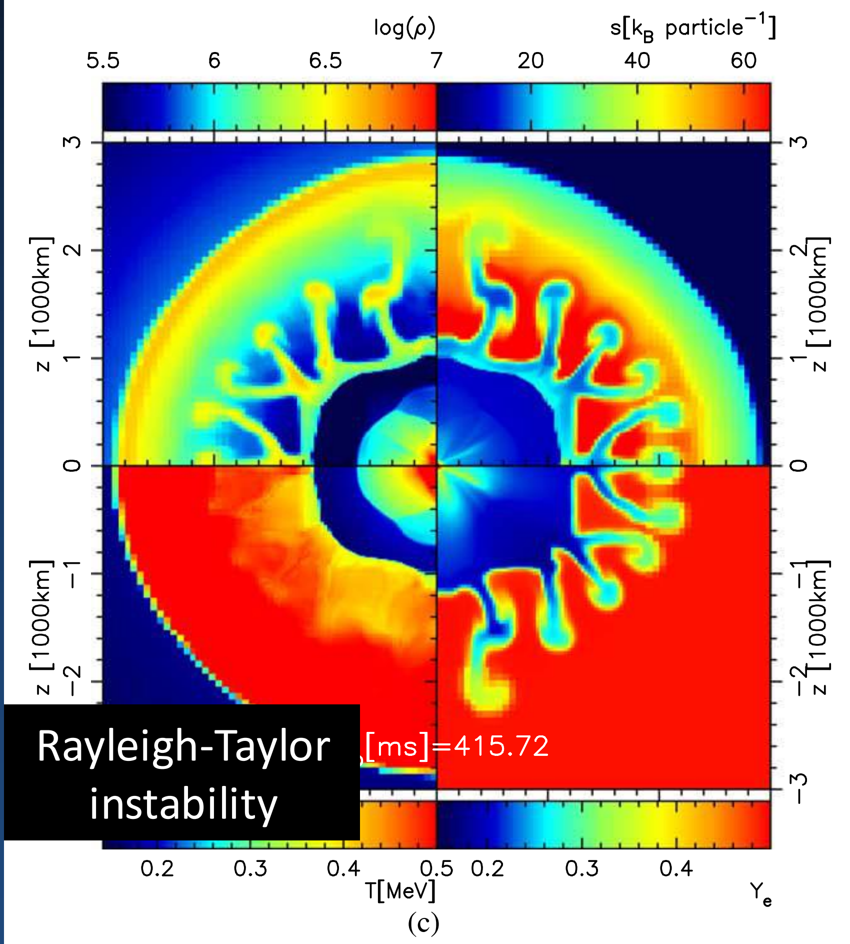


QCD phase transition



GW

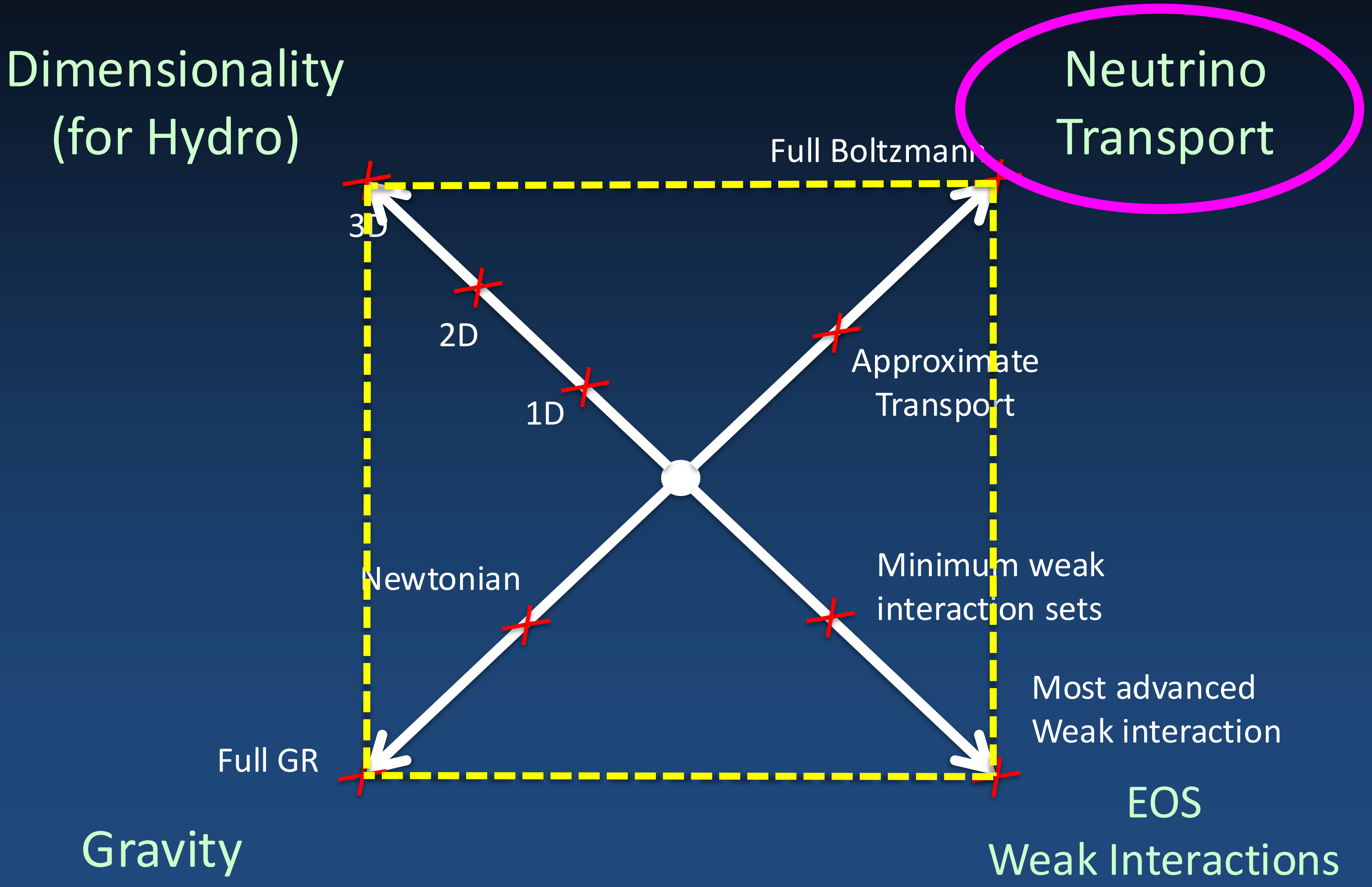
Post-phase transition phase



Rayleigh-Taylor instability

See also talks by Osakwe, Kumar, and Zhang

- There remain many ”holes” in CCSN simulations



Modeling of neutrino radiation field: necessitating a kinetic treatment

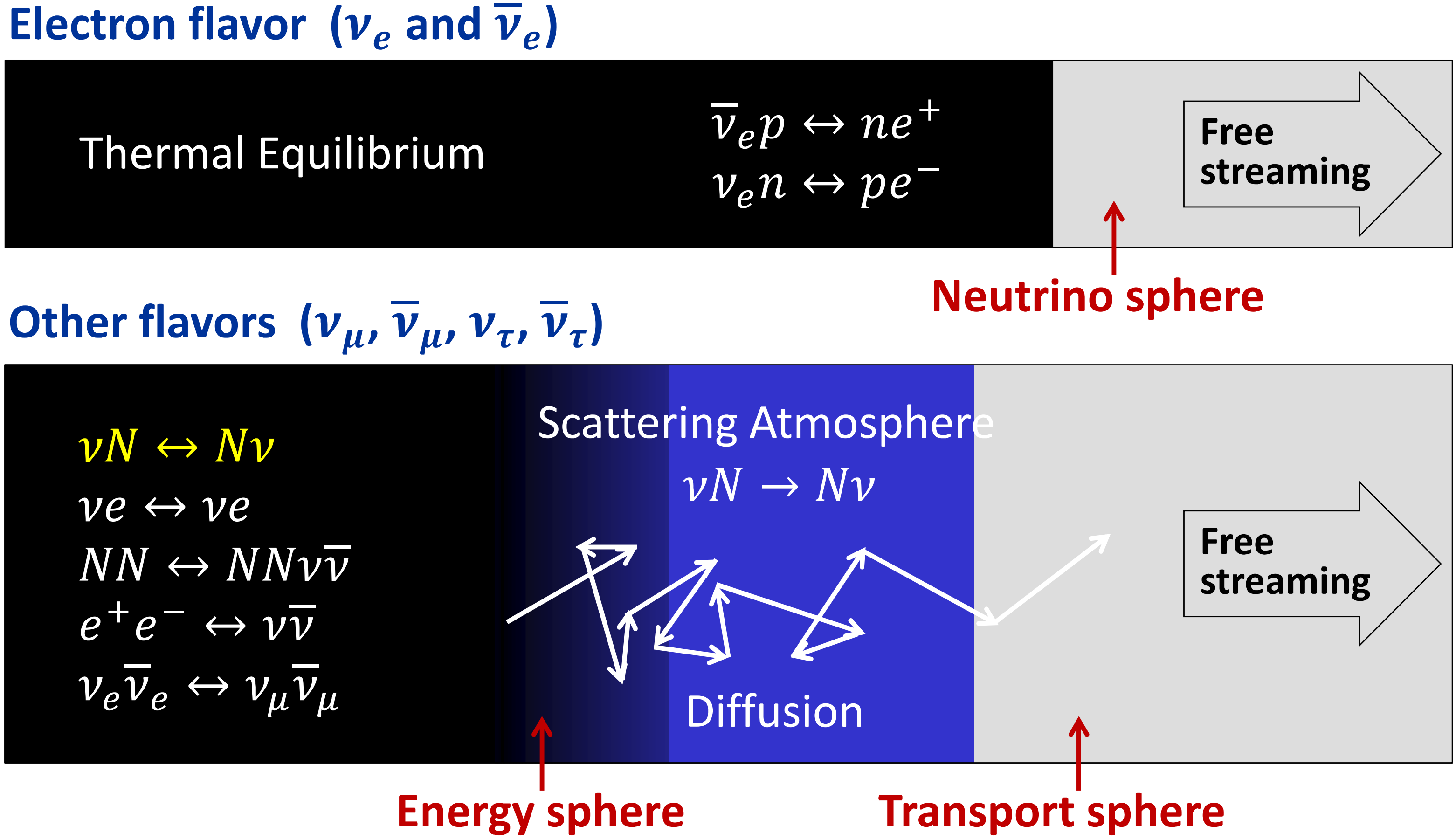


Figure by Janka 2017

Electron flavor (ν_e and $\bar{\nu}_e$)

Thermal Equilibrium

$\bar{\nu}_e p \leftrightarrow n e^+$

$\nu_e n \leftrightarrow p e^-$

Free streaming

Neutrino sphere

Other flavors ($\nu_\mu, \bar{\nu}_\mu, \nu_\tau, \bar{\nu}_\tau$)

$\nu N \leftrightarrow N \nu$

$\nu e \leftrightarrow \nu e$

$NN \leftrightarrow NN\nu\bar{\nu}$

$e^+ e^- \leftrightarrow \nu\bar{\nu}$

$\nu_e \bar{\nu}_e \leftrightarrow \nu_\mu \bar{\nu}_\mu$

Scattering Atmosphere

$\nu N \rightarrow N \nu$

Diffusion

Free streaming

Energy sphere

Transport sphere

Optically thick ↔ Optically thin

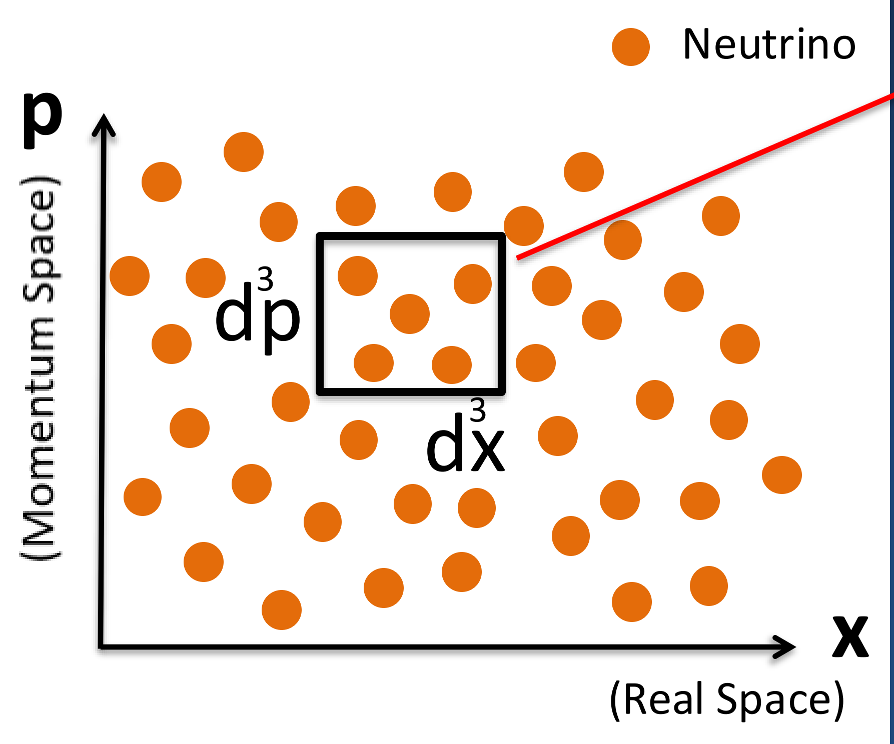
General relativistic
full Boltzmann neutrino transport

$$p^{\mu}\frac{\partial f}{\partial x^{\mu}} + \frac{dp^{i}}{d\tau}\frac{\partial f}{\partial p^{i}} = \left(\frac{\delta f}{\delta \tau}\right)_{\mathrm{col}},$$

Sumiyoshi and Yamada, 2012
Nagakura et al. 2014, 2017, 2019
Akaho et al. 2021

(Time evolution + Advection Term) (Collision Term)

6 dimensional Phase Space



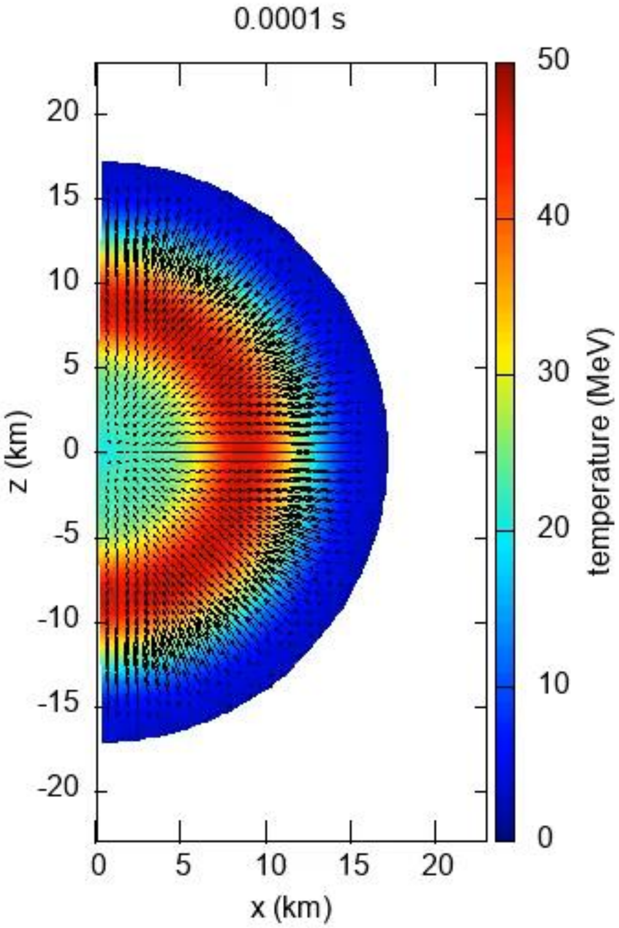
$$dN = f(t, \boldsymbol{p}, \boldsymbol{x})d^3p\,d^3x$$

Conservative form of GR Boltzmann eq.

$$\frac{1}{\sqrt{-g}}\frac{\partial(\sqrt{-g}\nu^{-1}p^{\alpha}f)}{\partial x^{\alpha}}\bigg|_{q_{(i)}} + \frac{1}{\nu^2}\frac{\partial}{\partial\nu}(-\nu f p^{\alpha}p_{\beta}\nabla_{\alpha}e^{\beta}_{(0)})$$
$$+\frac{1}{\sin\bar{\theta}}\frac{\partial}{\partial\bar{\theta}}\left(\nu^{-2}\sin\bar{\theta}f\sum_{j=1}^{3}p^{\alpha}p_{\beta}\nabla_{\alpha}e^{\beta}_{(j)}\frac{\partial\ell_{(j)}}{\partial\bar{\theta}}\right)$$
$$+\frac{1}{\sin^2\bar{\theta}}\frac{\partial}{\partial\bar{\varphi}}\left(\nu^{-2}f\sum_{j=2}^{3}p^{\alpha}p_{\beta}\nabla_{\alpha}e^{\beta}_{(j)}\frac{\partial\ell_{(j)}}{\partial\bar{\varphi}}\right) = S_{\mathrm{rad}},$$

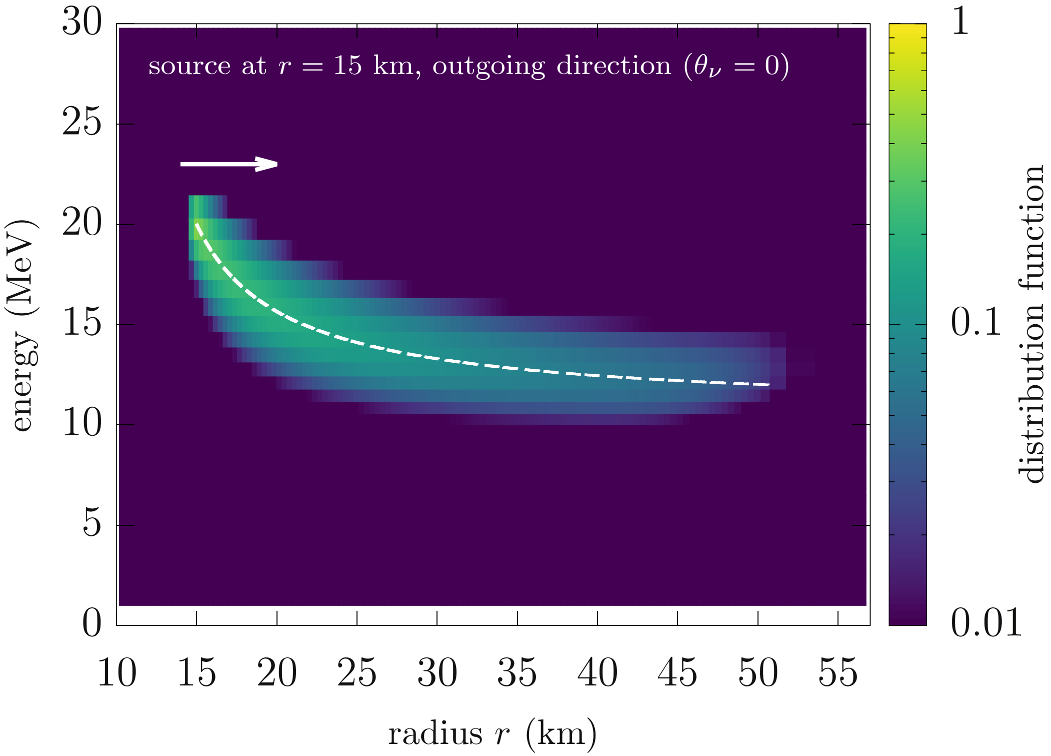
Shibata and HN et al. 2014, Cardall et al. 2013

ν GR simulations with full Boltzmann neutrino transport



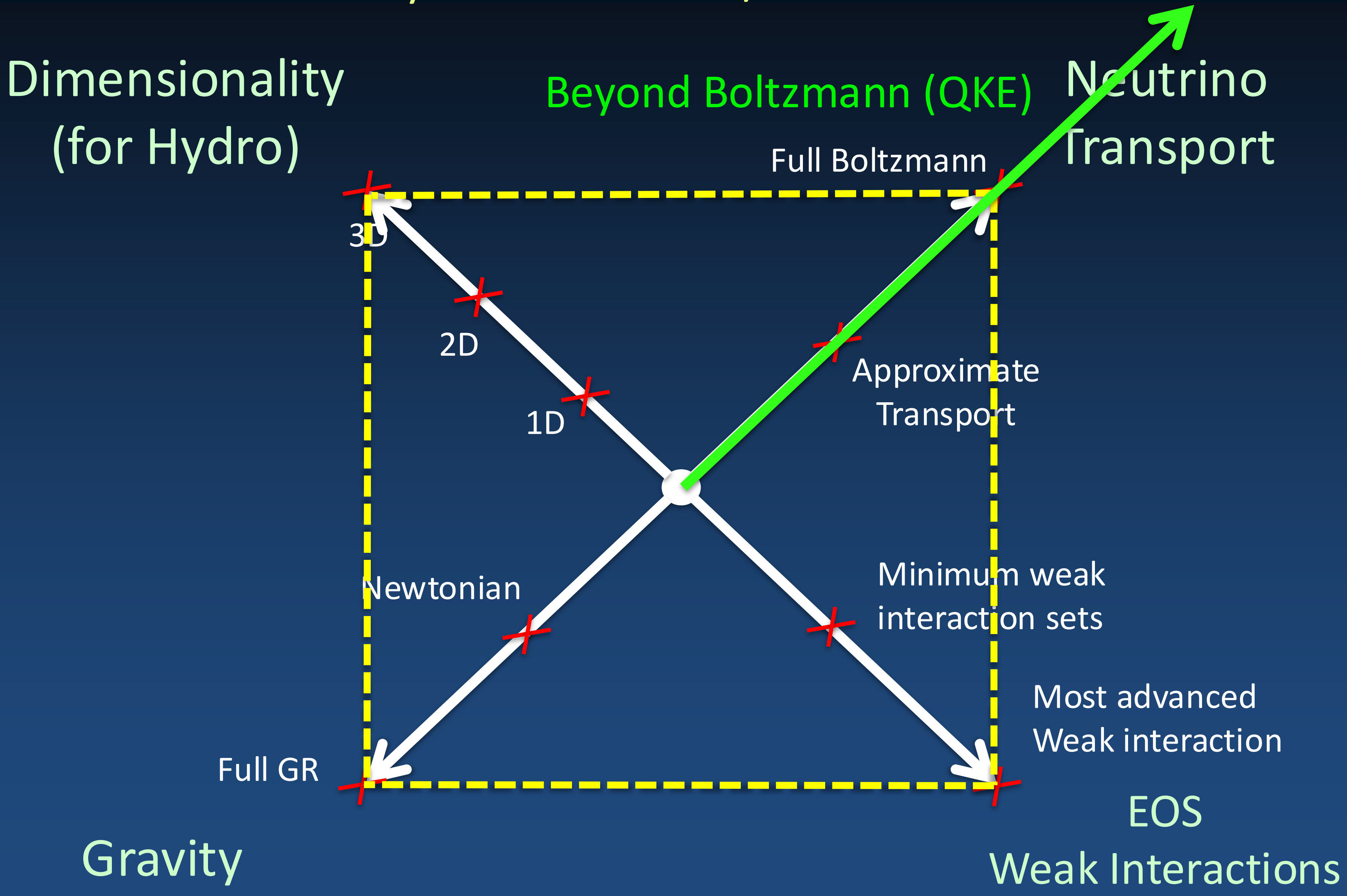
Akaho et al. 2023

Gravitational redshift in Black hole spacetime

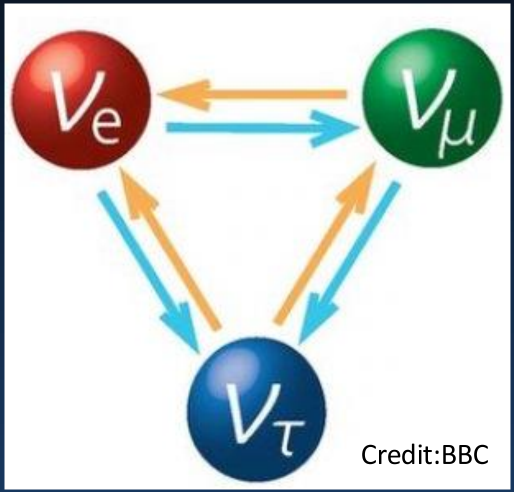


Akaho et al. 2020

- There remain many ”holes” in CCSN/BNSM simulations

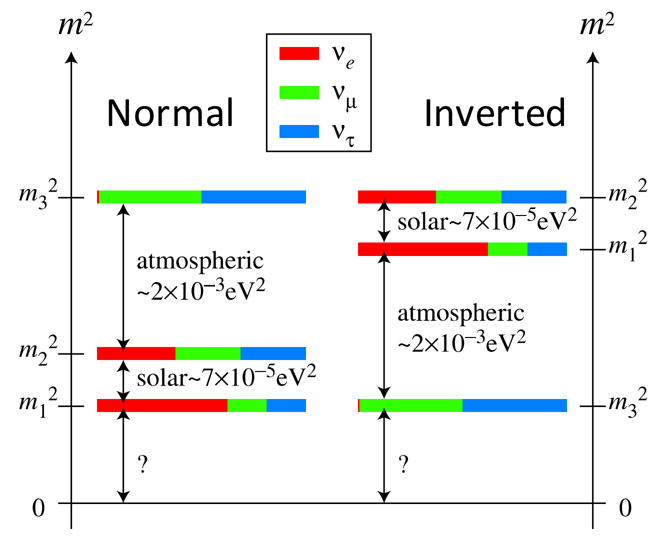


Neutrino oscillations



Credit:BBC

- There are many experimental evidences that neutrinos can go through flavor conversion.
- Neutrinos have at least three different masses.
- Flavor eigenstates are different from mass eigenstates.



Feruglio et al. 2003

$$|\nu_i\rangle = \sum_\alpha U^*_{\alpha i}\,|\nu_\alpha\rangle\,,$$

Mass state

$$|\nu_\alpha\rangle = \sum_i U_{\alpha i}\,|\nu_i\rangle\,,$$

Flavor state

U represents

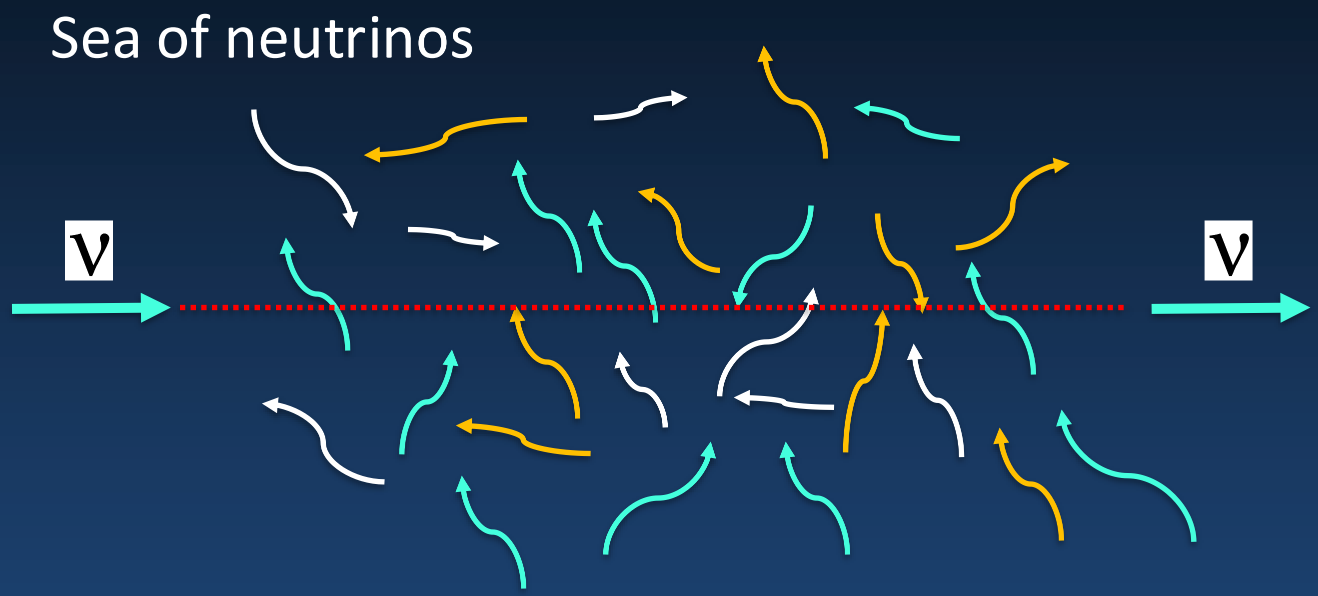
Pontecorvo–Maki–Nakagawa–Sakata matrix

(PMNS matrix)

$$\begin{aligned} U &= \begin{bmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu1} & U_{\mu2} & U_{\mu3} \\ U_{\tau1} & U_{\tau2} & U_{\tau3} \end{bmatrix} \\ &= \begin{bmatrix} 1 & 0 & 0 \\ 0 & c_{23} & s_{23} \\ 0 & -s_{23} & c_{23} \end{bmatrix} \begin{bmatrix} c_{13} & 0 & s_{13}e^{-i\delta} \\ 0 & 1 & 0 \\ -s_{13}e^{i\delta} & 0 & c_{13} \end{bmatrix} \begin{bmatrix} c_{12} & s_{12} & 0 \\ -s_{12} & c_{12} & 0 \\ 0 & 0 & 1 \end{bmatrix} \begin{bmatrix} e^{i\alpha_1/2} & 0 & 0 \\ 0 & e^{i\alpha_2/2} & 0 \\ 0 & 0 & 1 \end{bmatrix} \\ &= \begin{bmatrix} c_{12}c_{13} & s_{12}c_{13} & s_{13}e^{-i\delta} \\ -s_{12}c_{23} - c_{12}s_{23}s_{13}e^{i\delta} & c_{12}c_{23} - s_{12}s_{23}s_{13}e^{i\delta} & s_{23}c_{13} \\ s_{12}s_{23} - c_{12}c_{23}s_{13}e^{i\delta} & -c_{12}s_{23} - s_{12}c_{23}s_{13}e^{i\delta} & c_{23}c_{13} \end{bmatrix} \begin{bmatrix} e^{i\alpha_1/2} & 0 & 0 \\ 0 & e^{i\alpha_2/2} & 0 \\ 0 & 0 & 1 \end{bmatrix}, \end{aligned}$$

Neutrino oscillation induced by self-interactions

Pantalone 1992



1. Refractions by self-interactions induce neutrino flavor conversions, which is analogy to matter effects (e.g., MSW resonance).

2. The oscillation timescale is much shorter than the global scale of CCSN/BNSM.

3. Collective neutrino oscillation induced by neutrino-self interactions commonly occurs in CCSNe and BNSM environments.

Rich flavor-conversion phenomena driven by neutrino-neutrino self-interactions

Vacuum: $\omega = \frac{\Delta m^2}{2E_\nu},$

Matter: $\lambda = \sqrt{2}G_F n_e,$

Self-int: $\mu = \sqrt{2}G_F n_\nu,$

- Slow-mode (Duan et al. 2010)

- Energy-dependent flavor conversion occurs.
- The frequency of the flavor conversion is proportional to $\sqrt{\omega\mu}$

- Fast-mode (FFC) (Sawyer 2005)

- Collective neutrino oscillation in the limit of ω → 0.
- The frequency of the flavor conversion is proportional to μ
- Anisotropy of neutrino angular distributions drives the fast flavor-conversion.

- Collisional instability (Johns 2021)

- Asymmetries of matter interactions between neutrinos and anti-neutrinos drive flavor conversion.

$$\mathrm{Im}\,\boxtimes = \pm \frac{\Gamma - \bar{\Gamma}}{2} \sqrt{\frac{\mu S}{(\mu D)^2 + 4!\ \mu S}} - \frac{\Gamma + \bar{\Gamma}}{2}$$

Γ: Matter-interaction rate

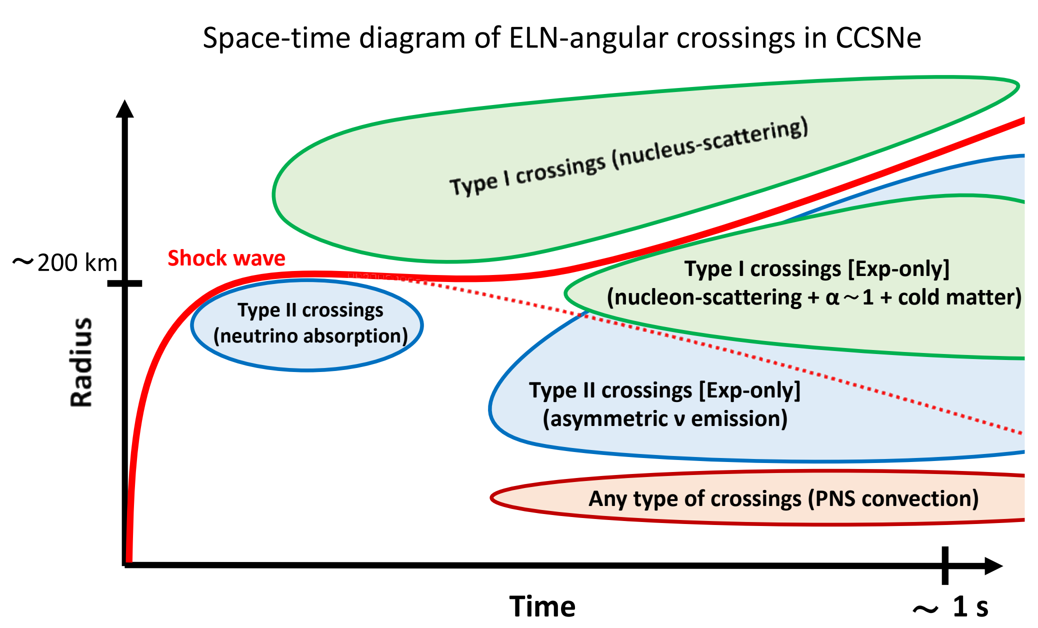
- Matter-neutrino resonance (Malkus et al. 2012)

- The resonance potentially occur in BNSM/Collapsar environment (but not in CCSN).
- Essentially the same mechanism as MSW resonance.

$$|\lambda + \mu| \sim |\omega|$$

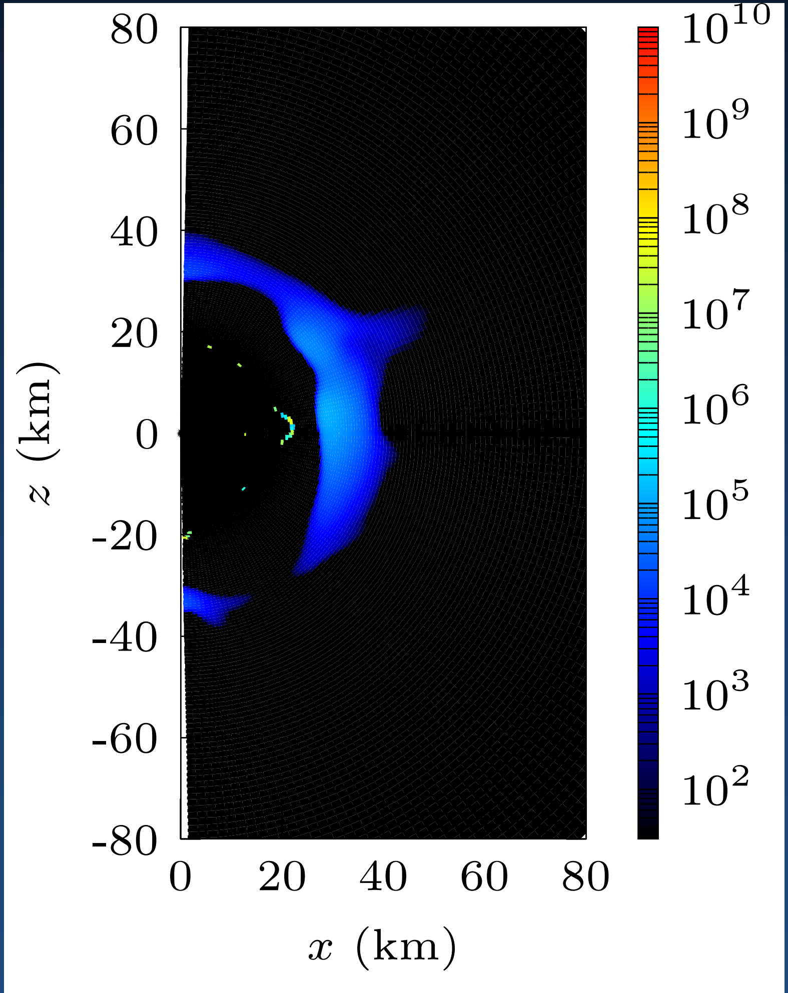
Neutrino flavor conversions are omnipresent in CCSN environments

Fast flavor conversions (FFC)



Nagakura et al. 2021

Collisional instability



Akaho et al. 2023

Quantum Kinetics neutrino transport:

Vlasenko et al. 2014, Volpe 2015,
Blaschke et al. 2016, Richers et al. 2019

$$\underline{p^\mu \frac{\partial \overset{(-)}{f}}{\partial x^\mu} + \frac{dp^i}{d\tau}\frac{\partial \overset{(-)}{f}}{\partial p^i}} = \underline{-p^\mu u_\mu \overset{(-)}{S}_{\text{col}}} + \underline{i p^\mu n_\mu [\overset{(-)}{H}, \overset{(-)}{f}]},$$

Advection terms
(Same as Boltz eq.)

Collision term

Oscillation term

f is not a
“distribution function”

Density matrix

$$\overset{(-)}{f} = \begin{bmatrix} \overset{(-)}{f}_{ee} & \overset{(-)}{f}_{e\mu} & \overset{(-)}{f}_{e\tau} \\ \overset{(-)}{f}_{\mu e} & \overset{(-)}{f}_{\mu\mu} & \overset{(-)}{f}_{\mu\tau} \\ \overset{(-)}{f}_{\tau e} & \overset{(-)}{f}_{\tau\mu} & \overset{(-)}{f}_{\tau\tau} \end{bmatrix}$$

Hamiltonian

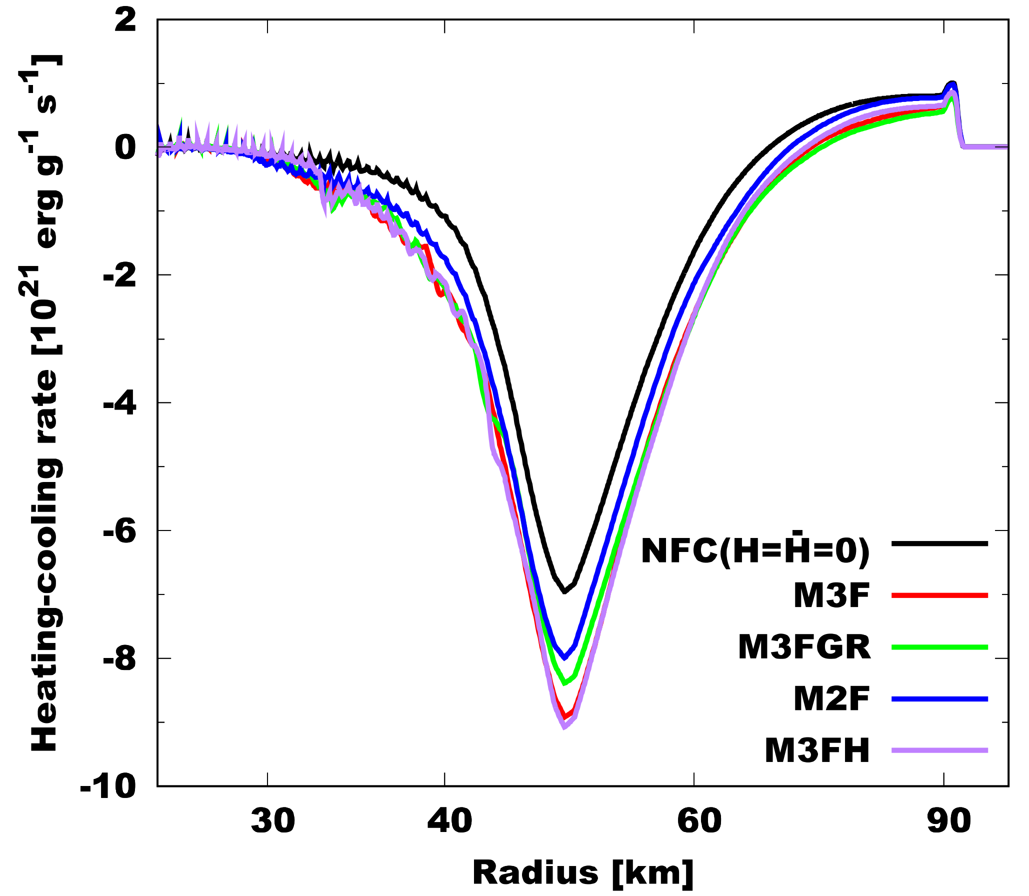
$$\overset{(-)}{H} = \overset{(-)}{H}_{\text{vac}} + \overset{(-)}{H}_{\text{mat}} + \overset{(-)}{H}_{\nu\nu},$$

$$H_{\text{vac}} = \frac{1}{2\nu} U \begin{bmatrix} m_1^2 & 0 & 0 \\ 0 & m_2^2 & 0 \\ 0 & 0 & m_3^2 \end{bmatrix} U^\dagger,$$

$$H_{\text{mat}} = D \begin{bmatrix} V_e & 0 & 0 \\ 0 & V_\mu & 0 \\ 0 & 0 & V_\tau + V_{\mu\tau} \end{bmatrix},$$

$$H_{\nu\nu} = \sqrt{2} G_F \int \frac{d^3 q'}{(2\pi)^3} \left(1 - \sum_{i=1}^{3} \ell'_{(i)} \ell_{(i)}\right)\left(f(q') - \bar{f}^*(q')\right),$$

- Global Simulations of FFC (in CCSN) Nagakura PRL 2023



Numerical setup:

Collision terms are switched on.

Fluid-profiles are taken from a CCSN simulation.

General relativistic effects are taken into account.

A wide spatial region is covered.

Three-flavor framework

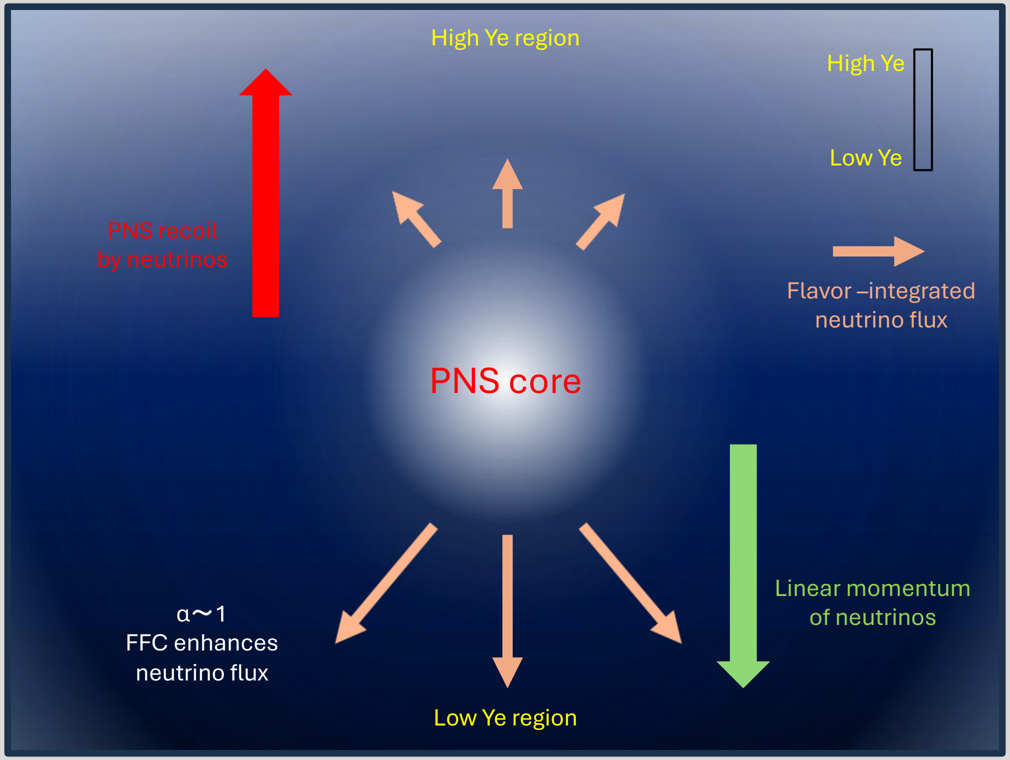
Neutrino-cooling is enhanced by FFCs
Neutrino-heating is suppressed by FFCs

➡

Impacts on explodabilities of CCSN

- Neutron star kick powered by neutrino flavor conversions

Nagakura and Sumiyoshi 2024



High Ye region

High Ye

Low Ye

PNS recoil by neutrinos

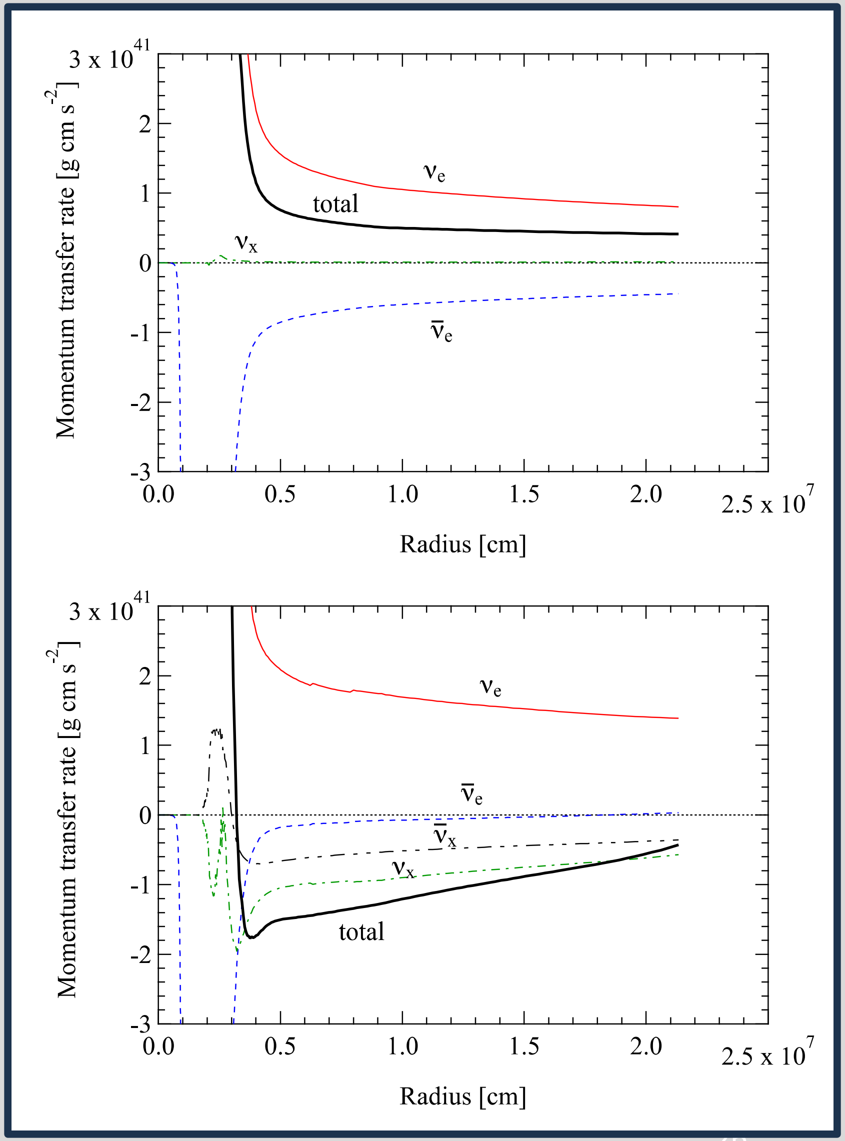
Flavor –integrated neutrino flux

PNS core

α∼1
FFC enhances neutrino flux

Linear momentum of neutrinos

Low Ye region



Summary:

- Remarkable progress on numerical modeling of CCSN have been made during the last decade.
- Observable signals can be discussed with realistic theoretical models.
- However, there are still many uncertainties in input physics.

  e.g., equation-of-state, weak interactions, and neutrino quantum kinetics
- These uncertainties should not be underestimated. They could be a game-changing ingredient in CCSN theory.
- Do not forget other mechanisms: MHD-driven, BH-driven, Phase-transition of NS etc.